

A unified model for AGN feedback in cosmological simulations of structure formation

Debra Sijacki,¹^{*} Volker Springel,¹ Tiziana Di Matteo² and Lars Hernquist³

¹Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Straße 1, 85740 Garching bei München, Germany

²Department of Physics, Carnegie-Mellon University, 5000 Forbes Ave., Pittsburgh, PA 15213, USA

³Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

Accepted 2007 June 26. Received 2007 June 21; in original form 2007 May 4

ABSTRACT

We discuss a numerical model for black hole growth and its associated feedback processes that for the first time allows cosmological simulations of structure formation to self-consistently follow the build up of the cosmic population of galaxies and active galactic nuclei (AGNs). Our model assumes that seed black holes are present at early cosmic epochs at the centres of forming haloes. We then track their growth from gas accretion and mergers with other black holes in the course of cosmic time. For black holes that are active, we distinguish between two distinct modes of feedback, depending on the black hole accretion rate itself. Black holes that accrete at high rates are assumed to be in a ‘quasar regime’, where we model their feedback by thermally coupling a small fraction of their bolometric luminosity to the surrounding gas. The quasar activity requires high densities of relatively cold gas around the black hole, as it is achieved through large-scale inflows triggered by galaxy mergers. For black holes with low accretion rates, we conjecture that most of their feedback occurs in mechanical form, where AGN-driven bubbles are injected into a gaseous environment. This regime of activity, which is subdominant in terms of total black hole mass growth, can be identified with radio galaxies in clusters of galaxies, and can suppress cluster cooling flows without the requirement of a triggering by mergers. Using our new model, we carry out *TREESPH* cosmological simulations on the scales of individual galaxies to those of massive galaxy clusters, both for isolated systems and for cosmological boxes. We demonstrate that our model produces results for the black hole and stellar mass densities in broad agreement with observational constraints. We find that the black holes significantly influence the evolution of their host galaxies, changing their star formation history, their amount of cold gas and their colours. Furthermore, the properties of intracluster gas are affected strongly by the presence of massive black holes in the cores of galaxy clusters, leading to shallower metallicity and entropy profiles, and to a suppression of strong cooling flows. Our results support the notion that AGNs are a key ingredient in cosmological structure formation. They lead to a self-regulated growth of black holes and bring the simulated properties of their host galaxies into much better agreement with observations.

Key words: black hole physics – methods: numerical – galaxies: clusters: general – galaxies: formation – cosmology: theory.

1 INTRODUCTION

It is now widely believed that most if not all galaxies with an spheroidal component harbour a supermassive black hole (BH) in their centres. Interestingly, the masses of these central BHs are found to be tightly linked with the stellar properties of their host galaxies, as expressed e.g. in the correlation of bulge velocity

dispersion with BH mass (Ferrarese & Merritt 2000; Gebhardt et al. 2000; Tremaine et al. 2002), or in the relation of bulge stellar mass (Kormendy & Richstone 1995; Magorrian et al. 1998; Marconi & Hunt 2003; Häring & Rix 2004) with BH mass. The existence of these relationships indicates that the formation and evolution of galaxies is fundamentally influenced by the presence of BHs and vice versa. We thus also expect that the environment and the cosmological evolution of galaxies will affect the way BHs grow.

In fact, there is a plethora of observational and theoretical studies that suggest that several different channels for interaction of BHs

*E-mail: deboras@mpa-garching.mpg.de

with their surroundings exist. At high redshift, mergers of gas-rich galaxies occur frequently and funnel copious amounts of cold gas towards the central regions of galaxies, such that the embedded BHs can reach high gas accretion rates. The radiation energy associated with the accretion can support the enormous luminosities of powerful quasars. Theoretically it has been hypothesized (Silk & Rees 1998; Fabian & Iwasawa 1999; King 2003) that quasars produce high velocity winds, which are expected to affect the properties of the host galaxy. The presence of quasar induced outflows has been observationally confirmed in a number of cases (e.g. Chartas, Brandt & Gallagher 2003; Crenshaw, Kraemer & George 2003; Pounds et al. 2003), and has first been demonstrated in simulations of merging galaxy pairs by Di Matteo, Springel & Hernquist (2005). Several numerical studies dealing with BH microphysics (e.g. Proga 2003; McKinney 2006) also predict existence of quasar outflows. Moreover, it appears that tidally disrupted galaxies are preferentially associated with active galactic nucleus (AGN) activity (for a review see Barnes & Hernquist 1992). Quasar activity hence appears to be directly linked to mergers of galaxies, and should represent the dominant mode of mass growth in the BH population.

Indeed, using semi-analytic models of galaxy formation, Kauffmann & Haehnelt (2000) have demonstrated that BH growth associated with mergers in cold dark matter (CDM) models can reproduce many properties of the observed quasar population as well as the inferred BH mass density today (see also Volonteri, Haardt & Madau 2003). Based on the detailed hydrodynamical simulations of BH growth in galaxy mergers (Di Matteo et al. 2005; Springel, Di Matteo & Hernquist 2005a; Cox et al. 2006a,b; Robertson et al. 2006a,b) that have recently become available, Hopkins et al. (2005, 2006a,c) have proposed a comprehensive picture of a unified, merger-driven model for the origin of quasars and their relation to spheroid formation, which also implies the existence of a ‘black hole fundamental plane’ (Hopkins et al. 2007b). Furthermore, rapid merging of the gas-rich progenitor systems of rare, massive galaxy clusters has been shown (Li et al. 2006) to be a viable formation path for supermassive BHs that are as massive at $z \sim 6$ as those seen in luminous high-redshift Sloan Digital Sky Survey (SDSS) quasars (Fan et al. 2001).

However, there also appears to exist another channel of BH interaction with host galaxies, which is neither related to powerful quasar activity nor associated with galaxy mergers. Evidence for this interaction can be seen in a number of local elliptical galaxies and central cluster galaxies, which contain X-ray cavities filled with relativistic plasma (Birzan et al. 2004; McNamara et al. 2005; Fabian et al. 2006; Forman et al. 2006) while harbouring seemingly ‘dormant’ BHs. These X-ray depressions, often referred to as ‘bubbles’, are thought to be inflated by relativistic jets launched from the central BH. Even though the radiative output from the central BH is not significant, the associated mechanical luminosity can be very important in these systems. There has been considerable effort from the theoretical point of view (e.g. Binney & Tabor 1995; Churazov et al. 2001, 2002; Quilis, Bower & Balogh 2001; Ruszkowski & Begelman 2002; Dalla Vecchia et al. 2004; Kawata & Gibson 2005; Sijacki & Springel 2006; Thacker, Scannapieco & Couchman 2006; Okamoto, Nemmen & Bower 2007) to understand the relevance of this mode of AGN feedback, which is widely considered to be leading candidate for resolving the cooling flow problem in groups and clusters of galaxies.

Given the complex physics of AGN–galaxy interactions, is it possible to construct a simple unified model that accounts for the different modes of BH feedback in a cosmological framework? First

attempts in this direction have already been made (Churazov et al. 2005; Croton et al. 2006; Merloni & Heinz 2006), motivated by the observational findings of X-ray binaries (Fender et al. 1999; Gallo, Fender & Pooley 2003). In particular, it has been shown that X-ray binaries switch between two states: in the so-called ‘low/hard’ state, a steady radio jet is present and the hard X-ray spectrum is observed, while in the ‘high/soft’ state, the jet vanishes and the X-ray spectrum shows a soft, thermal component. The transition between these two states is regulated by the accretion rate on to the BH itself, where the threshold value is of the order of 10^{-2} – $10^{-1} \dot{M}_{\text{Edd}}$. The ‘high/soft’ state can be explained by the standard, optically thick and geometrically thin accretion disc (Shakura & Sunyaev 1973), with BH accretion occurring at high rates and in a radiatively efficient mode. The ‘low/hard’ state on the other hand corresponds to optically thin, geometrically thick and radiatively inefficient accretion, as described by the theoretical Advection-Dominated Accretion Flow (ADAF) and Adiabatic Inflow-Outflow Solutions (ADIOS) models (Narayan & Yi 1994; Blandford & Begelman 1999).

The above suggests a rather simple, yet attractive scenario for distinguishing between different modes of BH feedback in models for the cosmological evolution of AGNs: we shall assume that for high accretion rates, a ‘quasar-like’ feedback occurs, while for states of low accretion, mechanical bubble feedback applies. It is clear that the simplicity of this model will not allow it to explain all kinds of AGN feedback phenomena, e.g. powerful radio galaxies that accrete at very high rates, as found in some protocluster environments, are not well represented in this simple scheme. Moreover, even though the physics of X-ray binaries is expected to be quite similar to the one of AGN (Heinz et al. 2005), an analogous transition between ‘high/soft’ and ‘low/hard’ states in AGN is observationally still not well established, although there are some encouraging observational findings (Maccarone, Gallo & Fender 2003; Körding, Jester & Fender 2006) in this direction.

With these caveats in mind, we explore in this study a new ‘two-mode’ AGN feedback scenario in fully self-consistent cosmological simulations of structure formation. This complements and extends our study of cosmological simulations with quasar feedback in Di Matteo et al. (2007), where we did not account for the ‘radio’ mode. Within our numerical model we represent BHs with collisionless ‘sink’ particles and we adopt subresolution methods to compute their gas accretion rate, estimated from a Bondi prescription. Furthermore, we allow BHs to grow via mergers with other BHs that happen to be in their vicinity and that have sufficiently low relative speeds. We assume that BH feedback is composed of two modes, as motivated above, and we track their growth and feedback with cosmic time. We consider a vast range of objects harbouring BHs, from galaxies with $\sim 10^8 M_{\odot}$ stellar mass to massive galaxy clusters with total mass $\sim 10^{15} M_{\odot}$. Even though our BH growth model is rather simple and crude due to the inevitable numerical limitations, it represents the first attempt to include AGN feedback effects in cosmological simulations of structure formation. As we shall see, this approach produces significant improvements in the properties of simulated galaxies.

The outline of this paper is as follows. In Section 2, we describe our numerical method for incorporating the BH growth and feedback in cosmological hydrodynamical simulations. In Section 3, we test our model in simulations of isolated galaxy clusters, and explore the numerical parameter space. We then discuss cosmological simulations of galaxy cluster formation and evolution, subject to AGN feedback, in Section 4, while in Section 5 we consider cosmological simulations of homogeneously sampled periodic boxes. Finally, we summarize and discuss our results in Section 7.

2 METHODOLOGY

We use a novel version of the parallel TREESPH code GADGET-2 (Springel, Yoshida & White 2001b; Springel 2005) in this study, which employs an entropy-conserving formulation of smoothed particles hydrodynamics (Springel & Hernquist 2002). Besides following gravitational and non-radiative hydrodynamical processes, the code includes a treatment of radiative cooling for a primordial mixture of hydrogen and helium, and heating by a spatially uniform, time-dependent ultraviolet (UV) background (as in Katz, Weinberg & Hernquist 1996). Star formation and associated supernovae feedback processes are calculated in terms of a subresolution multiphase model for the interstellar medium (ISM) (Springel & Hernquist 2003a). We use a simple prescription for metal enrichment, and optionally incorporate galactic winds powered by supernovae, as implemented by Springel & Hernquist (2003a). Furthermore, we follow the growth and feedback of BHs based on a new model that combines the prescriptions outlined in Springel, Di Matteo & Hernquist (2005b) and Sijacki & Springel (2006).

In the following, we briefly summarize the main features of the BH model relevant for this study (see Springel et al. 2005b, for further details), focusing on extensions that permit us to follow BH growth and feedback in a cosmological simulation of structure formation (see also Di Matteo et al. 2007). We then discuss a new BH feedback prescription at low accretion rates, based on a ‘bubble’ heating scenario for representing AGN heating in the radiatively inefficient regime of accretion.

2.1 Black hole formation and growth

The BHs in the code are represented by collisionless sink particles, which may accrete gas from their surroundings based on a prescribed estimate for the accretion rate. Two BH particles are also allowed to merge if they fall within their local smoothed particle hydrodynamics (SPH) smoothing lengths and if their relative velocities are smaller than the local sound speed.

In our cosmological simulations of structure formation, we assume that low-mass seed BHs are produced sufficiently frequently such that any halo above a certain threshold mass contains one such BH at its centre. Whether these seed BHs originate in exploding Population III stars, in the collapse of star clusters, or are of primordial origin is not important for our analysis, but there needs to be a process that produces initial seed BHs which can then grow to the masses of supermassive BHs by gas accretion in the course of our simulations. For definiteness, in most of our simulations we adopt a seed BH mass of $10^5 h^{-1} M_\odot$ and endow all haloes with a mass larger than $5 \times 10^{10} h^{-1} M_\odot$ with a seed if they do not contain any BH already. We identify haloes without BHs on the fly during a simulation by frequently calling a fast, parallel friends-of-friends (FOF) algorithm that is built into our simulation code.

State-of-the-art cosmological simulations of structure formation reach mass resolutions that are at best of order of our initial BH seed mass, while the resolution typically reached is still considerably coarser than that. This would mean that the initial growth of the BH mass could be significantly affected by numerical discreteness effects if the sink particle can only swallow full gas particles, as is the case in our scheme. In order to avoid that the BH growth and the accretion rate estimate is strongly affected by this numerical discreteness limitation, we treat the BH mass M_{BH} as an internal degree of freedom of the sink particle. In the beginning, M_{BH} may differ from the dynamical mass M_{dyn} of the sink particle itself. The variable M_{BH} is integrated smoothly in time based on the estimated

accretion rate on to the BH, while M_{dyn} increases in discrete steps when the sink particle swallows a neighbouring gas particle. The latter process is modelled stochastically such that M_{dyn} tracks M_{BH} in the mean, with small oscillations around it. With this prescription for the BH mass, we can follow the early growth of BHs accurately in a subresolution fashion even when their mass may be smaller than the mass resolution of the simulation, while at late times (or in the limit of very good mass resolution) the two mass variables coincide. Of course, in the event of a merger of two BH sink particles, both M_{BH} and M_{dyn} are added together.

Following Di Matteo et al. (2005), we estimate the accretion rate on to a BH particle according to the Bondi–Hoyle–Lyttleton formula (Hoyle & Lyttleton 1939; Bondi & Hoyle 1944; Bondi 1952):

$$\dot{M}_{\text{BH}} = \frac{4\pi\alpha G^2 M_{\text{BH}}^2 \rho}{(c_s^2 + v^2)^{3/2}}, \quad (1)$$

where α is a dimensionless parameter, ρ is the density, c_s the sound speed of the gas and v is the velocity of the BH relative to the gas. We account for the possibility that the BH accretion has an upper limit given by the Eddington rate:

$$\dot{M}_{\text{Edd}} = \frac{4\pi G M_{\text{BH}} m_p}{\epsilon_r \sigma_T c}, \quad (2)$$

where m_p is the proton mass, σ_T is the Thompson cross-section and ϵ_r is the radiative efficiency, that we assume to be 0.1, which is the mean value for the radiatively efficient Shakura & Sunyaev (1973) accretion on to a Schwarzschild BH. In some of our numerical models we specifically explore the imprints of the imposed Eddington limit on the BH properties.

2.2 Black hole feedback

In the model of Springel et al. (2005b), it is assumed that a fixed fraction of the BH bolometric luminosity couples thermally to the local gas, independent of the accretion rate and environment. In our model we extend this BH feedback prescription in order to obtain a physically refined model for AGN heating both at high and at low accretion rates. We are motivated by the growing theoretical and observational evidence (Fender et al. 1999; Gallo et al. 2003; Churazov et al. 2005; Heinz et al. 2005; Croton et al. 2006) that AGN feedback is composed of two modes, analogous to states of X-ray binaries. Specifically, at high redshifts and for high accretion rates we assume that the bulk of AGN heating is originating in the luminous quasar activity. In this regime, BHs accrete efficiently and power luminous quasars where only a very small fraction of their bolometric luminosity couples thermally to the gas. On the other hand, at lower redshifts and for BHs accreting at much lower rates than their Eddington limits, AGN heating proceeds via radiatively inefficient feedback in a mostly mechanical form.

To model the transition between these two accretion and feedback modes, we introduce a threshold $\chi_{\text{radio}} = \dot{M}_{\text{BH}}/\dot{M}_{\text{Edd}}$ for the BH accretion rate (BHAR) in Eddington units, above which ‘quasar heating’ is operating, and below which we deal with ‘radio’ mode feedback, which we model by injecting bubbles into the host galaxy/cluster. Typically, we adopt a value of 10^{-2} for χ_{radio} , and we impose no other criterion to distinguish between the two modes of feedback.

For BHAR that are higher than χ_{radio} , we parametrize the feedback as in Springel et al. (2005b), i.e. a small fraction of the bolometric luminosity is coupled thermally and isotropically to the surrounding gas particles, with an amount given by

$$\dot{E}_{\text{feed}} = \epsilon_r L_r = \epsilon_r \epsilon_r \dot{M}_{\text{BH}} c^2. \quad (3)$$

Here ϵ_f gives the efficiency of thermal coupling. The value of 5 per cent adopted here brings the simulated $M_{\text{BH}}-\sigma_*$ relation obtained for remnants of isolated galaxy mergers in agreement with observations, as shown by Di Matteo et al. (2005).

Below χ_{radio} we assume that the accretion periodically produces an AGN jet which inflates hot bubbles in the surrounding gas. We clearly lack the numerical resolution for self-consistent *ab initio* simulations of the detailed physics of BH accretion and the involved relativistic magnetohydrodynamics (MHD) that is responsible for the actual jet creation. However, we can nevertheless try to represent the relevant heating mechanism by directly injecting the energy contained in the AGN-inflated bubbles into the intracluster medium (ICM). To this end we need to link the bubble properties, like radius, duty cycle and energy content, directly with the BH physics.

Up to now there is no compelling theory that can satisfactorily explain AGN jet formation, bubble inflation by the jet and the duty cycle of jet activity. On the other hand, a growing body of observational evidence (e.g. Birzan et al. 2004; McNamara et al. 2005; Dunn & Fabian 2006; Fabian et al. 2006; Forman et al. 2006) shows that AGN-driven bubbles are present in many systems, at different redshifts and over a range of masses, which constrains the duty cycle to be of the order of 10^7-10^8 yr. However, it is still not clear whether and how AGN activity can be influenced by properties of the host galaxy like its mass, dynamical state or central gas cooling rate. In light of this uncertainty we propose a simple model for radio feedback where we assume that an AGN-driven bubble will be created if a BH has increased its mass by a certain fraction $\delta_{\text{BH}} \equiv \delta M_{\text{BH}}/M_{\text{BH}}$. Note that with this choice we do not constrain the possible duty cycle of the jet itself. We only conjecture that whenever the BH increases its mass by δM_{BH} , the thermodynamical state of the surrounding gas will be affected by the BH activity, and that this can be represented in the form of the bubbles.

Thus, we relate the energy content of a bubble to the BH properties as

$$E_{\text{bub}} = \epsilon_m \epsilon_r c^2 \delta M_{\text{BH}}, \quad (4)$$

where ϵ_m is the efficiency of mechanical heating by the bubbles. Moreover, we link the bubble radius both to δM_{BH} and to the density of the surrounding ICM in the following way:

$$R_{\text{bub}} = R_{\text{bub},0} \left(\frac{E_{\text{bub}}/E_{\text{bub},0}}{\rho_{\text{ICM}}/\rho_{\text{ICM},0}} \right)^{1/5}, \quad (5)$$

where $R_{\text{bub},0}$, $E_{\text{bub},0}$ and $\rho_{\text{ICM},0}$ are normalization constants for the bubble radius, energy content and ambient density, respectively. The scaling of the bubble radius is motivated by the solutions for the radio cocoon expansion in an spherically symmetric case (Scheuer 1974; Begelman & Cioffi 1989; Heinz, Reynolds & Begelman 1998). With this parametrization for R_{bub} , we mimic a scenario in which a more powerful jet will inflate bigger radio lobes, and where a higher ICM density will confine the size of the buoyant bubbles more. Finally, the spatial injection of the bubbles is random within a sphere with radius twice the bubble radius and centred on the BH particle in consideration.

3 SELF-REGULATED BUBBLE FEEDBACK IN ISOLATED HALO SIMULATIONS

We have first carried out a number of simulations of isolated galaxy clusters in order to test our BH feedback model and to explore its parameter space. These simulations consist of a static NFW dark matter halo (Navarro, Frenk & White 1996, 1997) with a gas component initially in hydrostatic equilibrium. The initial gas density

Table 1. Numerical parameters of our simulations of isolated galaxy clusters. The virial masses and radii of the haloes, evaluated at $200\rho_{\text{crit}}$, are given in the first two columns. The assumed values for the concentration parameter are in the third column, while the number and the mass of the gas particles are shown in the fourth and the fifth column, respectively. The mass of the star particles is half that of the gas particles, because we set the number of generations of star particles that a gas particle may produce to two. Note that there are no parameters for the dark matter particles in these run, because we here modelled the dark halo with a static NFW potential. Finally, in the last column, the gravitational softening length ϵ for the gas and star particles is given.

M_{200} ($h^{-1} M_{\odot}$)	R_{200} (h^{-1} kpc)	c	N_{gas}	m_{gas} ($h^{-1} M_{\odot}$)	ϵ (h^{-1} kpc)
10^{13}	444	8.0	3×10^5	4.0×10^6	2.0
10^{14}	957	6.5	3×10^5	4.0×10^7	5.0
10^{15}	2063	5.0	3×10^5	4.0×10^8	10.0
10^{15}	2063	5.0	1×10^6	1.2×10^8	6.5

profile has a similar form as the dark matter, but with a slightly softened core, as explained in more detail in Sijacki & Springel (2006). We construct initial conditions for haloes with a range of masses (see Table 1 for the details of the numerical set-up) and evolve them non-radiatively for a time $0.25 t_{\text{Hubble}}$ to damp out possible initial transients. Then, we ‘switch on’ radiative gas cooling and star formation, and also introduce a seed BH particle in the halo centre. In the following sections, we describe the results for the subsequent growth of the BHs.

3.1 Exploring the parameter space

In this section, we analyse the sensitivity of our model and of the resulting BH and host halo properties with respect to the parameter choices we have adopted. We begin by considering an isolated halo of mass $10^{14} h^{-1} M_{\odot}$, comparing simulations with and without BHs. For these numerical tests, we only consider AGN feedback in the form of bubbles, while in all full cosmological simulations we will include both modes of AGN feedback introduced in Section 2.2.

First, we vary the threshold for bubble triggering in terms of accreted mass, δ_{BH} , from 0.1 to 5 per cent, which will affect both the number of bubble events and their mean energy. The other two free parameters in the BH feedback model, ϵ_m and the normalization value for R_{bub} , are kept fixed in this first series of simulations in order to facilitate the comparison. For definiteness, for most of the isolated halo simulations we select $\epsilon_m = 1$, $R_{\text{bub},0} = 30 h^{-1}$ kpc, $E_{\text{bub},0} = 5 \times 10^{60}$ erg and $\rho_{\text{ICM},0} = 10^6 h^2 M_{\odot} \text{ kpc}^{-3}$. Note that our choice for these parameters will be slightly different in cosmological simulations, as we will discuss in Section 4. In most of our numerical experiments we start with a small BH seed, typically equal to $10^5 h^{-1} M_{\odot}$, to establish if the BH growth is self-regulated and whether it leads to a realistic BH mass when a quasi-stationary state is reached. In a number of test simulations we also explore a scenario in which a massive BH is already present at the very beginning.

In panels (a) and (b) of Fig. 1, we show the BH mass and the accretion rate expressed in Eddington units as a function of time, for three different values of $\delta_{\text{BH}} = 0.1, 1$ and 5 per cent. It can be seen that the BH is initially growing rapidly starting from the seed mass of $10^5 h^{-1} M_{\odot}$, and the accretion rate on to the BH is quite high, reaching the Eddington level at $t = 0.06 t_{\text{Hubble}}$. However, after this initial phase of rapid growth, AGN feedback starts

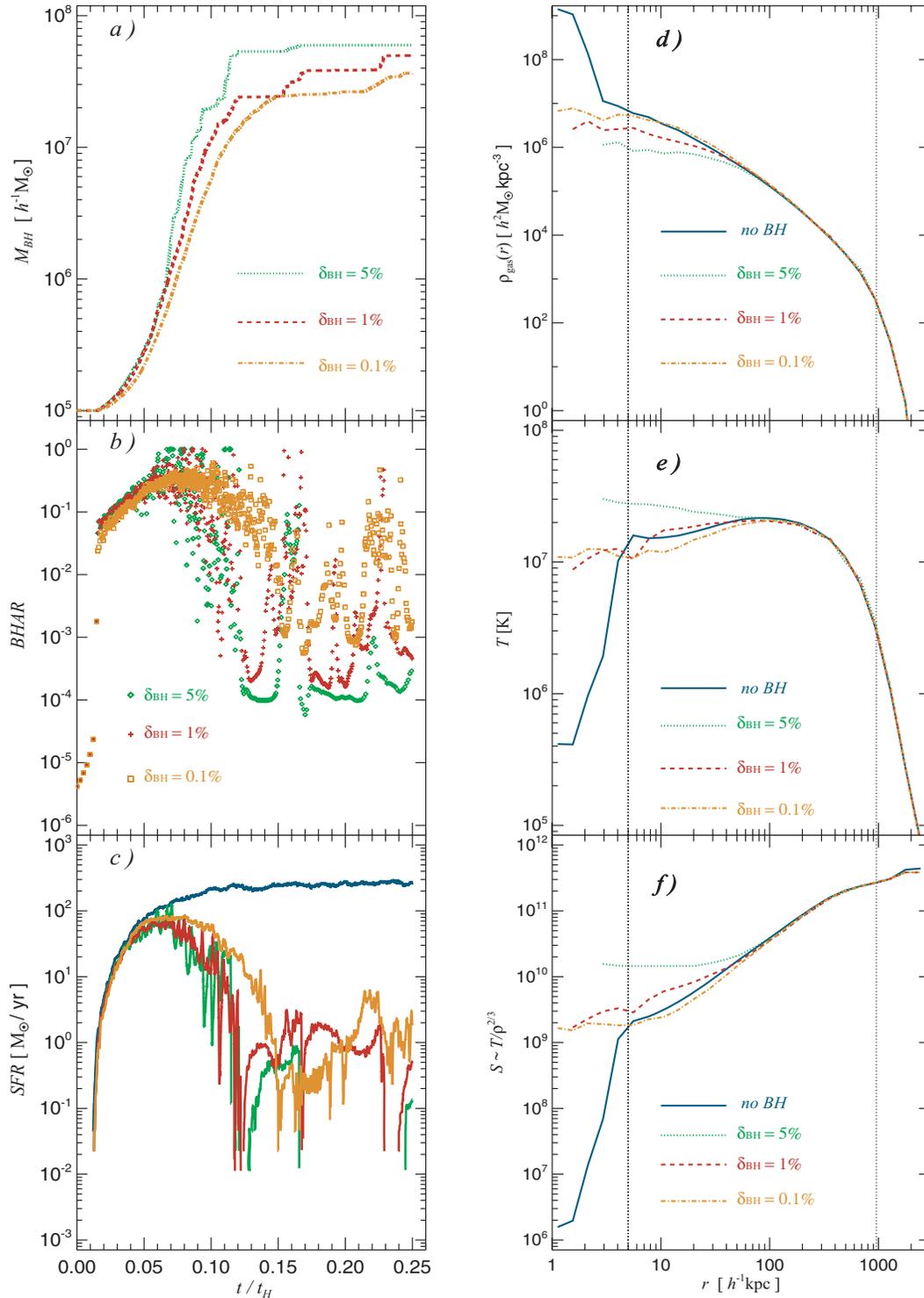


Figure 1. Panels (a) and (b) show the time evolution of the BH mass and the accretion rate in Eddington units for a BH that has been introduced in the centre of an isolated cluster of mass $10^{14} h^{-1} M_{\odot}$. The different curves and symbols give results for increasing δ_{BH} , as labelled in the panels. Panel (c) illustrates the SFR as a function of time for the same set of simulations (same colour coding as panel d). Finally, panels (d), (e) and (f) give the gas density, mass-weighted temperature and entropy radial profiles for this cluster at $t = 0.2 t_{\text{Hubble}}$ in runs without AGN feedback (blue continuous lines) and with AGN heating, where the different curves are for the increasing value of δ_{BH} , as indicated on the panels. The vertical dotted lines mark the gravitational softening length and the virial radius, respectively.

to reduce the further BH growth, because at this point enough heating is supplied to the surrounding medium, preventing it to quickly cool and sink towards the most inner regions. Thus, the feedback provided by AGN-driven bubbles reduces the supply of gas available for accretion by the central BH. Consequently, \dot{M}_{BH} drops and the mass of the BH shows no significant growth with time for a while. However, after a certain amount of time which depends on the heating efficiency, the central cluster gas starts to cool again, causing an increase of the BHAR and the triggering of another bubble episode. This mechanism establishes a self-regulated cycle, in which a balance between the gas cooling rate, the bubble heating rate and the residual BHAR is achieved. In panel (b) of Fig. 1, this cycle of AGN activity can be clearly seen. Here the peaks in BHAR correspond to the bubble injection events. Note that the higher values of δ_{BH} lead to fewer and more powerful AGN outbursts that cause the accretion rate to drop to very low values of order of $10^{-4} \dot{M}_{\text{Edd}}$.

The δ_{BH} parameter also affects the properties of the ICM. In panels (d), (e) and (f) of Fig. 1, we illustrate this effect by plotting the radial profiles of gas density, temperature and entropy, for the different values of δ_{BH} assumed in the simulations, as indicated in the panels. We also show the corresponding profiles for the run without AGN heating, denoted by the continuous blue lines. Clearly, there is a trend in the simulated profiles with δ_{BH} : more frequent and more gentle bubble heating events increase the central cluster entropy less than more powerful but rarer feedback episodes. A similar result has been found in the simulations of Omma & Binney (2004), where a higher jet power was heating the cluster for a longer time interval. However, regardless of the exact value of the δ_{BH} parameter, the central cooling catastrophe is prevented in our model cluster in all simulated cases with feedback. Once the feedback becomes effective, the BH mass stops growing when it reaches $3\text{--}6 \times 10^7 h^{-1} M_{\odot}$. Note, however, that the final value of the BH mass depends also on the ϵ_{m} parameter, which we have taken to be as high as possible in these test runs. Consequently, the BH masses obtained here represent a lower limit for the actual BH mass that one would expect in a cluster of this size.

In panel (c) of Fig. 1, we show how the star formation rate (SFR) is affected by AGN-driven bubbles. Without AGN heating, the SFR of the $10^{14} h^{-1} M_{\odot}$ isolated cluster grows with time to values as high as $\sim 250 M_{\odot} \text{ yr}^{-1}$. This high SFR results from the large amounts of cool gas that accumulate in the central regions, forming a reservoir for intense star formation activity. However, in the models with AGN feedback, the SFR drops dramatically to values of order of $\sim 1 M_{\odot} \text{ yr}^{-1}$. For these runs it can be also seen that the SFR shows a number of short spikes of somewhat enhanced activity, and these peaks are directly related with episodes of increased BHAR which lead to a triggering of bubbles and a successive reduction of the SFR.

We have also tested to which extent the AGN feedback is affected by changes in the normalization value for the bubble radius. As before, we adopt as default values for bubble energy and ICM density $5 \times 10^{60} \text{ erg}$ and $10^6 h^2 M_{\odot} \text{ kpc}^{-3}$, respectively, but we vary $R_{\text{bub},0}$ for different runs. In the range of 15 to $30 h^{-1} \text{ kpc}$ for $R_{\text{bub},0}$, our results are not altered significantly by modifications of this parameter. However, for substantially larger values of $R_{\text{bub},0}$, e.g. $60 h^{-1} \text{ kpc}$, the heating energy per particle drops to considerably smaller values, changing the initial rapid growth phase of the BH. For large values of $R_{\text{bub},0}$, the AGN feedback is less efficient in the early phase of growth such that the BH can grow to a somewhat larger mass before a self-regulated heating loop is established. We will further discuss the importance of the bubble radius when we consider haloes of different mass in Section 3.2.

Finally, we have numerically explored scenarios where an already massive BH is introduced at the very beginning of a simulation. For δ_{BH} equal to 0.1 per cent, the BH mass is found to increase from its initial value of 5×10^7 to $9 \times 10^7 h^{-1} M_{\odot}$ over a time-span equal to one quarter of the Hubble time. Most of this BH growth, however, can be attributed to the intermittent nature of bubble heating, while in the continuous AGN feedback regime, the BH mass increases by less than 1 per cent over the same simulated time-span. Thus, our numerical scheme indeed leads to a stable galaxy cluster solution where overcooling in the central parts is prevented and the BH mass grows to a self-regulated value.

3.2 AGN heating in haloes of different mass

In this section we examine AGN feedback effects in cluster haloes spanning a range in mass, from 10^{13} to $10^{15} h^{-1} M_{\odot}$. We adopt one set of feedback parameters for all simulations, namely ϵ_{m} , $R_{\text{bub},0}$, $E_{\text{bub},0}$ and $\rho_{\text{ICM},0}$, as in the previous section, and for δ_{BH} we choose the intermediate value of 1 per cent. We are interested in the question whether for a fixed set of parameters our model produces satisfactory solutions for haloes of very different masses.

Before we delve into a detailed quantitative analysis of the cluster simulations, we illustrate the visual morphology of simulated AGN-inflated bubbles in Fig. 2. For this purpose, we show projected mass-weighted temperature maps (first column), pressure fluctuation maps (middle column) and fluctuation maps of the X-ray luminosity (last column). Different rows are for clusters of increasing mass, from top to bottom. We call the maps in the second and the third columns ‘fluctuation’ maps because they have been constructed by dividing the original map by a version that has been smoothed on a scale of $16 h^{-1} \text{ kpc}$. This highlights local departures from the mean ICM properties, similar to what is known as unsharp masking technique in observational analysis.

A number of interesting features can be seen from the panels in Fig. 2. First, AGN-driven bubbles have characteristic mushroom-like and cap-like morphologies. These are particularly evident for the $10^{13} h^{-1} M_{\odot}$ halo, where several bubbles injected briefly after one another can be noticed. Note that the reason why there are so many bubbles in this particular map is in part due to the selected time for the image. At this moment \dot{M}_{BH} happens to grow rapidly, thus the bubble duty cycle is rather short. In the last row of Fig. 2, two bubbles can be seen in the central region of the $10^{15} h^{-1} M_{\odot}$ cluster, one south-west (SW) from the cluster centre and the other north-west (NW) and more distant. The latter one is uplifting some residual cool gas from the cluster centre, which forms a filamentary structure in the wake of the bubble. Looking at the pressure map of this cluster a nicely defined mushroom-like bubble can be seen that corresponds to the NW red blob in the temperature map. In general, the pressure and X-ray luminosity maps show numerous irregularities, sound waves and weak shocks that have been generated by the bubbles. The ripples in these maps become progressively weaker for more massive cluster, which is a consequence of the parametrization adopted in our model, as we will clarify in the following paragraphs.

In the uppermost panel of Fig. 3, we show the mass growth of the central BH as a function of time for the different cluster simulations. It can be seen that haloes of increasing mass exhibit a qualitatively similar behaviour for their BH growth: an initial rapid growth phase is followed by self-regulated stagnation. However, the starting point of the rapid growth is dictated by the gas cooling time, which is shortest for the smallest cluster mass considered. Therefore, for the $10^{13} h^{-1} M_{\odot}$ halo, the BH starts growing first, but it is also the first one to reach stagnation, while the BH sitting in the centre of the

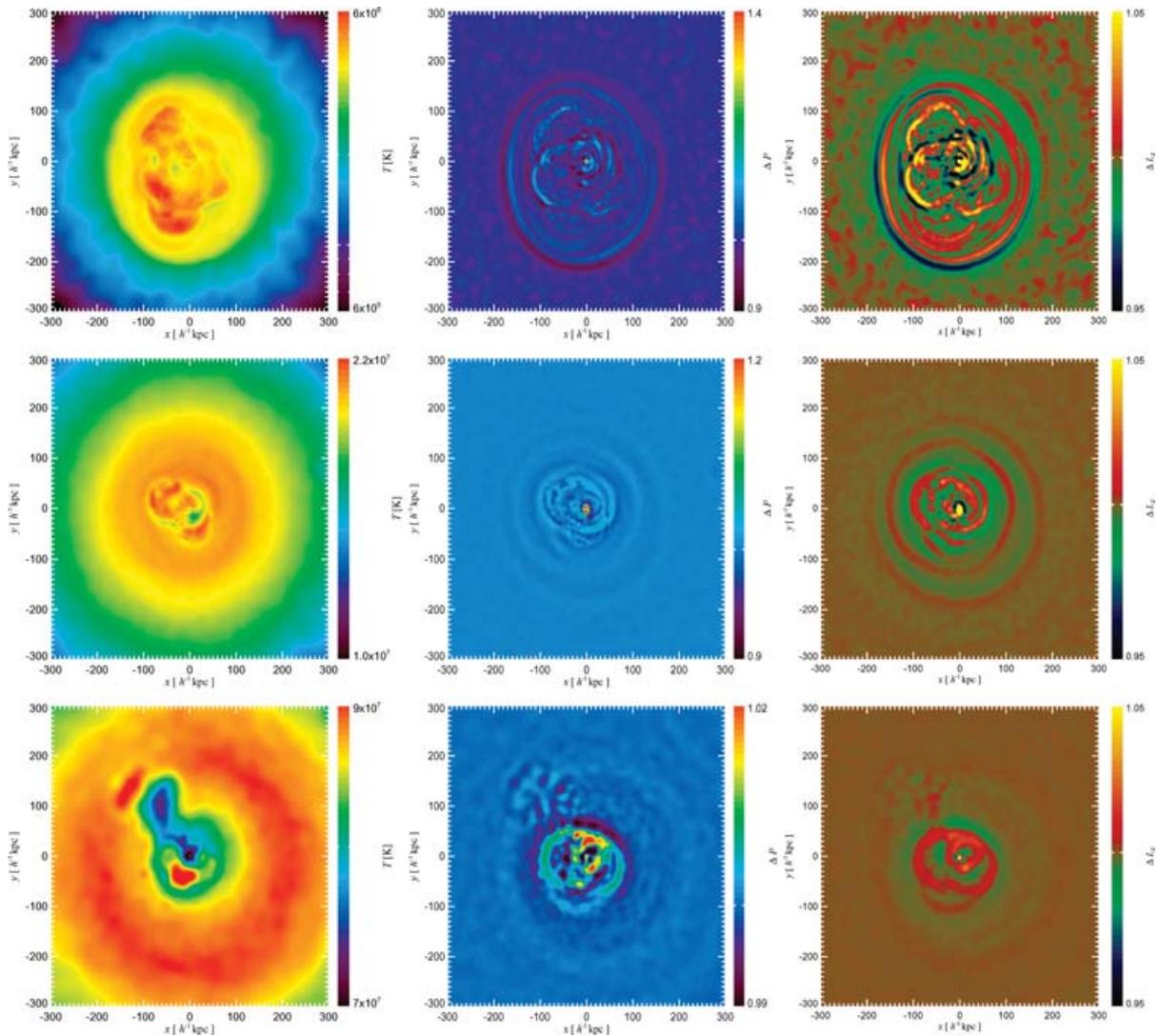


Figure 2. The panels on the left-hand side show mass-weighted projected temperature maps of three isolated clusters with increasing mass: $10^{13} h^{-1} M_{\odot}$ (uppermost panel), $10^{14} h^{-1} M_{\odot}$ (central panel) and $10^{15} h^{-1} M_{\odot}$ (lowermost panel). It can be seen that the bubbles injected in the central cluster regions have mushroom-like morphologies and are uplifting residual cool material from the centre (lowermost panel). The middle and right-hand panels give projected pressure and X-ray emission maps, divided by the corresponding map smoothed on a scale of $16 h^{-1} \text{kpc}$. A number of bubble-induced sound waves and weak shocks are clearly visible, and in the case of the $10^{15} h^{-1} M_{\odot}$ cluster, NW from the cluster centre a mushroom-shaped bubble in the pressure map is clearly visible which corresponds to the red blob in the temperature map.

$10^{15} h^{-1} M_{\odot}$ cluster waits for $\sim 0.05 t/t_{\text{Hubble}}$ before increasing its mass significantly. At the end of the simulated time-span, all three BHs reach equilibrium states and their final masses are $\sim 6.5 \times 10^6$, $\sim 5 \times 10^7$ and $\sim 1.8 \times 10^8 h^{-1} M_{\odot}$, respectively.

In the other two panels of Fig. 3, we show how certain bubble properties evolve with time for different cluster simulations. In the middle panel, $\Delta E_{\text{bub, thermal}}$ represents the thermal energy contrast of the bubble just before and after the energy was injected into it. $\Delta E_{\text{bub, thermal}}$ is reflecting mainly the mass growth of the BH, since it determines how much energy will be thermally coupled to the bubbles. It can be seen that $\Delta E_{\text{bub, thermal}}$ initially grows significantly with time, but when the BH growth saturates, it scatters around a constant value. Furthermore, there is a systematic trend of $\Delta E_{\text{bub, thermal}}$ with cluster mass, with smaller mass clusters showing a larger energy contrast for the bubbles. This is primarily due to the smaller radius of bubbles in less massive clusters, as can be seen in the bottom panel of Fig. 3. Given that we have adopted the same set of parameters

for our different cluster simulations, it follows from equation (5) that BHs of smaller mass will generate smaller bubbles. This explains why R_{bub} is growing in time and why it is bigger for more massive haloes that harbour larger BHs. This consequence of our model has also repercussions for the presence and intensity of sound waves generated by the inflation of bubbles, as seen in the maps of Fig. 2. Smaller bubbles with a higher energy contrast relative to the surrounding ICM are more efficient in generating weak shocks and sound waves than large bubbles with low $\Delta E_{\text{bub, thermal}}$.

In Fig. 4, we show radial profiles of gas density, temperature and entropy for our isolated clusters of mass 10^{13} and $10^{15} h^{-1} M_{\odot}$. The radial profiles of the $10^{14} h^{-1} M_{\odot}$ cluster can be found in Fig. 1 with the same colour coding. The blue, continuous lines denote the runs without AGN, while the red, dashed curves are for the simulations with AGN heating. The upper set of profiles in each panel is for the more massive cluster. The results show that for a fixed choice of parameters in our AGN heating model the ICM properties of

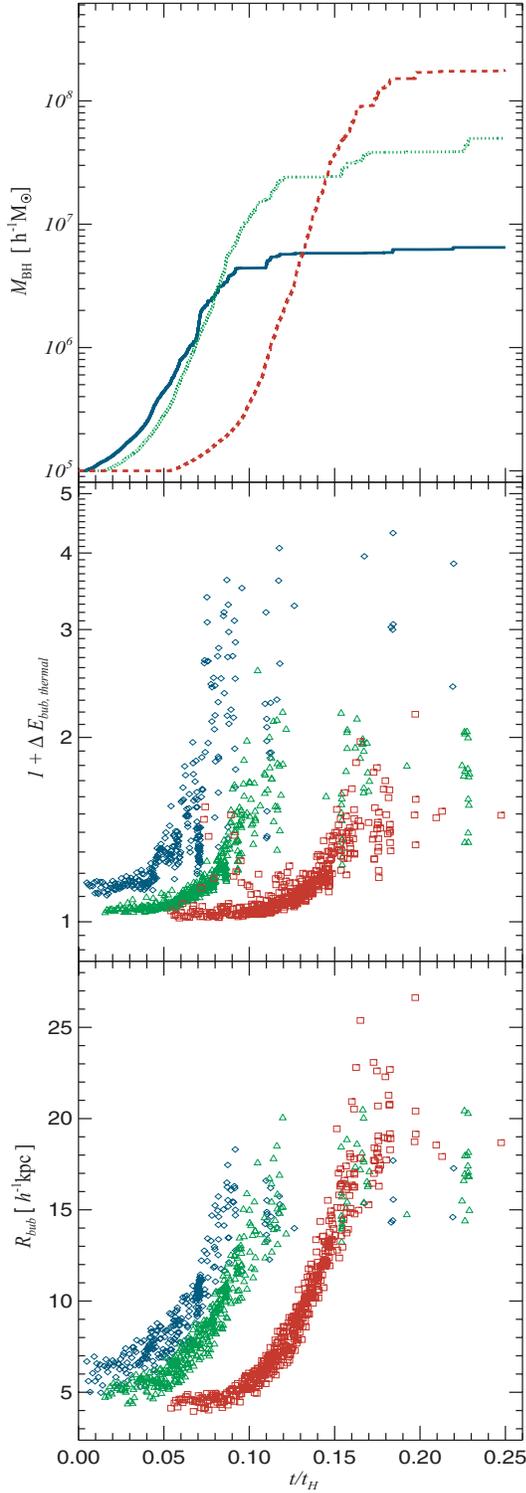


Figure 3. Growth of the BH mass with time (top panel) for three clusters of different mass. The middle panel shows how the thermal energy contrast, $\Delta E_{\text{bub, thermal}}$, depends on time and the BH mass in our cluster simulations. The bubble radius as a function of time is shown in the bottom panel, for the same set of runs. The blue, continuous line and the diamonds are for a cluster of mass $10^{13} h^{-1} M_{\odot}$, the green, dotted line and triangles are for a $10^{14} h^{-1} M_{\odot}$ halo, while the red, dashed line and the squares denote results for a $10^{15} h^{-1} M_{\odot}$ cluster.

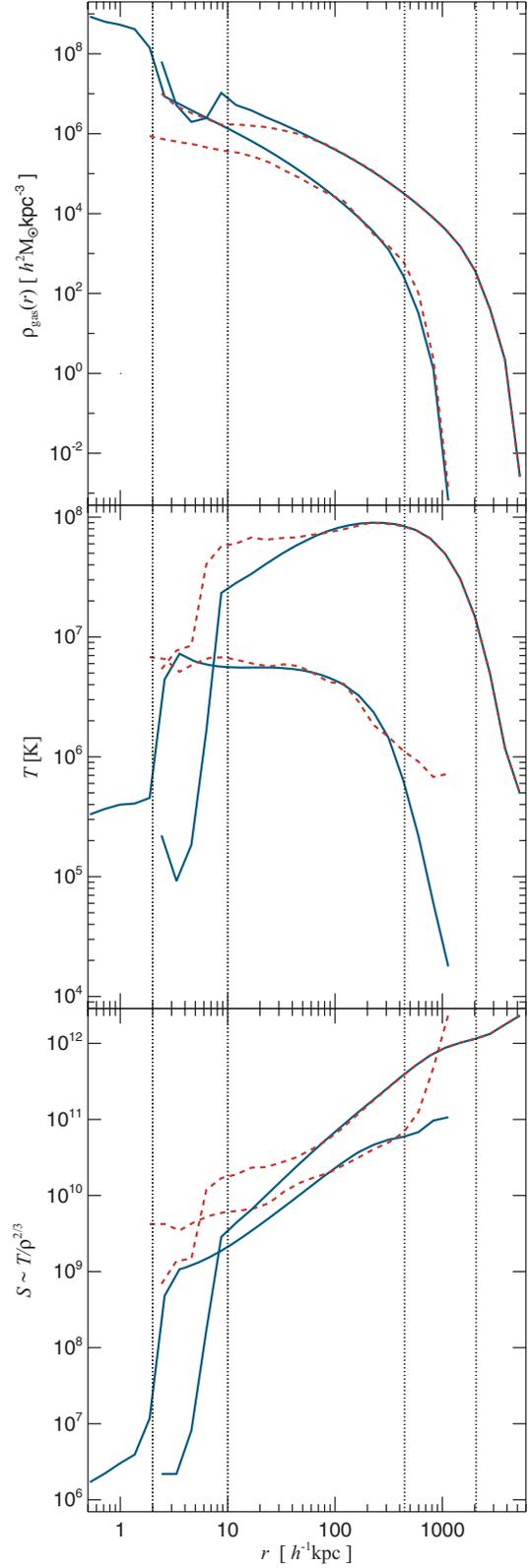


Figure 4. Radial profiles of gas density (uppermost panel), temperature (central panel) and entropy (lowermost panel) for 10^{13} and $10^{15} h^{-1} M_{\odot}$ galaxy clusters at $t = 0.2 t_{\text{Hubble}}$. The blue, continuous lines correspond to the simulations without AGN heating, while red, dashed lines give the case where AGN feedback is present. On every panel, the upper set of profiles is for the more massive cluster. Vertical dotted lines mark the gravitational softening length and the virial radii of these clusters.

Table 2. Numerical parameters of the cosmological galaxy cluster simulations analysed in this study. The values listed from the second to the fifth column refer to the number and to the mass of high-resolution dark matter particles and of gas particles. Note that the actual values of N_{gas} and m_{gas} vary in time due to star formation. The last three columns give the initial and final redshifts of the runs, and the gravitational softening length ϵ .

Simulation	N_{HR}	N_{gas}	$m_{\text{DM}} (h^{-1} M_{\odot})$	$m_{\text{gas}} (h^{-1} M_{\odot})$	z_{start}	z_{end}	$\epsilon (h^{-1} \text{kpc})$
g676	314 518	314 518	1.13×10^9	0.17×10^9	60	0	5.0
g1	4 937 886	4 937 886	1.13×10^9	0.17×10^9	60	0	5.0

clusters over a range of masses appear satisfactory: (i) the gas density is somewhat suppressed in central regions; (ii) very cool gas in the innermost regions of clusters is absent, and at the same time the AGN heating is not increasing the ICM temperature too much; (iii) the gas entropy is boosted in central regions, but without generating entropy inversions that are typically not observed in real systems. Together with realistic BH masses and BHAR, these promising features of the simulated ICM properties encourage us to proceed studying our model using fully cosmological simulations of galaxy cluster formation, which we consider in the next section.

4 COSMOLOGICAL SIMULATIONS OF AGN FEEDBACK IN GALAXY CLUSTERS

We now turn to self-consistent cosmological simulations of galaxy cluster formation, with the aim to understand how AGN feedback influences these objects at different cosmological epochs and how this influence depends on the mass and the evolutionary history of clusters. We focus our analysis on two galaxy cluster simulations that have quite different merging histories and different present-day masses. The clusters have been selected from a cosmological Λ CDM model with a box size of $479 h^{-1} \text{Mpc}$ (Jenkins et al. 2001; Yoshida, Sheth & Diaferio 2001), and were prepared by Dolag (2004) for resimulation at higher resolution using the Zoomed Initial Condition (ZIC) technique (Tormen, Bouchet & White 1997). The cosmological parameters of the simulations are those of a Λ CDM concordance cosmology with $\Omega_{\text{m}} = 0.3$, $\Omega_{\Lambda} = 0.7$, $\Omega_{\text{b}} = 0.04$, $\sigma_8 = 0.9$ and Hubble constant $H_0 = 70 \text{ km s}^{-1} \text{Mpc}^{-1}$ at the present epoch. In respect to the isolated haloes simulations we have adopted somewhat different parameters of the BH feedback model, with the normalization values for $E_{\text{bub},0} = 10^{55} \text{ erg}$, and $\rho_{\text{ICM},0} = 10^4 h^2 M_{\odot} \text{ kpc}^{-3}$ being lower, with $\delta_{\text{BH}} = 0.01$ per cent, and with a more realistic mechanical feedback efficiency of $\epsilon_{\text{m}} = 0.20$, while the other parameters were kept exactly the same.

The primary numerical parameters of our simulations are summarized in Table 2, and the main physical properties at $z = 0$ of the formed clusters are listed in Table 3. For both simulated galaxy clusters, the smaller one (‘g676’) and the more massive one (‘g1’), the

virial radius, virial mass, total gas mass and stellar mass have been measured in runs without AGN feedback (labelled ‘csf’) and when it is included (labelled ‘csfbh’). Additionally, the mass- and emission-weighted temperature, X-ray luminosity and total SFR are given in the sixth to the ninth column of Table 3, respectively. Looking at the gas properties listed in Table 3 it can be seen that AGN feedback does not significantly alter the global gas temperature, or the X-ray luminosity. However, there are still important local changes of these properties when AGN heating is included, as we will discuss in more detail in Section 4.3. On the other hand, the global stellar properties are strongly affected by AGN feedback: the total stellar mass is reduced considerably, as well as the total SFR, while the total gas mass is increased. This implies that the relative amount of gas versus stars is a function of AGN heating efficiency.

4.1 Black hole growth in clusters

In Fig. 5, we show a projected gas density map of the g676 galaxy cluster simulation at $z = 0$, with the positions of BHs more massive than $1.5 \times 10^7 h^{-1} M_{\odot}$ overlaid as black dots. The density map is $20 h^{-1} \text{Mpc}$ on a side and is centred on the most massive cluster galaxy. Interestingly, the surrounding large-scale structure shows many smaller haloes with embedded BHs, many of them residing in a filamentary region NW of the central object. In fact, most haloes are harbouring a central, massive BH. However, the most massive BH among them is sitting at the centre of the biggest halo, and this is also true at higher redshifts.

The mass as a function of time of this most massive BH is shown in Fig. 6. To produce this measurement, we have first constructed the full merger tree of all BHs formed in the simulated volume, allowing us to extract the main progenitor trunk corresponding to the most massive BH at $z = 0$, which is plotted as a thick continuous blue line. Furthermore, we show the mass growth of the secondary progenitors that at the moment of their last merger have a mass greater than $5 \times 10^7 h^{-1} M_{\odot}$. It can be seen that the most massive BH at $z = 0$ is not the first one to form, but it merges with two BHs that formed somewhat earlier and that are more massive for $z > 3.5$. This is similar to our findings in Di Matteo et al. (2007), where

Table 3. Physical properties of our sample of simulated galaxy clusters at $z = 0$, selected with a mean overdensity of $200\rho_{\text{crit}}$. For two different galaxy clusters, as labelled in the first column, the cluster radius, total mass, gas mass and stellar mass are given. Furthermore, the mass- and emission-weighted gas temperatures, X-ray luminosity and total SFR are listed, in the sixth to ninth column, respectively. Note that the values give results both for the simulations with cooling and star formation only (denoted with ‘csf’), and for runs performed with AGN heating (labelled ‘csfbh’).

Cluster	R_{200} ($h^{-1} \text{kpc}$)	M_{200} ($h^{-1} M_{\odot}$)	$M_{\text{gas},200}$ ($h^{-1} M_{\odot}$)	$M_{\text{stars},200}$ ($h^{-1} M_{\odot}$)	T_{mw} (K)	T_{ew} (K)	L_{X} (erg^{-1})	SFR ($\text{yr}^{-1} M_{\odot}$)
g676_csf	1176	1.13×10^{14}	9.3×10^{12}	4.7×10^{12}	1.4×10^7	2.6×10^7	1.6×10^{43}	51
g676_csfbh	1165	1.10×10^{14}	1.1×10^{13}	1.4×10^{12}	1.3×10^7	2.4×10^7	1.0×10^{43}	1
g1_csf	2857	1.63×10^{15}	1.4×10^{14}	6.3×10^{13}	7.3×10^7	1.3×10^8	1.0×10^{45}	742
g1_csfbh	2859	1.63×10^{15}	1.7×10^{14}	2.5×10^{13}	7.3×10^7	1.3×10^8	1.0×10^{45}	144

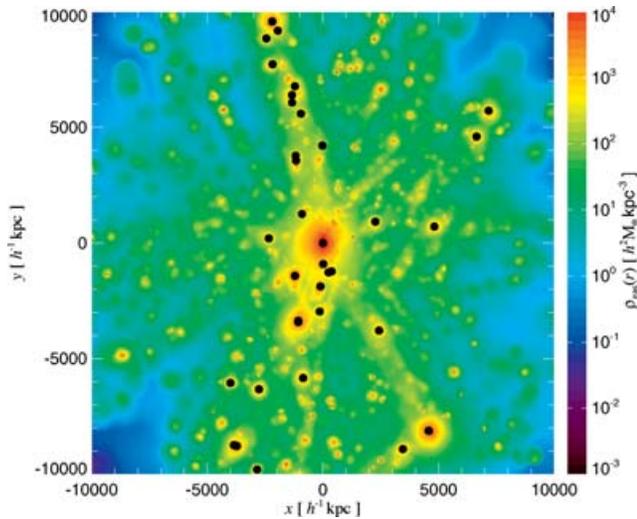


Figure 5. Projected gas density map of the g676 galaxy cluster simulation at $z = 0$, subject to AGN heating. Black dots mark the positions of BH particles that are more massive than $1.5 \times 10^7 h^{-1} M_{\odot}$. It can be seen that the BHs are residing not only in the centre of the most massive galaxy cluster but also in almost all smaller haloes that are visible on the map. Interestingly, there is a concentration of BHs in a filament that is extending from the central cluster towards NW, due to the ongoing structure formation there.

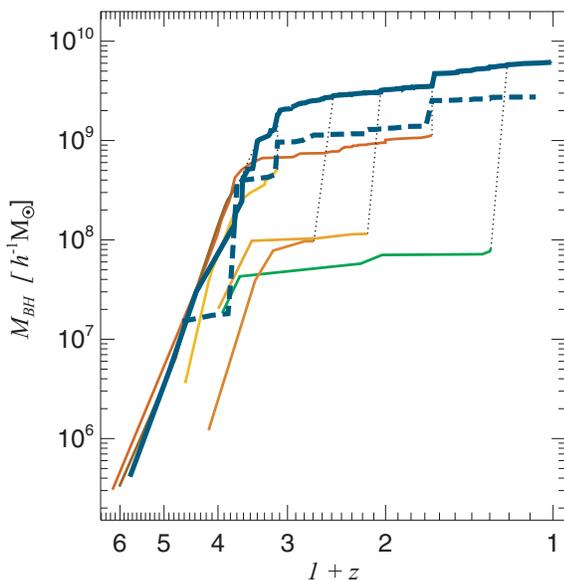


Figure 6. Merger tree of the most massive BH at $z = 0$ in the g676 galaxy cluster simulation. The thick blue line is showing the mass growth of the BH's first progenitor as a function of redshift, while the thin lines of different colour represent the evolutionary history of secondary progenitors, until the merger with the main progenitor occurs. For clarity, only secondary progenitors that at the moment of the last merger have a mass greater than $5 \times 10^7 h^{-1} M_{\odot}$ are shown. Vertical dotted lines indicate when merger events between the BHs in consideration occur. The dashed, thick blue line represents the cumulative mass of all secondary progenitors.

we showed that the most massive BHs found at high redshift in a cosmological volume need not necessarily be the most massive ones at low redshift. The thick, dashed blue line in Fig. 6 represents the cumulative mass of all secondary progenitors of the most massive BH at $z = 0$. This shows that mergers contribute up to 45 per cent to the final mass of the BH that defines the trunk, while the remaining

55 per cent are due to gas accretion by the BH on the trunk itself. Note, however, that ultimately all the BH mass is built-up by gas accretion (the mass of the initial seed BHs is negligible), but the part we here attributed to mergers was built-up by gas accretion in other systems.

We also note that the mass growth via mergers is more important at higher redshifts, with the last significant merging event occurring at $z \sim 0.7$. The merger events of BHs are mainly driven by the merging history of the host galaxy clusters. Indeed, there is a quite close correspondence between the evolution of the main progenitor of the most massive BH and the main progenitor of the biggest halo in the simulation. In particular, the main progenitor of the g676 galaxy cluster undergoes a number of mergers between redshifts $3.5 < z < 2.8$, when also the BH grows significantly via mergers with other BHs. Then at $z \sim 0.75$, the g676 cluster experiences its last major merger, visible as a jump in the main trunk in Fig. 6. At lower redshifts, the cluster becomes fairly relaxed and isolated, such that its central BH grows mostly by accreting gas that is cooling off from the hot cluster atmosphere. As we will discuss in Section 6 this mode of accretion (at low accretion rates) does however not contribute significantly to the overall BH mass growth.

Now we consider in more detail the mass growth and the accretion rate, not only of the most massive BH but also of all the BHs belonging to the simulated volume. Because of the better statistics provided by the g1 galaxy cluster simulation, we focus on this simulation, noting that the results for the g676 cluster are very similar. In Fig. 7, we plot the BHAR, expressed in Eddington units, as a function of the BH mass. The colour coding denotes the redshift, over the redshift interval $z = 4.3$ to 0 considered here. A number of interesting features can be noticed in this plot: (i) at a given BH mass, the accretion rate is decreasing with redshift; (ii) for intermediate redshifts, $2 < z < 3$, the BHAR is lowest for intermediate-mass BHs; (iii) at still lower redshifts, the BHAR for intermediate masses

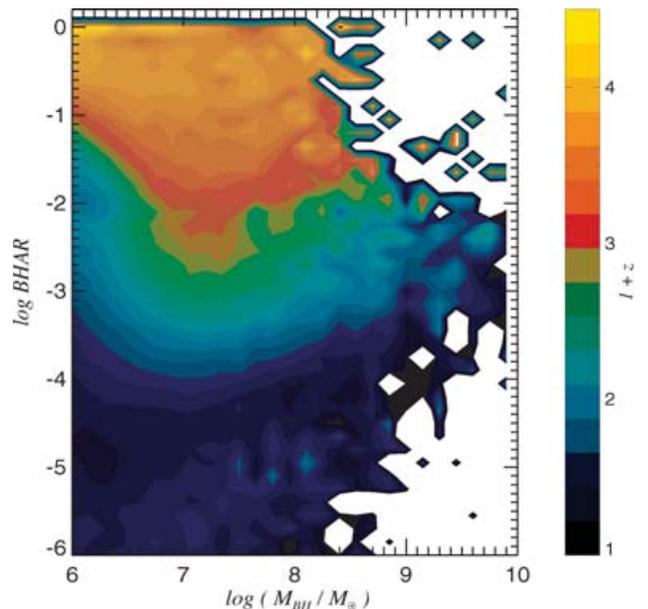


Figure 7. BHAR in Eddington units as a function of BH mass, for all BHs belonging to the g1 galaxy cluster simulation. The colour coding expresses the redshift, and the redshift interval considered here ranges from 4.3 to 0.0. For a given BH mass, the accretion rate is highest at early cosmic times. At lower redshifts, there is an upturn in the BHAR at the massive BH end, corresponding to BHs that are fed by gas that cools off from the atmosphere of the cluster and its massive progenitor systems.

keeps falling strongly, while it stays comparatively high for low-mass BHs as well as for very massive BHs; (iv) at the present epoch there is a large population of low-mass BHs that is still accreting efficiently. These features are qualitatively consistent with observational findings both from optical and X-rays surveys (e.g. Steffen et al. 2003; Ueda et al. 2003; Heckman et al. 2004; Barger et al. 2005; Hasinger, Miyaji & Schmidt 2005) that indicate that BHs are growing in a so-called ‘antihierarchical’ manner. In Section 5, we will come back to this issue, and discuss how the BHAR depends on the BH mass in a galactic environment.

At low redshifts, the increase of the BHAR with mass at the high BH mass end is driven by BHs that sit at the centres of massive galaxy clusters. These BHs are fed by gas from the hot cluster atmosphere, which would develop a strong cooling flow without the periodic heating by the AGN feedback. Interestingly, we also find that our BHAR of massive BHs at $z = 0$ agrees very well with a recent estimate of the Bondi accretion rate for a sample of X-ray luminous elliptical galaxies by Allen et al. (2006).

However, we need to emphasize that the simulations performed in this study were not designed to address the problem of BH formation and growth in the very early Universe, for which they lack the required resolution and volume. This becomes also apparent in the plot of Fig. 7, where the upper right corner is not populated. This means that these simulations cannot account for the existence of supermassive BHs of mass $\sim 10^9 M_{\odot}$ already at $z = 6$, as inferred from the luminosity of high redshift quasars (Fan et al. 2001). Using high-resolution compound galaxy models to populate a merger tree measured from a zoom simulation of a protocluster region, Li et al. (2006) have however recently demonstrated that at least in principle such early supermassive BHs can grow from accretion in gas-rich mergers at high redshift.

4.2 Heating the cluster outskirts

In Fig. 8, we examine the radial dependence of AGN heating in our simulated clusters for different redshift bins. To this end, we evaluate the AGN luminosity at the given epoch and plot it as a function of distance from the centre of the most massive cluster in the simulation. Depending on the BHAR, we decide whether the BH is in the ‘quasar’ or ‘radio’ mode, which we denote with star symbols or circles, respectively (see upper panel). In the lower panel, we plot the total AGN luminosity outside of a given radius, regardless in which mode the BHs accrete. It can be seen that the AGN heating at all redshifts considered and regardless of the feedback mode is most important in the cluster centre. However, at early times, and in particular for $1.5 < z < 4.3$, BHs that happen to reside in cluster outskirts during this time could provide an important additional source of ICM heating. These BHs are essentially all in the quasar phase where they accrete gas efficiently. Most of these BHs reside in satellite haloes that are entering the most massive system for the first time, and thus possibly are pre-heating the gas of smaller haloes prior to and during the merger with the massive cluster. These findings and the spatial distribution of AGN in galaxy clusters, illustrated in Fig. 5, are in a good agreement with observational evidence (e.g. Cappi et al. 2001; Cappelluti et al. 2005; Ruderman & Ebeling 2005; Martini et al. 2006) for the presence of AGN in galaxy cluster environments. At low redshifts, though, the AGN sitting at the cluster centre appears to be the dominant source of heating, as we confirm in Section 4.3. Interestingly, the feedback luminosity of the central AGN, regardless of its accretion regime, peaks at the similar value for the whole time-span considered.

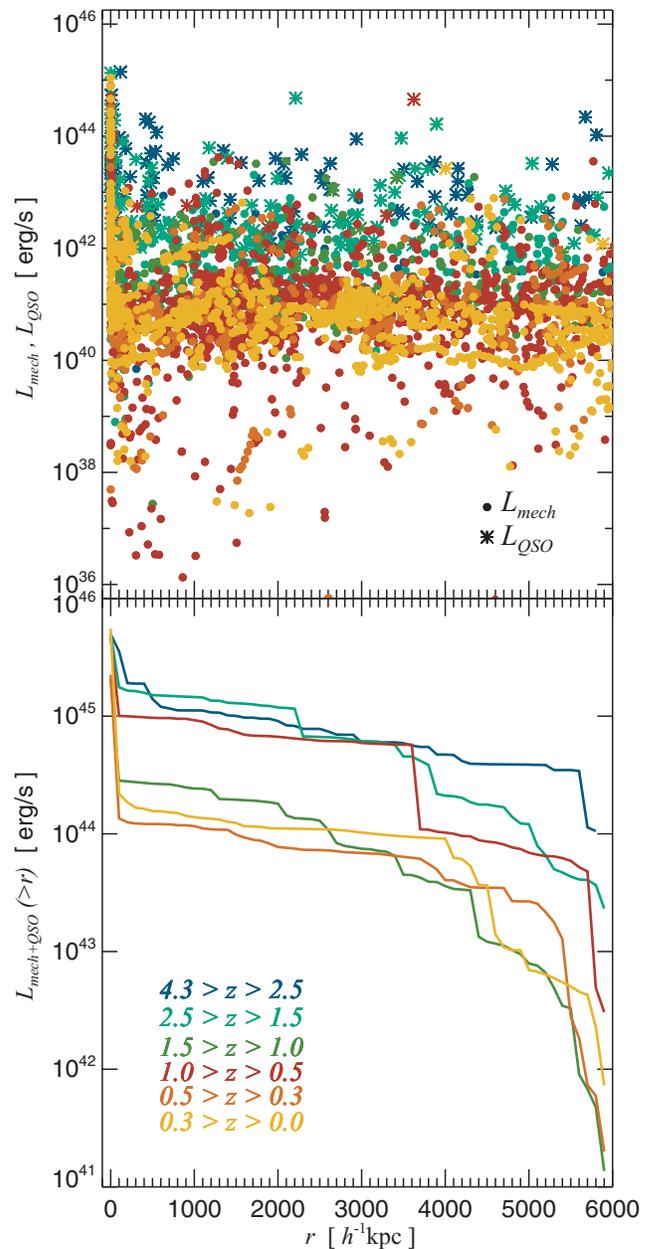


Figure 8. In the upper panel, the mechanical and quasar AGN luminosities are plotted as a function of distance from the most massive halo in the g676 galaxy cluster simulation. Depending on the BHAR, a BH is assumed to be either in a ‘quasar’ or in a ‘radio’ mode, and the BH luminosity is calculated accordingly and denoted with different symbols. The colour coding is according to redshift bins, as indicated in the lower panel. The total AGN luminosity outside a given radius, regardless of the feedback mode, is shown in the lower panel. It can be seen that the highest AGN luminosities always correspond to BHs sitting in the centre of the most massive clusters, and that at higher redshifts (especially $1.5 < z < 4.3$) heating from the quasars in cluster outskirts can be important.

4.3 The impact of AGN heating on the ICM

We here analyse how AGN feedback affects the ICM and the stellar properties of the cluster galaxies. In Fig. 9, we show radial profiles of gas density, temperature and entropy of the g676 cluster at three different epochs, both without (continuous blue lines) and with AGN heating (dashed red lines). Additionally, we have explored a model

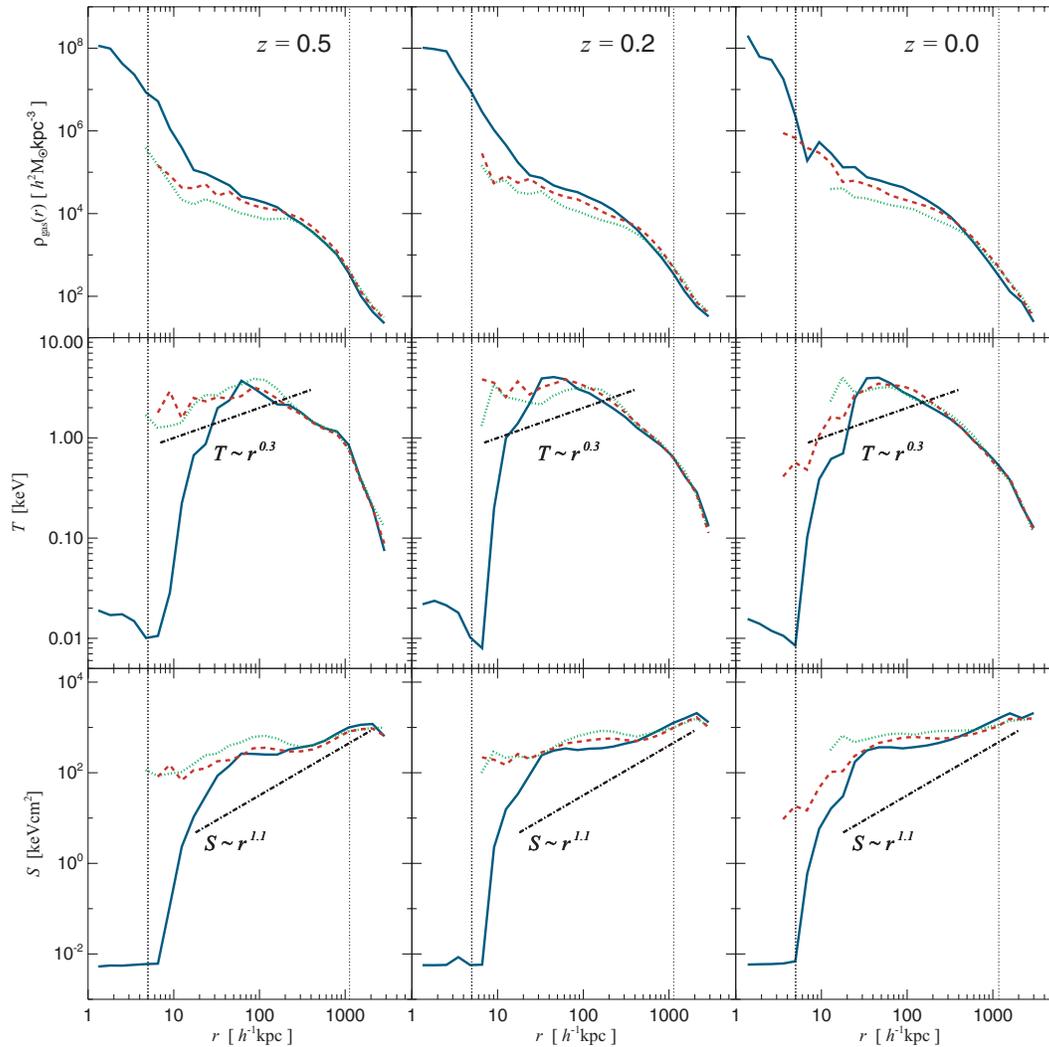


Figure 9. Radial profiles of gas density (upper row), mass-weighted temperature (central row) and entropy (lower row) of the g676 galaxy cluster at $z = 0.5$, 0.2 and 0.0 , respectively. Continuous blue lines illustrate the case without AGN heating, while dashed red lines are for simulations where AGN feedback is included. The run where the Eddington limit was not imposed on the BHAR is shown with a dotted green line. The vertical dotted lines denote the gravitational softening and the virial radius of this cluster. The dash-dotted lines in the central row show the slope of the central temperature profiles of the cool core clusters found by Sanderson et al. (2006), while the dash-dotted lines in the lower row illustrate the entropy scaling with radius, i.e. $S \propto r^{1.1}$.

in which the BHAR was not limited to the Eddington rate, which is drawn with a dotted green line. It can be seen that at all redshifts considered, the central gas density is reduced significantly as a result of AGN feedback. We also observe that the gas temperature in the central regions is fluctuating around the ‘isothermal’ value, sometimes for short time intervals increasing towards the centre, but most of the time being roughly constant or decreasing. These trends in gas temperature are due to the periodic nature of the bubble feedback. If a very energetic bubble is injected in the innermost cluster regions, the gas temperature increases for a short time, and reduces the BHAR considerably. After some time has elapsed, the gas will begin to cool again such that the central gas temperature will gradually start to decrease towards the centre. When enough gas has cooled off from the hot cluster atmosphere and has become available for accretion on to the supermassive BH, a new bubble will be triggered, establishing a self-regulated loop for the growth of the BH and the heating of the ICM. In some sense, this feedback loop acts as a thermometer for the ICM, preventing it to develop strong cooling flows.

On top of our simulated gas temperature profiles, we plot the slopes of the central temperature profiles of the cool core clusters recently found by Sanderson, Ponman & O’Sullivan (2006), showing a very encouraging agreement. Indeed, a number of observational studies (e.g. De Grandi & Molendi 2002; Vikhlinin et al. 2005; Dunn & Fabian 2006; Pratt et al. 2007) have shown that the central temperature profiles of relaxed, cool core clusters decrease with decreasing radius, while at the same time these systems require an additional source of heating in order to avoid excessive cooling flows. Theoretically, it is not trivial to explain this observed feature of the central cluster temperature – a self-regulated feedback mechanism is apparently needed which is sensitive to the local properties of the intracluster gas. Our model contains such a mechanism and represents to our knowledge the first successful attempt to simulate such a process in a cosmological environment.

As a result of the non-gravitational bubble heating, the ICM entropy is increased in the central regions, but maintains a monotonically increasing radial profile. Given the temperature of this cluster, the overall shape and the normalization of the simulated

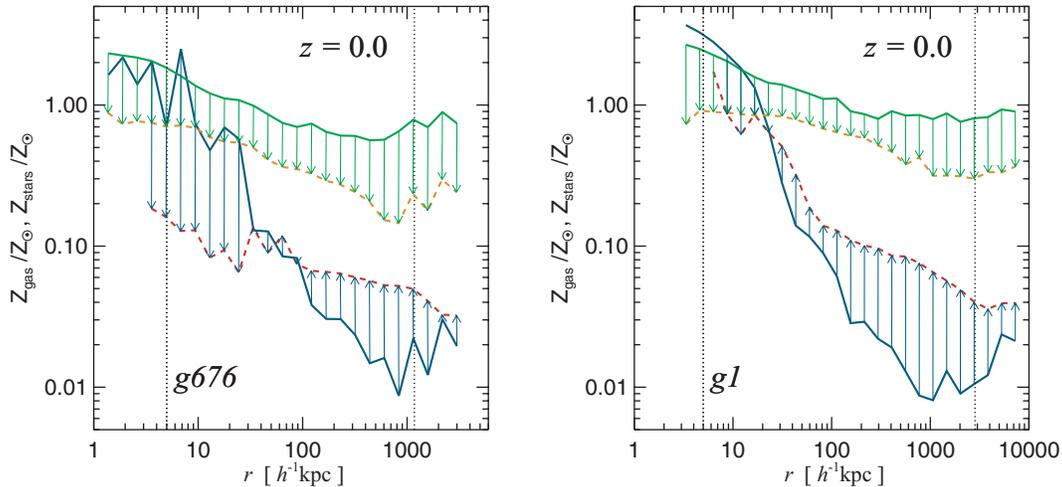


Figure 10. Radial profiles of mass-weighted gas metallicity and stellar metallicity of the g676 (left-hand panel) and g1 (right-hand panel) galaxy cluster simulations at $z = 0$. Continuous lines show the metallicity profiles without AGN heating, while dashed lines are for the runs with AGN feedback. The arrows illustrate how the ICM metallicity is affected by the BHs. The stellar metallicity is reduced at all radii, while the gas metallicity shows a tilt in the profile, decreasing in the inner regions and increasing for $r > 30 h^{-1}$ kpc, while still exhibiting a residual gradient.

entropy profiles are consistent with observational findings (e.g. Ponman, Sanderson & Finoguenov 2003; Pratt & Arnaud 2003). Observationally, it is found that the entropy scales with radius roughly as $S \propto r^{1.1}$ in the cluster outskirts, while departures from this scaling are seen in the central regions, out to $r \sim 0.1\text{--}0.2 R_{200}$. As can be noticed from the bottom row of Fig. 9, our simulated entropy profiles depart from this scaling out to somewhat larger radii, and this is also found for the ‘g1’ cluster simulation which has quite similar radial profiles. However, to decide whether this constitutes a real discrepancy between the simulations and observations, a larger set of simulated galaxy clusters is required, which we plan to compile in forthcoming work.

Looking at the radial gas profiles for $r > 400 h^{-1}$ kpc it is clear that the AGN heating does not significantly change the ICM properties at larger radii. Thus, heating of the cluster outskirts does not appear to be relevant for $z < 0.5$, confirming our findings in Section 4.2. Moreover, the radial profiles of our simulated galaxy cluster at $z = 0.5$ do not differ significantly from the ones at $z = 0$. This is consistent with observational findings (Bauer et al. 2005), where already

formed cool core clusters are seen at $z \sim 0.15\text{--}0.4$, with similar properties to the local population.

In Fig. 10, we show radial profiles of gas and stellar metallicity, in solar units, for the g676 and g1 cluster simulations at $z = 0$. The continuous lines show the metallicity profiles without AGN feedback, while the dashed lines illustrate how these profiles change when our BH model is ‘switched-on’. The stellar metallicity is reduced at all radii, while the gas metallicity shows a tilt: it is reduced in the innermost regions, for $r < 30 h^{-1}$ kpc, while it is increased for larger radii. These changes in the metallicity gradients are caused by two effects: first, due to the AGN feedback the total number of stars is reduced, not only in the central regions but also all over the cluster, as illustrated in Fig. 11; second, some of the metals accumulated in dense, star-forming regions are expelled into the hot ICM. This mechanism drives the tilt in the gas metallicity, increasing the metal content of the hot intracluster gas.

The effect of AGN heating on the metal production and mixing in clusters brings the simulated gradients into a much better agreement with observational results. However, the simulated gas

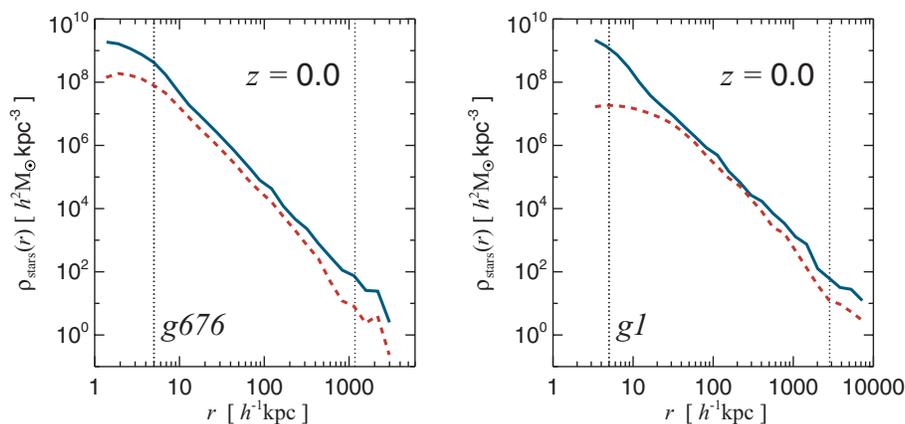


Figure 11. Stellar density profile of the g676 (left-hand panel) and g1 (right-hand panel) galaxy cluster simulations at $z = 0$. Continuous blue lines show the case without AGN heating, while dashed red lines are for simulations where AGN feedback is included. It can be seen that the stellar density is decreased at all radii, and in the case of the g1 cluster, the central stellar density develops a core when AGN feedback is operating.

Table 4. Properties of the cD galaxy of the g676 and g1 galaxy clusters at $z = 0$. The listed values refer both to simulations with cooling and star formation only (denoted with ‘csf’), and to the runs where AGN heating was included (labelled ‘csfbh’). From the second to the fourth column, we give the stellar mass, the gas mass and the cold gas mass with $T < 1$ keV of the cD galaxy. The SFR of the cD galaxy is found in the fifth column, while the M_r magnitude and the $u - r$ and $g - r$ colours are given in the last three columns, respectively.

Cluster	M_{stars} ($h^{-1} M_{\odot}$)	M_{gas} ($h^{-1} M_{\odot}$)	$M_{\text{gas, cold}}$ ($h^{-1} M_{\odot}$)	SFR ($\text{yr}^{-1} M_{\odot}$)	M_r	$u - r$	$g - r$
g676_csf	1.3×10^{12}	2.3×10^{10}	1.9×10^{10}	30	−24.64	1.83	0.59
g676_csfbh	3.4×10^{11}	3.5×10^9	1.2×10^9	0	−22.55	2.68	0.84
g1_csf	4.2×10^{12}	8.4×10^{10}	7.8×10^{10}	163	−26.37	0.76	0.29
g1_csfbh	3.4×10^{11}	7.5×10^9	2.5×10^8	0	−22.71	2.64	0.83

metallicity in the cluster outskirts appears still too low in comparison with the metallicity levels of cool core clusters at similar radii (e.g. De Grandi & Molendi 2001). This discrepancy may in part be attributed to the fact that in these simulations we have not included feedback effects from galactic winds and outflows powered by star formation. As we have shown in Sijacki & Springel (2006), galactic winds help spreading metals throughout the cluster environment, especially into the outer parts. Another restriction of our model for metal enrichment lies in its highly simplified treatment of supernovae, where we neglect the time delay of Type Ia supernova. More sophisticated models of chemical enrichment (e.g. Tornatore et al. 2004; Scannapieco et al. 2005; Kobayashi, Springel & White 2007) that take this into account might therefore help as well.

We now turn to the effects of AGN feedback on the stellar properties of the simulated clusters. In Fig. 11, it can be seen that the stellar density is reduced at all radii in both clusters, as we have anticipated above. This reduction is caused not only by the activity of the central cluster BH but also by other BHs that are harboured at the centres of individual galaxies throughout the cosmic evolution. As we will discuss in more detail in Section 5.2, when a BH at the centre of a given galaxy becomes massive enough to influence its host, it acts on the stellar properties by reducing both the instantaneous SFR as well as the integrated total stellar mass by a significant amount.

In Table 4, we list a number of properties we have measured for the central cluster galaxy. For the identification of the cD galaxy, we have simply taken its radius to be $20 h^{-1}$ kpc, based on the apparent break in the stellar density profile at this radius. A number of interesting results can be noticed from the values in the table: due to the AGN feedback, the stellar mass of the cD galaxy is significantly reduced, as well as the central total gas mass. There is some residual cold gas in the cD galaxy, but star formation is completely quenched. This has an immediate and important consequence for the photometric colours of the cD galaxies. In order to estimate the colours, we have used the stellar population synthesis models of Bruzual & Charlot (2003) and computed the rest-frame magnitudes in the SDSS bands, assuming solar metallicity and a Chabrier initial mass function. Both cD galaxies in our cluster simulations become much less luminous in the r , u and g bands as a result of AGN feedback, such that the $u - r$ and $g - r$ colour indices are substantially increased. This reflects the fact that the stellar populations of the cD galaxies have become much older in our BH model, with very little recent star formation. With these properties, the simulated cD galaxies are in quite good agreement with observational findings (e.g. von der Linden et al. 2007), in terms of their stellar mass, their SFR and their colours.

Finally, we briefly discuss whether the Eddington limit that we have imposed on the BHAR has a significant effect on the BH growth

and the feedback in clusters. For this purpose we have rerun the g676 galaxy cluster simulation with exactly the same feedback parameters, but relaxing the maximum accretion rate assumption. In Fig. 9, we have included results for the radial profiles of gas density, temperature and entropy in this run as dotted green lines. It can be seen that the ICM properties are modified slightly more strongly in this case with respect to our default model, i.e. the AGN feedback effects appear to be somewhat stronger. A closer look at the detailed mass growth and accretion rate history of the BH as function of redshift shows that the feedback is regulating the BHAR even without an upper limit in the form of the Eddington rate. In fact, the total mass of all BHs at $z = 0$ in the simulated volume increases only by roughly 20 per cent if the Eddington limit is disregarded. The most massive BH in the cluster centre shows an increase of its mass of the same order. This additional growth of the BH mass mostly comes from brief episodes of very high accretion rate (typically reaching a few times up to 10 times \dot{M}_{Edd}), which are quickly terminated by the onset of AGN heating. Without an explicit Eddington limit, a bit more growth can happen before the ‘suicidal’ AGN activity shuts off the accretion, but the limit itself is not required for establishing a stable self-regulation loop. The Eddington limit therefore does not seem to pose an important restriction on the growth of BHs in clusters. However, this may still be quite different in the early phases of BH accretion at high redshifts, when the BH mass is so low that feedback effects are weak. Then the Eddington limit sets the shortest growth time-scale that can be realized, which makes it a challenge to grow BHs quickly enough to the observationally inferred large masses already at very high redshifts. We plan to address this interesting problem in full cosmological simulations in future work.

5 SIMULATIONS OF GALAXY FORMATION WITH AGN FEEDBACK

In this section we analyse BH growth and feedback in simulations of cosmic structure formation in homogeneously sampled volumes. We are mainly interested in the question whether our numerical model for a two-mode AGN growth produces realistic results for a range of object masses, from the scales of isolated galaxies to the ones of massive galaxy clusters. In Section 4, we have already confirmed that the BH model works well for clusters of galaxies, where it in fact drastically improves the properties of simulated galaxy clusters with respect to observations, and yields at the same time plausible evolutionary histories for the BH masses and the accretion rates. Here, we want to see whether this success also carries over to the scale of individual galaxies, where the ‘quasar’ mode of feedback will be dominant. In particular, we want to establish whether our unified model for AGN feedback maintains and extends

Table 5. Numerical parameters of our cosmological simulations in periodic boxes. The values listed in the second to the fifth column refer to the size of the box, to the number of gas and dark matter particles and to their masses. The last three columns give the initial and final redshifts of the runs, and their gravitational softening lengths ϵ .

Simulation	$L_{\text{box}} (h^{-1} \text{ Mpc})$	N_{part}	$m_{\text{DM}} (h^{-1} M_{\odot})$	$m_{\text{gas}} (h^{-1} M_{\odot})$	z_{start}	z_{end}	$\epsilon (h^{-1} \text{ kpc})$
R1	25	2×176^3	1.72×10^8	0.35×10^8	100	1	3.0
R2	25	2×256^3	5.58×10^7	1.14×10^7	100	1	2.0
R3	25	2×384^3	1.65×10^7	0.34×10^7	100	1	1.3

the successes we have found in Di Matteo et al. (2007) for the high-redshift galaxy population in a model that only accounts for the ‘quasar’ mode.

We have performed hydrodynamical simulations of a cosmological box with $25 h^{-1} \text{ Mpc}$ on a side, at three different mass resolutions, as summarized in Table 5. The small box size allows us to study comparatively small galaxies with stellar masses in the range of $\sim 10^8$ to $\sim 10^{11} M_{\odot}$. However, since the fundamental mode of this small box becomes non-linear at around $z = 1$, we had to stop the simulations at this redshift, since they would not be representative any more for the low-redshift Universe. For the initial conditions, we have used a ‘glass’ as unperturbed particle load and the power spectrum fit of Eisenstein & Hu (1999) for imposing initial perturbations with *WMAP-3* (*Wilkinson Microwave Anisotropy Probe*) cosmological parameters (Spergel et al. 2007). In particular, we have adopted a ΛCDM cosmology with $\Omega_{\text{m}} = 0.26$, $\Omega_{\text{b}} = 0.044$, $\Omega_{\Lambda} = 0.74$, $\sigma_8 = 0.75$ and $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ at the present epoch, and a primordial spectral index of $n_{\text{s}} = 0.938$.

For all three mass resolutions R1–R3 that we considered, we have computed two runs, one only with cooling and star formation and the other with the BH model included as well. The parameters of the BH model were exactly the same as the ones adopted in our cosmological simulations of galaxy cluster formation. For our intermediate resolution case R2, we have carried out additional simulations to study the influence of further parameters. This includes a run where we used the same cosmological parameters as in Section 4 in order to gauge the importance of the cosmological model. We have also explored the relative importance of the bubble feedback on galactic scales relative to the quasar feedback, that we will discuss more in detail in Section 6. Finally, we have studied the influence of includ-

ing a model for supernova-driven galactic winds with velocity $v \sim 480 \text{ km s}^{-1}$, which can be especially important in low-mass systems, as we will discuss below.

5.1 Black hole growth

In Fig. 12, we consider the comoving BH mass density (left-hand panel) and BHAR density (right-hand panel) as a function of cosmic time. For comparison, we also show the redshift evolution of the SFR density (right-hand panel, dashed lines), which we have rescaled by a constant factor of 10^3 to put the curves on to the same plot. It can be seen that the BHAR and SFR densities approach each other asymptotically towards low redshifts when the numerical resolution is increased. The fact that numerical convergence can be more easily achieved at low than at high redshift is not surprising in light of the hierarchical growth of structure. At high redshift, most of the small-mass systems that undergo star formation and BH accretion in our highest resolution simulation are unresolved in our low-resolution simulation, which hence underpredicts the BHAR and SFR density. However, towards later times, the mass scale where most of the accretion happens shifts to larger mass objects, which can be resolved in all of our simulations, such that approximate convergence can be reached even with moderate resolution. This behaviour is also similar to the one seen by Springel & Hernquist (2003b) in a comprehensive simulation study of the cosmic SFR history. We consider it highly encouraging that we find such a behaviour for our BH growth model as well, which shows that the model is numerically well posed and can be meaningfully applied to cosmological volumes with their comparatively poor resolution per galaxy.

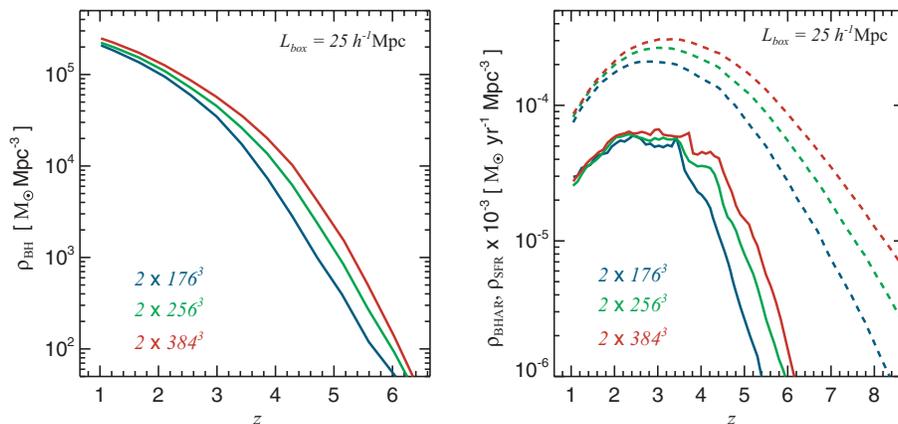


Figure 12. Redshift evolution of the comoving BH mass density in a cosmological box of $25 h^{-1} \text{ Mpc}$ on a side (left-hand panel). The three different curves show runs with increasing resolution as labelled in the panel. The right-hand panel gives the BHAR (continuous lines) and SFR (dashed lines) densities as a function of redshift for the same set of runs. For plotting purposes, ρ_{SFR} has been rescaled by a factor of 10^3 to fit on to the same scale. It can be seen that numerical convergence is achieved at low redshift, where BHAR and SFR densities of the different runs asymptotically approach each other with increasing resolution.

Nevertheless, it is clear that the numerical resolution in our cosmological runs remains too poor to correctly resolve the internal structure of galaxies, in particular the thin gaseous and stellar discs in spiral galaxies. This presumably affects the detailed time evolution of star formation and BH accretion in galaxy mergers. For example, our cosmological runs will largely blend over the time delay we see in high-resolution galaxy mergers between BH activity and the peak of the starburst. However, the final BH mass reached in a merger is comparatively insensitive to these details and probes primarily the depth of the gravitational potential of the forming spheroid, which our simulations can still represent quite well. While numerical resolution is clearly an important limitation in our cosmological results, we therefore believe they are reasonably accurate for the basic properties of the BH population.

Our estimated BH mass density at $z = 1$ in our highest resolution run is $2.5 \times 10^5 M_{\odot} \text{Mpc}^{-3}$. This is in very good agreement with a number of observational estimates, based both on optical and X-ray data, which typically find values in the range of $2\text{--}6 \times 10^5 M_{\odot} \text{Mpc}^{-3}$ (e.g. Fabian & Iwasawa 1999; Merritt & Ferrarese 2001; Yu & Tremaine 2002; Cowie et al. 2003; Graham et al. 2007). The SFR density shows a peak at $z \sim 3$, while the BHAR density peaks at somewhat lower redshift, i.e. $z \sim 2.5$. The BHAR and SFR then decline to lower redshifts, with the SFR density decreasing slightly more steeply. If we extrapolate the BHAR and SFR densities to $z = 0$, the ratio of the two is of order of 2×10^3 , a value that is roughly in agreement with estimates from the local population of galaxies (Heckman et al. 2004). In contrast, at high redshift, the BHAR density is rising more steeply than the SFR density, in broad agreement with the estimates by Merloni, Rudnick & Di Matteo (2004).

However, Hopkins, Richards & Hernquist (2007a) find that the quasar luminosity density declines more steeply for $z > 2$ than in our simulations. It is quite plausible that this discrepancy at high redshift is due to the limited numerical resolution available in our cosmological simulations, and in particular, their inability to correctly reproduce galactic discs of the right sizes, as discussed above. While the bulk of the BH accretion happens during major merger events in our simulations, we expect that there is a significant contribution from minor merger episodes or secular processes, which boost the BHAR at high z and induce a less steep rise. If the galaxies

could be fully spatially resolved, the gas would instead not be able to easily reach the galactic centres in these minor merger events. None the less, it is encouraging that the shapes and the peaks of the SFR and BHAR densities we find are significantly different, clearly indicating that the BHAR is not simply following the evolution of the SFR with time. Finally, by comparing the simulations with and without AGN feedback we conclude that the SFR density starts to be noticeably reduced by BH heating for $z < 4$, and at $z = 1$, this reduction is of order ~ 20 per cent.

In Di Matteo et al. (2007), similar results for the BH mass density and BHAR density are obtained for a cosmological box performed at somewhat higher numerical resolution, further supporting the trends we find here. The peak of the BHAR density appears moderately broader here than in Di Matteo et al. (2007), which could be a result of our smaller simulation box, or simply due to different choices of cosmological parameters, or due to the absence of galactic winds in Fig. 12. As we shall discuss in detail in Section 6, the ‘radio mode’ does not play a significant role at high redshift and hence is unlikely to modify the shape of the BHAR evolution at the peak of the quasar activity.

In the left-hand panel of Fig. 13, we show how the BHAR evolves with redshift as a function of the BH mass. We have already considered an analogous plot for our galaxy cluster simulations in Fig. 7. While the two plots are consistent for the overlapping range of BH masses, the most striking difference can be seen at the high BH mass end: in the galaxy cluster simulation we saw an upturn of the BHAR at low redshifts for very massive BHs, while here instead the BHAR keeps decreasing with increasing BH mass at all redshifts considered. This can be understood as a consequence of the limited size of the cosmological box. As a result of the small volume, rare big objects are absent, and hence massive BHs residing in cluster-sized objects are missing as well.

The characteristic time of BH growth, as a function of the BH mass, is illustrated in the right-hand panel of Fig. 13. For each BH, we have defined the characteristic growth time as the ratio of its mass to its accretion rate at the given redshift, normalized to the Hubble time at the present epoch. It can be seen that the BH growth time is shortest at high redshifts and for the smallest BHs. At $z = 1$, BHs with mass of order of $10^7 M_{\odot}$ reach a characteristic growth time of the order of the Hubble time for the first time. At different redshifts,

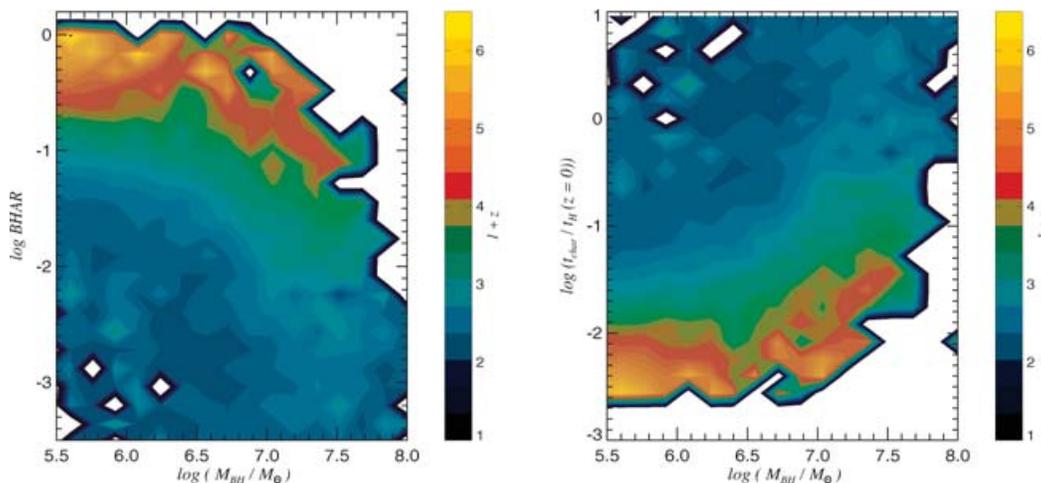


Figure 13. Left-hand panel: BHAR in Eddington units as a function of BH mass, for all BHs belonging to the R2 cosmological simulation. The colour coding expresses the redshift, and the redshift interval considered is from 5.0 to 1.0. Right-hand panel: characteristic growth time of the BHs as a function of their mass, colour coded according to redshift. The growth time has been defined as the ratio of the BH mass and the BHAR at the given epoch, normalized to the Hubble time at $z = 0$.

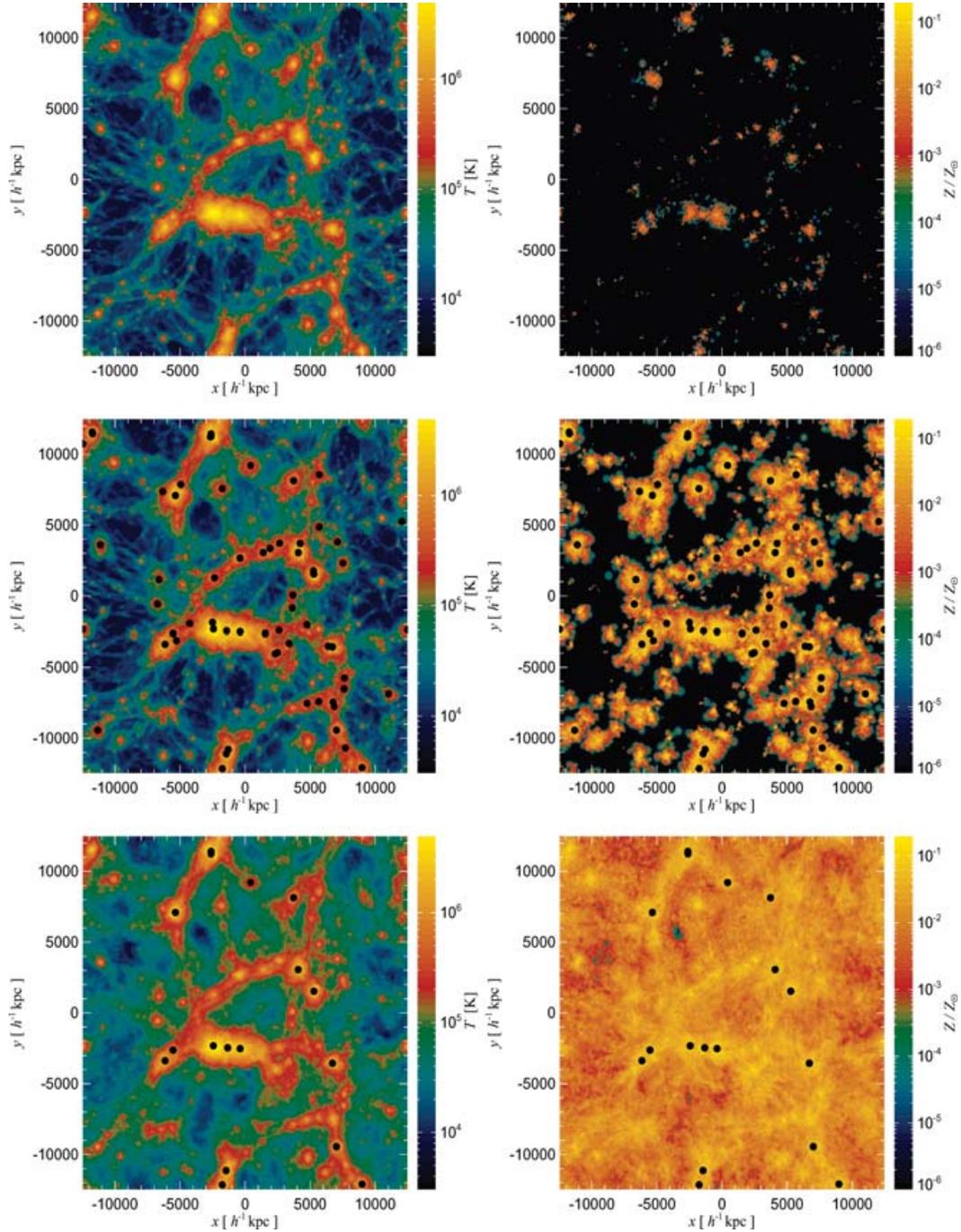


Figure 14. Projected mass-weighted temperature maps (left-hand panels) and gas metallicity maps (right-hand panels) at $z = 1$ of the R2 run. Top row: simulation with cooling and star formation only; middle row: simulation with additional BH model; bottom row: run with BH model and galactic winds ‘switched-on’. The positions of BH particles more massive than $2 \times 10^7 h^{-1} M_{\odot}$ are marked with black dots.

the $\log t_c - \log M_{\text{BH}}$ curve has a similar shape, which is consistent with results obtained for the local galaxy population (Heckman et al. 2004).

5.2 Effects on the intergalactic medium (IGM)

In Fig. 14, we show projected mass-weighted temperature and metallicity maps of our cosmological box at $z = 1$. In the upper row, the run with cooling and star formation only is plotted, while the

other two rows give results for simulations where the AGN feedback was included. In the lower row, feedback by galactic winds was additionally included. The positions of BHs more massive than $2 \times 10^7 h^{-1} M_{\odot}$ are marked with black dots. It can be seen that galactic winds modify the gas temperature and metallicity properties more significantly than the BH feedback alone. As expected, galactic winds are especially important in low-mass systems, given that they can expel gas from objects with small escape speeds relatively easily due to their shallow potential wells. Galactic

winds can therefore pollute a sizable fraction of the volume of the IGM with metals, to a level of 10^{-3} – 10^{-2} of the solar metallicity, in agreement with a number of observational findings (e.g. Rauch, Haehnelt & Steinmetz 1997; Cowie & Songaila 1998; Ellison et al. 2000; Schaye et al. 2000).

However, as can be seen from Fig. 14, AGN contribute to the process of metal pollution of the low-density gas, albeit with lower efficiency. If only AGN feedback is considered, we find that metals can be transported up to a distance of $\sim 1.5 h^{-1}$ Mpc from a galactic centre, but the large-scale metallicity distribution remains extremely patchy. Thus, the main difference relative to enrichment by galactic winds lies not only in the level of IGM metal enrichment but also in the volume filling factor of the metal-enriched regions. Nevertheless, by redshift $z = 3$, AGN feedback is making an important contribution to the enrichment of the IGM, where the characteristic size of high metallicity regions is of the order of $\sim 1 h^{-1}$ Mpc. However, we should point out that the AGN impact on the metal distribution found in our simulations is probably a lower limit, due to the fact that we cannot account for the contribution of very bright quasars powered by massive BHs at high redshifts.

Interestingly, even though AGN heating is affecting the gas metallicity of the IGM, it hardly affects its mean temperature, which is probably a reflection of the small volume filling factor of AGN-driven outflows. In contrast, galactic winds raise the intergalactic temperature significantly. Since this can delay the formation of small galaxies, the number and the mass of galaxies at a given epoch is reduced as well, which in turn slows the growth of the BH mass density.

5.3 Galaxy properties and evolution

We now study the question how the properties of galaxies are affected by BH feedback, and how the BH masses relate to their hosts. For this purpose, we identify the galaxies present in our cosmological box at different redshifts using a special group finder, and compute a number of properties for them, e.g. their stellar mass M_* , stellar velocity dispersion σ_* , their total SFR and their colours in the rest-frame SDSS bands. This also allows us to look for correlations with their central BH masses. Our scheme for identifying simulated galaxies as isolated groups of stars is based on a simplified variant of the SUBFIND algorithm (Springel et al. 2001a). We first compute an adaptively smoothed baryonic density field for all stars and gas particles, allowing us to robustly identify centres of individual galaxies as isolated density peaks. Using SUBFIND's approach for growing these centres by attaching nearby particles of even lower density, we then obtain a list of cleanly separated galaxies. By restricting the set of gas particles processed in this way to those that have at least a density of 0.01 times the threshold density for star formation, we avoid that spatially separated galaxies are linked together, while all the stars present in the simulation are assigned to the different galaxies. A gravitational unbinding procedure is then not necessary in this procedure. Note that we are not interested in the gaseous components of galaxies here. We only include the gas particles in the galaxy finding procedure because the dense, cold gas that most galaxies contain makes the method more robust.

In Fig. 15, we show the relation between BH mass and stellar velocity dispersion σ_* (left-hand panels), and between BH mass and M_* (right-hand panels), at three different redshifts, $z = 1, 2$ and 3. Here we evaluate M_* and σ_* within the effective radius, R_e , chosen as the half-mass radius. Green star symbols are for the run without galactic winds, while red diamonds are for the simulation that also includes galactic winds. The results shown are from our

intermediate-resolution box, but we have verified good convergence in these quantities by comparing with the R3 box. With the dashed lines we overplot the locally observed $M_{\text{BH}}-\sigma_*$ and $M_{\text{BH}}-M_*$ relations, as determined by Tremaine et al. (2002) and Häring & Rix (2004), respectively. A number of interesting features are noteworthy in this figure.

(i) Both relations show some evolution with redshift that however seems most prominent at the low BH mass end and for the run with galactic winds. The most massive objects found at every epoch stay very close to the local relation over the redshift interval considered.

(ii) At high redshifts, galactic winds give rise to less massive galaxies with somewhat lower σ_* , and for a given M_* or σ_* , BHs are less massive when the winds are ‘switched-on’. However, by redshift $z = 1$ and for $M_{\text{BH}} > 5 \times 10^6 M_\odot$, the central BH masses are comparable for a given stellar mass of a galaxy, regardless of the presence of winds. This suggests that the feedback by galactic winds is delaying BH growth by expelling significant amounts of gas at high redshift. However, this gas becomes available for BH accretion at latter epochs, when it has been reincorporated from the IGM back into the galaxies.

(iii) For BH masses smaller than $5 \times 10^6 M_\odot$ (as indicated with the arrow in Fig. 15) there is a tilt in both relations, especially noticeable for the $M_{\text{BH}}-M_*$ relation. This tilt is significantly reduced in the run with galactic winds. This can be understood by considering that the AGN feedback in these small galaxies is very modest, and does not modify the properties of the hosts significantly. On the other hand, galactic winds are very efficient for these low-mass objects, reducing their stellar mass and bringing them into better agreement with the expected relation. However, we caution that this result of our modelling may be sensitive to our prescription for BH seeding and the early initial growth, which could be affected by numerical resolution effects. If small mass galaxies could grow somewhat bigger BHs in their centre, they would also be more affected by BH feedback so that it may not be necessary to invoke galactic winds to match the low-mass end of the $M_{\text{BH}}-\sigma_*$ relation. In any case, it is interesting that galactic winds shift these smaller objects in the desired direction while at the same time more massive galaxies are unaffected and still lie on the local relation.

(iv) Finally, given that the simulated $M_{\text{BH}}-M_*$ and $M_{\text{BH}}-\sigma_*$ relations at $z = 1$ match the locally observed relations well, it would be interesting to explore how the model will evolve further until $z = 0$. It seems quite likely, especially for more massive galaxies, that not much evolution will occur for $z < 1$, given that these galaxies have a reduced star formation and depleted gas content due to the AGN feedback. However, considering the complex interplay between BHs and galactic winds, a reliable answer of this question will require explicit numerical simulations of a bigger cosmological box at comparable or better resolution, which we will tackle in future work.

In Di Matteo et al. (2007), a measurement of the $M_{\text{BH}}-M_*$ and $M_{\text{BH}}-\sigma_*$ relations is presented as well, for an overlapping redshift range. Their results extend to somewhat higher BH mass, due to their larger box size and their higher normalization σ_8 of the power spectrum. In general, we find a reassuring agreement with the findings of Di Matteo et al. (2007), indicating that the $M_{\text{BH}}-M_*$ and $M_{\text{BH}}-\sigma_*$ relations are not significantly affected by the bubble feedback or the underlying cosmology, as we will discuss latter on in more detail. Only for the $M_{\text{BH}}-\sigma_*$ relations at $z = 3$, our measurements lie somewhat below those of Di Matteo et al. (2007), a trend that we attribute to the fact that we required dark matter haloes to

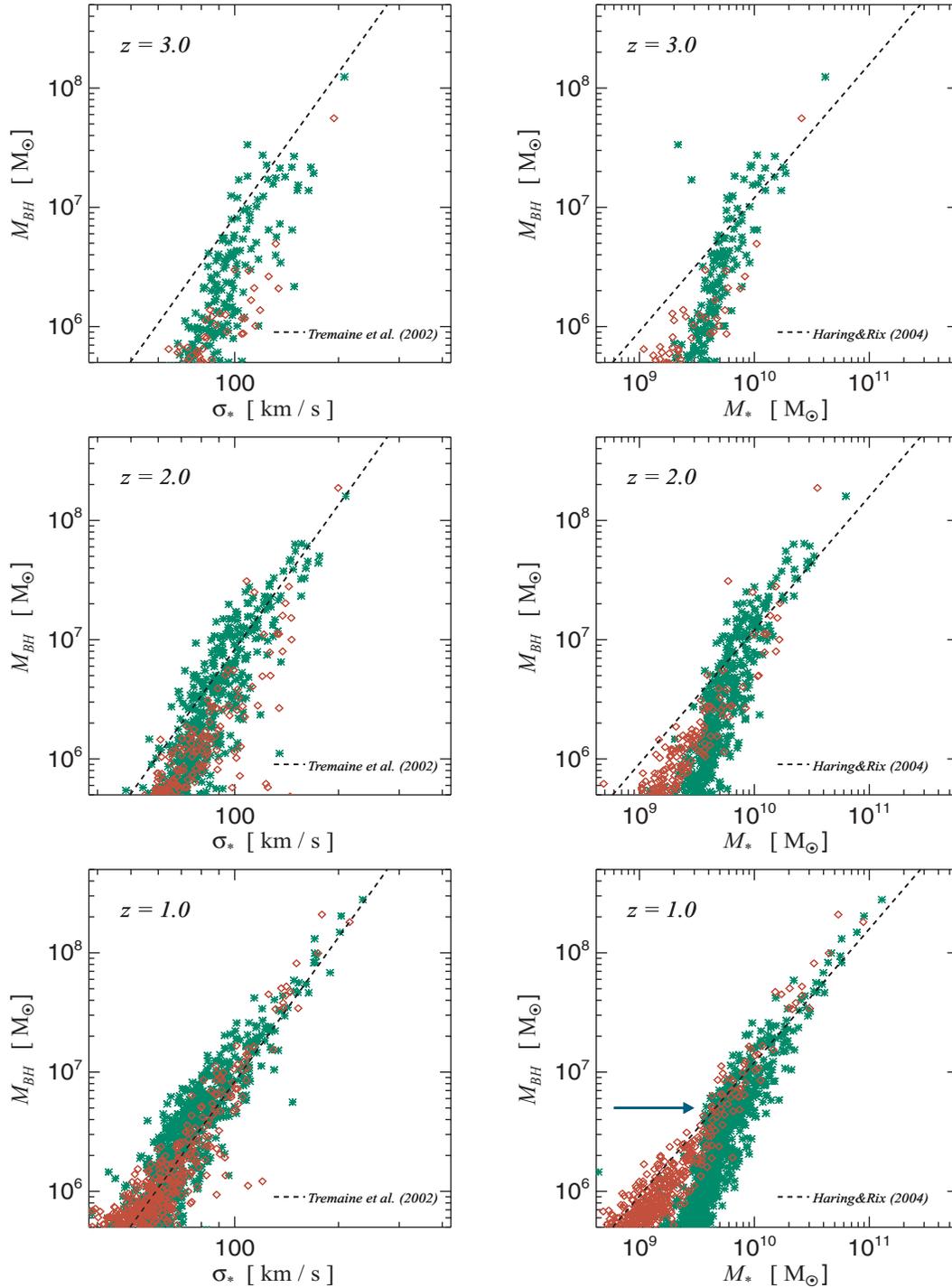


Figure 15. BH mass–stellar velocity dispersion relation (left-hand panels) and BH mass–stellar mass relation (right-hand panels), at redshifts $z = 1, 2$ and 3 . Green star symbols denote the run without galactic winds, while red diamonds are for a run where galactic winds have been included as well. Both simulations have been performed at the resolution of the R2 cosmological box. The dashed lines give the locally observed relationships between the considered quantities, as determined by Tremaine et al. (2002) and Häring & Rix (2004).

be five times more massive than Di Matteo et al. (2007) before we endowed them with a seed BH, which may have delayed their early growth.

In Fig. 16, we show the specific SFRs and the $g-r$ colours of our simulated galaxies at $z = 1$ as a function of their stellar mass. Blue dots and blue continuous contours show the results when only cool-

ing and star formation is considered. Green dots and long-dashed contours give the results for the run with AGN feedback, while red dots and short-dashed contours represent the case with additional galactic winds. While AGN feedback does not affect the bulk of the population significantly, it is evident that the stellar masses and the SFRs of most massive galaxies are reduced. When feedback

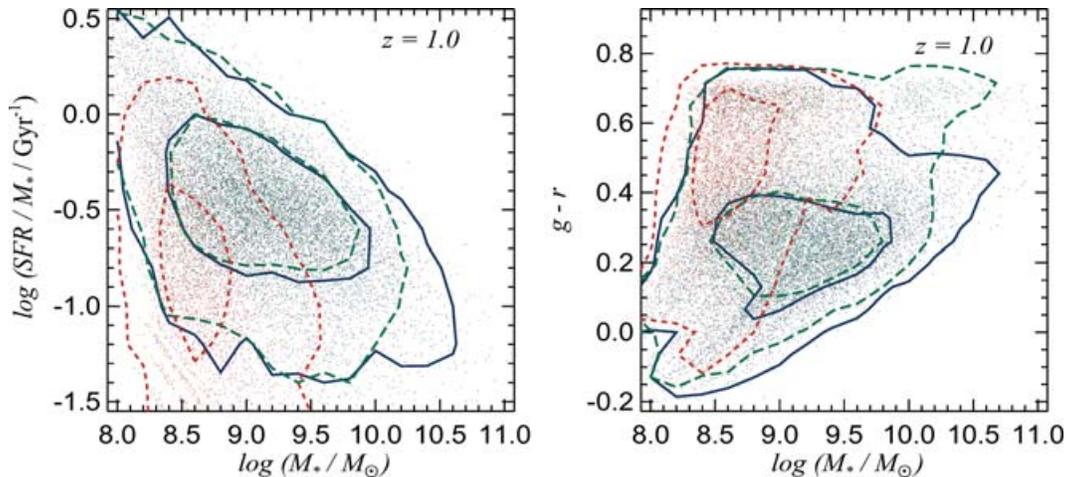


Figure 16. Specific SFR (left-hand panel) and $g-r$ colour (right-hand panel) for simulated galaxies as a function of stellar mass at $z = 1$. Blue dots and continuous contours are for the run with cooling and star formation only. Green dots and long-dashed contours indicate the case where our BH model was included, while red dots and short-dashed contours are for the simulation which in addition contains galactic winds.

by galactic winds is included as well this trend is even more pronounced, and extends to the lower mass systems. Note also that the galaxy colours change in our BH model – a number of galaxies become much redder, forming a red sequence, as it is shown in the right-hand plot of Fig. 16. However, if the galactic wind model is included, the clearly defined red sequence disappears. Instead, all galaxies are in general less massive and redder. The fact that the red sequence becomes largely unpopulated in the run with galactic winds is not necessarily problematic, however. A larger cosmological box may again fill this region of the diagram with a population of more massive galaxies.

Finally, in order to assess whether the colours of our galaxies in the AGN feedback model are realistic, we have computed the mean $u-r$ colour of the red sequence at $z = 1$, obtaining $u-r \sim 2.2$. We can compare this result to the work of Bell et al. (2004) on the COMBO-17 (Classifying Objects by Medium-Band Observations in 17 Filters) survey. Using the conversion between $U-V$ and $u-r$ colours suggested by these authors, we find that the colour of our simulated galaxies along the red sequence is in good agreement with the early-type galaxies from the COMBO-17 survey.

5.4 Dependence on cosmological parameters

In order to evaluate how sensitively our BH model depends on the underlying cosmological model, we have performed two additional simulations at intermediate resolution (R2) with exactly identical BH parameters. In one case we have adopted cosmological parameters consistent with the *WMAP* first year data analysis, while for the other run we have selected the updated parameters of *WMAP*'s third year data release. The main difference between the corresponding cosmological models lies in a reduction of the normalization parameter σ_8 ($0.9 \rightarrow 0.75$), a lowering of the matter density Ω_m ($0.3 \rightarrow 0.26$) and in the introduction of a tilt in the primordial power spectrum slope n_s ($1.0 \rightarrow 0.938$).

Our results show that BHs form earlier in the *WMAP*-1 cosmology, as expected from the fact that host haloes above a given mass threshold appear earlier in this model due to the higher normalization, and are therefore seeded with a BH at higher z . Furthermore, the number density of BHs is higher and there are more high-mass BHs at lower redshift with respect to the *WMAP*-3 run. Consequently,

the comoving BH mass density is somewhat higher in the *WMAP*-1 cosmology. At $z = 1$, it is $3.9 \times 10^5 M_\odot \text{Mpc}^{-3}$, while for *WMAP*-3 we obtain $2.2 \times 10^5 M_\odot \text{Mpc}^{-3}$ at the same numerical resolution. The BHAR density is higher in the *WMAP*-1 model as well, with its peak shifted to a slightly earlier redshift of ~ 3.5 . However, for lower redshifts this is compensated by a steeper decline, such that at $z = 1$ the BHAR densities have a very similar value in both cosmologies. These trends can simply be understood as a result of ‘delayed’ structure formation in the *WMAP*-3 cosmology. Increasing σ_8 will let the BHs form earlier, evolve faster and reach their peak activity at higher redshift. Interestingly, within the scatter, the $M_{\text{BH}}-\sigma_*$ and $M_{\text{BH}}-M_*$ relations are however not significantly affected by the change in the cosmological parameters.

6 RELEVANCE OF RADIO VERSUS QUASAR FEEDBACK MODE

In this section, we explore the relative importance of our two different modes of BH growth and quasar feedback for the build-up of the cosmological BH mass density. For this purpose, we show in Fig. 17 the distribution function of the net BH mass accretion rate in $M_\odot \text{yr}^{-1}$, as a function of the Eddington-normalized accretion rate. In the left-hand panel, we give results for our cosmological box R2 at redshifts $z = 1, 2$ and 3, while on the right-hand side we show results for the g676 galaxy cluster simulation at different redshifts.

It can be seen that at $z = 3$ the bulk of the BH growth occurs in the ‘quasar’ mode, where the accretion rate is more than 0.01 times the Eddington rate. In fact, the ‘radio’ mode contributes less than 2 per cent to the total BHAR at this epoch, and for the g676 cluster simulation, the contribution from growth in this radiatively inefficient mode is similarly small (~ 1 per cent). At lower redshifts, BHs accreting at low rates make up for an increasingly higher fraction of the total BHAR density (which however declines with time). For the R2 simulation at $z = 2$ they contribute ~ 13 per cent, and at $z = 1$ they are at ~ 37 per cent. Instead, for the g676 galaxy cluster, these numbers are somewhat higher, with 65 per cent at $z = 1$ and 96 per cent at $z = 0$ coming from BHs in a low accretion state. Thus, while in the cosmological box BHs in the quasar regime are the main channel of BH growth at all epochs considered, for the

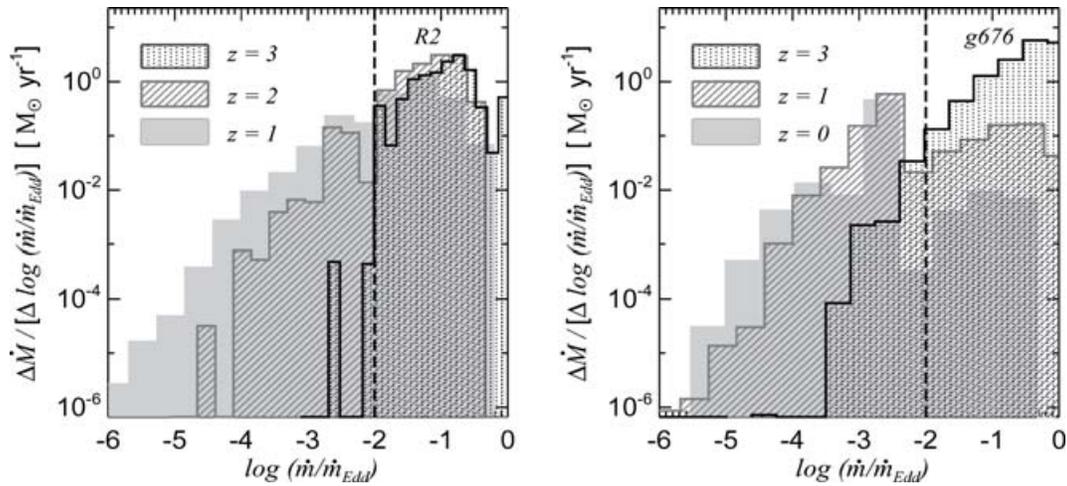


Figure 17. Distribution of the BHAR, expressed in $M_{\odot} \text{yr}^{-1}$, as a function of accretion rate in Eddington units, for three different redshifts, as indicated on the panels. The left-hand panel shows our measurements for the cosmological box simulation at the intermediate resolution (R2), while the right-hand panel gives analogous results for the g676 galaxy cluster simulation. The vertical dashed line denotes the transition adopted in our model between BHs accreting in the ‘quasar’ mode and those accreting in the radiatively inefficient ‘radio’ mode.

galaxy cluster run there appears to be a transition epoch at $z \sim 1$, below which BHs in the ‘radio’ mode are driving the BH growth at low redshifts. This difference between the BHAR distributions for the R2 run and the g676 cluster simulation can be explained by the presence of a dominating, very massive BH in the centre of the main progenitor of the cluster simulation, while the cosmological box contains a much fairer sample of the BH mass function.

However, integrating the BHAR over cosmic time for all BHs in the simulated volumes, we find that the bulk of BH mass is always grown during phases of quasar activity, i.e. in the high accretion rate regime. In fact, BH accreting in the ‘radio’ mode contribute less than ~ 5 per cent to the integrated BH mass density in the cosmological box, while for the g676 cluster simulation this number is ~ 20 per cent. Performing the same calculations for our g1 galaxy cluster simulation we obtain a very similar result (~ 15 per cent), further strengthening this point. Thus, we conclude that most of the BH mass is assembled during periods of rapid growth in the radiatively efficient ‘quasar’ mode. This implies that our models are consistent with the Soltan argument (Soltan 1982; Yu & Tremaine 2002; Marconi et al. 2004; Merloni 2004; Hopkins, Narayan & Hernquist 2006b), and shows that merger-driven quasar activity is still dominating the BH assembly in our unified model for AGN feedback.

Furthermore, by repeating our cosmological box simulations without ‘radio’ mode feedback we have explicitly checked that the inclusion of bubble feedback has a negligible influence on the $M_{\text{BH}}-M_*$ and $M_{\text{BH}}-\sigma_*$ relations. Only at $z < 2$, the ‘radio’ mode makes BHs slightly less massive due to the somewhat more efficient heating by the bubbles. Thus, quasar activity triggered by the merging of host galaxies is the primary process responsible for establishing the scaling relationships between the central supermassive BHs and their host galaxies, and this process remains intact and is essentially unaffected if an additional ‘radio’ mode feedback is invoked in the low accretion state of BHs.

On the other hand, our simulations also show that the ‘radio’ mode feedback in the form of AGN-driven bubbles is essential in massive groups and clusters. In order to demonstrate this point directly, we have rerun our g676 galaxy cluster simulation by switching off the radio channel of feedback mode, only allowing the ordinary

quasar activity to proceed until $z = 0$. This results in simulated ICM properties that are in disagreement with observations. In particular, the gas temperature profile keeps rising towards the very centre. Thus the intermittent nature of our bubble feedback appears necessary in this respect, as well of course for explaining the observed phenomena of radio galaxies and X-ray cavities in clusters of galaxies.

7 DISCUSSION AND CONCLUSIONS

In this study, we have proposed a new model for following BH growth and feedback in cosmological simulations of structure formation. Within our prescription, BH seeds are introduced at early cosmic times in the centres of all haloes once their mass exceeds a certain threshold value for the first time. The BH seeding is accomplished by frequently running a FOF algorithm on the fly in our simulation code, which identifies newly formed haloes and their properties. The BHs themselves are represented as collisionless sink particles that can grow their mass via two channels, by (1) gas accretion with an accretion rate estimated by a simple Bondi formula, and (2) via mergers with other BHs that are close enough. Motivated by the observed phenomenology of radio galaxies and quasars, and by analogy with X-ray binaries, we have assumed that the AGN feedback that accompanies gas accretion can be decomposed into two physically distinct modes: BHs accreting at high accretion rates in terms of their Eddington limit are operating in a radiatively efficient ‘quasar’ regime, while BHs accreting at much lower rates live in a radiatively inefficient ‘radio’ mode, which however produces significant mechanical luminosity that gives rise to jet-inflated bubbles.

In our initial tests of the model, we have applied it to simulations of isolated galaxy clusters consisting of a static dark matter potential and gas in hydrostatic equilibrium. Considering a range of cluster masses from 10^{13} to $10^{15} h^{-1} M_{\odot}$, we have found that the model leads to a self-regulated AGN feedback that produces realistic ICM properties and prevents the overcooling problem in the central cluster regions. At the same time, it yields reasonable BH masses and accretion rates.

We have then applied our model to fully self-consistent cosmological simulations of galaxy cluster formation, where the seeding of BHs and their subsequent growth is followed self-consistently. Together with our work in Di Matteo et al. (2007), these simulations are the first cosmological hydrodynamical models that simultaneously follow the growth of structure in the dark matter as well as the baryonic physics of star formation and AGN feedback. These simulations provide information on how the BHs are spatially distributed in the simulated volume, how they grow with cosmic time as a function of their environment and how they influence the surrounding medium and their host galaxies.

Our analysis of these calculations has focused on the history of the BH in the central cD galaxy. We have shown that while essentially all the mass in supermassive BHs originates in gas accretion, a large fraction of the mass of the BH in the cD grows from mergers with other BHs, as opposed to being all accreted by a single massive progenitor. This suggests that mergers of supermassive BHs are an important factor for building up very massive BHs. In line with our findings for the isolated cluster simulations we found that feedback from BHs is preventing the formation of excessive cooling flows out of hot cluster atmospheres, and is changing the central ICM properties substantially in the process. These changes are all acting to bring the simulated temperature and entropy profiles into much better agreement with observational findings of cool core clusters (e.g. De Grandi & Molendi 2002; Ponman et al. 2003; Pratt & Arnaud 2003; Bauer et al. 2005; Vikhlinin et al. 2005; Dunn & Fabian 2006; Sanderson et al. 2006; Pratt et al. 2007). Moreover, the stellar properties of the central galaxy are affected as well, with the forming cD galaxies being less massive, and having older and redder stellar populations which are consistent with the results found in recent studies of a large sample of BCGs in the SDSS survey (von der Linden et al. 2007).

We have also studied our unified AGN feedback model in simulations of homogeneously sampled cosmological volumes, thereby investigating how successful it is on the scale of galaxies, and with respect to the mean properties of the cosmic population of BHs. In our simulation box of size $25 h^{-1}$ Mpc on a side, we were able to resolve galaxies with stellar masses greater than $\sim 10^8 M_{\odot}$. We have found that our model produces a comoving BH mass density equal to $2.8 \times 10^5 M_{\odot} \text{Mpc}^{-3}$, which is in very good agreement with observational estimates (Fabian & Iwasawa 1999; Merritt & Ferrarese 2001; Yu & Tremaine 2002; Cowie et al. 2003). Moreover, the AGN feedback influences the properties and the formation of the host galaxies and vice versa, and this mutual coupling establishes $M_{\text{BH}}-\sigma_*$ and $M_{\text{BH}}-M_*$ relationships already at early epochs. Again, these relations are in broad agreement with observations, at least for the more massive and better resolved systems. Interestingly, a model that besides BH feedback includes feedback from supernova-driven galactic winds produces a better match to the expected relationships at the low-mass end. At the same time, this model also leads to a much more widespread enrichment of the IGM, which appears required by the data on quasar absorption line systems. Additionally, AGN heating reduces the stellar mass of the most luminous galaxies identified in the simulated volume, and it affects their SFR and colours as well. The galaxies in simulations with AGN feedback show a clearly defined red sequence already at $z = 1$.

Finally, we have explored the relative importance of ‘radio’ versus ‘quasar’ mode. We have found that the bulk of the BH growth occurs at high accretion rates, corresponding to radiatively efficient AGN activity. While the relative importance of the ‘radio’ mode grows towards late times, and becomes large in clusters of galaxies at $z < 1$,

this mode of accretion is unimportant for the total BH mass density today. This also implies that our model is consistent with the Soltan argument. We also found that the feedback of the quasar regime is the key factor for establishing the $M_{\text{BH}}-\sigma_*$ and $M_{\text{BH}}-M_*$ relationships, a process that appears unaffected by the bubble feedback. On the other hand, the mechanical luminosity of AGN-driven bubbles in groups and cluster of galaxies is necessary in order to successfully reproduce the observed ICM properties.

These results represent highly encouraging successes for hydrodynamical models of hierarchical galaxy formation in Λ CDM cosmologies, and provide support for the theoretical conjecture that galaxy formation and supermassive BH growth are intimately linked. In spite of these successes we need to emphasize that our BH model clearly represents a drastically simplified picture of BH physics. In part, the simplifications are driven by numerical limitations, because even state-of-the-art simulations of galaxy formation in cosmologically representative regions of the Universe cannot directly resolve the gravitational sphere of influence of individual BHs. We are therefore forced to adopt a subresolution approach for the representation of BH accretion, and also for the modelling of bubble feedback where the initial stages of bubble inflation by a jet are not resolved. However, this subresolution approach can still capture the essential parts of the physics relevant for the coupling of the BHs to their larger-scale environment, and this is what allows us to study their cosmological significance for galaxy formation.

While we have tried in this study to shed some light on to the entangled histories of galaxies and their central BHs, a number of important questions remain open for future work, and our model appears suitable to help answering them. They include (i) the early growth of BHs – can the observed ‘downsizing’ of BH growth be reproduced by our model? (ii) Do the galaxy cluster scaling relations change as a result of the AGN feedback? (iii) Can a galaxy luminosity function with a bright end consistent with observational constraints be reproduced in cosmological simulations, and what is the respective AGN luminosity function? (iv) Do the quasar clustering properties come out right?

Indeed, AGN appear to be a key ingredient of structure formation, and future numerical simulations should be a very helpful theoretical tool to unveil the intriguing and complex picture of their interactions and growth.

ACKNOWLEDGMENTS

We are grateful to Simon White, Eugene Churazov, Andrea Merloni and Martin Haehnelt for very stimulating discussions and useful comments on the manuscript. We thank Klaus Dolag for providing us with galaxy cluster initial conditions used in this study. DS acknowledges the PhD fellowship of the International Max Planck Research School in Astrophysics, and the support received from a Marie Curie Host Fellowship for Early Stage Research Training.

REFERENCES

- Allen S. W., Dunn R. J. H., Fabian A. C., Taylor G. B., Reynolds C. S., 2006, *MNRAS*, 372, 21
- Barger A. J., Cowie L. L., Mushotzky R. F., Yang Y., Wang W.-H., Steffen A. T., Capak P., 2005, *AJ*, 129, 578
- Barnes J. E., Hernquist L., 1992, *ARA&A*, 30, 705
- Bauer F. E., Fabian A. C., Sanders J. S., Allen S. W., Johnstone R. M., 2005, *MNRAS*, 359, 1481

- Begelman M. C., Cioffi D. F., 1989, *ApJ*, 345, L21
- Bell E. F. et al., 2004, *ApJ*, 608, 752
- Binney J., Tabor G., 1995, *MNRAS*, 276, 663
- Birzan L., Rafferty D. A., McNamara B. R., Wise M. W., Nulsen P. E. J., 2004, *ApJ*, 607, 800
- Blandford R. D., Begelman M. C., 1999, *MNRAS*, 303, L1
- Bondi H., 1952, *MNRAS*, 112, 195
- Bondi H., Hoyle F., 1944, *MNRAS*, 104, 273
- Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000
- Cappelluti N., Cappi M., Dadina M., Malaguti G., Branchesi M., D'Elia V., Palombo G. G. C., 2005, *A&A*, 430, 39
- Cappi M. et al., 2001, *ApJ*, 548, 624
- Chartas G., Brandt W. N., Gallagher S. C., 2003, *ApJ*, 595, 85
- Churazov E., Brügggen M., Kaiser C. R., Böhringer H., Forman W., 2001, *ApJ*, 554, 261
- Churazov E., Sunyaev R., Forman W., Böhringer H., 2002, *MNRAS*, 332, 729
- Churazov E., Sazonov S., Sunyaev R., Forman W., Jones C., Böhringer H., 2005, *MNRAS*, 363, L91
- Cowie L. L., Songaila A., 1998, *Nat*, 394, 44
- Cowie L. L., Barger A. J., Bautz M. W., Brandt W. N., Garmire G. P., 2003, *ApJ*, 584, L57
- Cox T. J., Di Matteo T., Hernquist L., Hopkins P. F., Robertson B., Springel V., 2006a, *ApJ*, 643, 692
- Cox T. J., Dutta S. N., Di Matteo T., Hernquist L., Hopkins P. F., Robertson B., Springel V., 2006b, *ApJ*, 650, 791
- Crenshaw D. M., Kraemer S. B., George I. M., 2003, *ARA&A*, 41, 117
- Croton D. J. et al., 2006, *MNRAS*, 365, 11
- Dalla Vecchia C., Bower R. G., Theuns T., Balogh M. L., Mazzotta P., Frenk C. S., 2004, *MNRAS*, 355, 995
- De Grandi S., Molendi S., 2001, *ApJ*, 551, 153
- De Grandi S., Molendi S., 2002, *ApJ*, 567, 163
- Di Matteo T., Springel V., Hernquist L., 2005, *Nat*, 433, 604
- Di Matteo T., Colberg J., Springel V., Hernquist L., Sijacki D., 2007, preprint *Dolag K.*, 2004, in Reiprich T. H., Kempner J. C., Soker N., eds, *The Riddle of Cooling Flows in Galaxies and Clusters of Galaxies*. Available online at <http://www.astro.virginia.edu/coolflow/>
- Dunn R. J. H., Fabian A. C., 2006, *MNRAS*, 373, 959
- Eisenstein D. J., Hu W., 1999, *ApJ*, 511, 5
- Ellison S. L., Songaila A., Schaye J., Pettini M., 2000, *AJ*, 120, 1175
- Fabian A. C., Iwasawa K., 1999, *MNRAS*, 303, L34
- Fabian A. C., Sanders J. S., Taylor G. B., Allen S. W., Crawford C. S., Johnstone R. M., Iwasawa K., 2006, *MNRAS*, 366, 417
- Fan X. et al., 2001, *AJ*, 122, 2833
- Fender R. et al., 1999, *ApJ*, 519, L165
- Ferrarese L., Merritt D., 2000, *ApJ*, 539, L9
- Forman W. et al., 2006, preprint (astro-ph/0604583)
- Gallo E., Fender R. P., Pooley G. G., 2003, *MNRAS*, 344, 60
- Gebhardt K. et al., 2000, *ApJ*, 539, L13
- Graham A. W., Driver S. P., Allen P. D., Liske J., 2007, *MNRAS*, 378, 198
- Häring N., Rix H.-W., 2004, *ApJ*, 604, L89
- Hasinger G., Miyaji T., Schmidt M., 2005, *A&A*, 441, 417
- Heckman T. M., Kauffmann G., Brinchmann J., Charlot S., Tremonti C., White S. D. M., 2004, *ApJ*, 613, 109
- Heinz S., Reynolds C. S., Begelman M. C., 1998, *ApJ*, 501, 126
- Heinz S., Merloni A., Di Matteo T., Sunyaev R., 2005, *Ap&SS*, 300, 15
- Hopkins P. F., Hernquist L., Cox T. J., Di Matteo T., Martini P., Robertson B., Springel V., 2005, *ApJ*, 630, 705
- Hopkins P. F., Hernquist L., Cox T. J., Di Matteo T., Robertson B., Springel V., 2006a, *ApJS*, 163, 1
- Hopkins P. F., Narayan R., Hernquist L., 2006b, *ApJ*, 643, 641
- Hopkins P. F., Somerville R. S., Hernquist L., Cox T. J., Robertson B., Li Y., 2006c, *ApJ*, 652, 864
- Hopkins P. F., Richards G. T., Hernquist L., 2007a, *ApJ*, 654, 731
- Hopkins P. F., Hernquist L., Cox T. J., Robertson B., Krause E., 2007b, preprint (astro-ph/0701351)
- Hoyle F., Lyttleton R. A., 1939, *Proc. Camb. Phil. Soc.*, 34, 405
- Jenkins A., Frenk C. S., White S. D. M., Colberg J. M., Cole S., Evrard A. E., Couchman H. M. P., Yoshida N., 2001, *MNRAS*, 321, 372
- Katz N., Weinberg D. H., Hernquist L., 1996, *ApJS*, 105, 19
- Kauffmann G., Haehnelt M., 2000, *MNRAS*, 311, 576
- Kawata D., Gibson B. K., 2005, *MNRAS*, 358, L16
- King A., 2003, *ApJ*, 596, L27
- Kobayashi C., Springel V., White S. D. M., 2007, *MNRAS*, 376, 225
- Körding E. G., Jester S., Fender R., 2006, *MNRAS*, 372, 1366
- Kormendy J., Richstone D., 1995, *ARA&A*, 33, 581
- Li Y. et al., 2006, preprint (astro-ph/0608190)
- Maccarone T. J., Gallo E., Fender R., 2003, *MNRAS*, 345, L19
- McKinney J. C., 2006, *MNRAS*, 368, 1561
- McNamara B. R., Nulsen P. E. J., Wise M. W., Rafferty D. A., Carilli C., Sarazin C. L., Blanton E. L., 2005, *Nat*, 433, 45
- Magorrian J. et al., 1998, *AJ*, 115, 2285
- Marconi A., Hunt L. K., 2003, *ApJ*, 589, L21
- Marconi A., Risaliti G., Gilli R., Hunt L. K., Maiolino R., Salvati M., 2004, *MNRAS*, 351, 169
- Martini P., Kelson D. D., Kim E., Mulchaey J. S., Athey A. A., 2006, *ApJ*, 644, 116
- Merloni A., 2004, *MNRAS*, 353, 1035
- Merloni A., Heinz S., 2006, preprint (astro-ph/0611202)
- Merloni A., Rudnick G., Di Matteo T., 2004, *MNRAS*, 354, L37
- Merritt D., Ferrarese L., 2001, *MNRAS*, 320, L30
- Narayan R., Yi I., 1994, *ApJ*, 428, L13
- Navarro J. F., Frenk C. S., White S. D. M., 1996, *ApJ*, 462, 563
- Navarro J. F., Frenk C. S., White S. D. M., 1997, *ApJ*, 490, 493
- Okamoto T., Nemmen R. S., Bower R. G., 2007, preprint (astro-ph/0704.1218)
- Omma H., Binney J., 2004, *MNRAS*, 350, L13
- Ponman T. J., Sanderson A. J. R., Finoguenov A., 2003, *MNRAS*, 343, 331
- Pounds K. A., Reeves J. N., King A. R., Page K. L., O'Brien P. T., Turner M. J. L., 2003, *MNRAS*, 345, 705
- Pratt G. W., Arnaud M., 2003, *A&A*, 408, 1
- Pratt G. W., Böhringer H., Croston J. H., Arnaud M., Borgani S., Finoguenov A., Temple R. F., 2007, *A&A*, 461, 71
- Proga D., 2003, *ApJ*, 585, 406
- Quilis V., Bower R. G., Balogh M. L., 2001, *MNRAS*, 328, 1091
- Rauch M., Haehnelt M. G., Steinmetz M., 1997, *ApJ*, 481, 601
- Robertson B., Cox T. J., Hernquist L., Franx M., Hopkins P. F., Martini P., Springel V., 2006a, *ApJ*, 641, 21
- Robertson B., Hernquist L., Cox T. J., Di Matteo T., Hopkins P. F., Martini P., Springel V., 2006b, *ApJ*, 641, 90
- Ruderman J. T., Ebeling H., 2005, *ApJ*, 623, L81
- Ruszkowski M., Begelman M. C., 2002, *ApJ*, 581, 223
- Sanderson A. J. R., Ponman T. J., O'Sullivan E., 2006, *MNRAS*, 372, 1496
- Scannapieco C., Tissera P. B., White S. D. M., Springel V., 2005, *MNRAS*, 364, 552
- Schaye J., Rauch M., Sargent W. L. W., Kim T.-S., 2000, *ApJ*, 541, L1
- Scheuer P. A. G., 1974, *MNRAS*, 166, 513
- Shakura N. I., Sunyaev R. A., 1973, *A&A*, 24, 337
- Sijacki D., Springel V., 2006, *MNRAS*, 366, 397
- Silk J., Rees M. J., 1998, *A&A*, 331, L1
- Soltan A., 1982, *MNRAS*, 200, 115
- Spergel D. N. et al., 2007, *ApJS*, 170, 377
- Springel V., 2005, *MNRAS*, 364, 1105
- Springel V., Hernquist L., 2002, *MNRAS*, 333, 649
- Springel V., Hernquist L., 2003a, *MNRAS*, 339, 289
- Springel V., Hernquist L., 2003b, *MNRAS*, 339, 312
- Springel V., White S. D. M., Tormen G., Kauffmann G., 2001a, *MNRAS*, 328, 726
- Springel V., Yoshida N., White S. D. M., 2001b, *New Astron.*, 6, 79
- Springel V., Di Matteo T., Hernquist L., 2005a, *ApJ*, 620, L79
- Springel V., Di Matteo T., Hernquist L., 2005b, *MNRAS*, 361, 776
- Steffen A. T., Barger A. J., Cowie L. L., Mushotzky R. F., Yang Y., 2003, *ApJ*, 596, L23
- Thacker R. J., Scannapieco E., Couchman H. M. P., 2006, *ApJ*, 653, 86

- Tormen G., Bouchet F. R., White S. D. M., 1997, MNRAS, 286, 865
Tornatore L., Borgani S., Matteucci F., Recchi S., Tozzi P., 2004, MNRAS, 349, L19
Tremaine S. et al., 2002, ApJ, 574, 740
Ueda Y., Akiyama M., Ohta K., Miyaji T., 2003, ApJ, 598, 886
Vikhlinin A., Markevitch M., Murray S. S., Jones C., Forman W., Van Speybroeck L., 2005, ApJ, 628, 655
Volonteri M., Haardt F., Madau P., 2003, ApJ, 582, 559
von der Linden A., Best P. N., Kauffmann G., White S. D. M., 2007, MNRAS, 379, 867
Yoshida N., Sheth R. K., Diaferio A., 2001, MNRAS, 328, 669
Yu Q., Tremaine S., 2002, MNRAS, 335, 965

This paper has been typeset from a \TeX/L\TeX file prepared by the author.