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# Tracking Volatile Behaviour in Sub-volcanic Plumbing Systems Using Apatite and Glass: Insights into Pre-eruptive Processes at Campi Flegrei, Italy

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#### **ABSTRACT**

Volatile elements play an important role in many aspects of the physicochemical architecture of sub-volcanic plumbing systems, from the liquid line of descent to the dynamics of magma storage and eruption. However, it remains difficult to constrain the behaviour of magmatic volatiles on short timescales before eruption using established petrological techniques (e.g. melt inclusions); specifically, in the final days to months of magma storage. This study presents a detailed model of pre-eruptive volatile behaviour in the Campi Flegrei system (Italy), through combined analyses of apatite crystals and glass. The deposits of eight eruptions were examined, covering the full spectrum of melt compositions, eruptive styles and periods of activity at Campi Flegrei in the past 15 kyr. Measured apatite compositions are compared with thermodynamic models that predict the evolution of the crystal compositions during different fractional crystallization scenarios, including (1) volatile-undersaturated conditions, (2)  $H_2O$ -saturated conditions and (3) varying P-T conditions. The compositions of clinopyroxene-hosted and biotite-hosted apatite inclusions are consistent with crystallization under volatile-undersaturated conditions that persisted until late in magmatic evolution. Apatite microphenocrysts show significantly more compositional diversity, interpreted to reflect a mixed cargo of crystals derived from volatile-undersaturated melts at depth and melts that have undergone cooling and degassing in discrete shallow-crustal magma bodies. Apatite microphenocrysts from lavas show some re-equilibration during cooling at the surface. Clinopyroxene-hosted melt inclusions within the samples typically contain 2-4 wt % H<sub>2</sub>O, indicating that they have been reset during temporary magma storage at 1-3 km depth, similar to the depth of sill emplacement during recent seismic crises at Campi Flegrei. Comparable apatite compositional trends are identified in each explosive eruption analysed, regardless of volume, composition or eruption timing. However, apatites from the different epochs of activity appear to indicate subtle changes in the H<sub>2</sub>O content of the parental melt feeding the Campi Flegrei system over time. This study demonstrates the potential utility of integrated apatite and glass analysis for investigating pre-eruptive volatile behaviour in apatite-bearing magmas.

Key words: apatite; melt inclusions; Campi Flegrei; volatiles; thermodynamic modelling

## INTRODUCTION

Magmatic volatiles (H<sub>2</sub>O, CO<sub>2</sub>, halogens and SO<sub>2</sub>) are minor but important constituents of most silicate melts, and affect almost every aspect of magmatic evolution and eruption. Volatiles influence mineral phase stability and the liquid line of descent (Grove *et al.*, 2003; Zimmer *et al.*, 2010), as well as melt density (Lange & Carmichael, 1990) and viscosity (Giordano *et al.*, 2008), thus exerting a major control on the depths of magma storage and the crustal-scale structure of sub-volcanic systems (Annen *et al.*, 2006). Volatile exsolution and expansion drives volcanic eruptions, and the pre-eruptive behaviour of volatiles plays an important role in controlling the style and tempo of volcanism at the Earth's surface (Roggensack *et al.*, 1997; Huppert & Woods, 2002; Cashman, 2004; Edmonds, 2008).

Given the fundamental role of volatiles in controlling volcanic processes, quantification of their pre-eruptive concentrations remains a high priority for any investigation. A variety of petrological methods have been used to decipher the volatile histories of past eruptions, giving access to distinct snapshots of melt volatile contents over variable pre-eruptive timescales. For example, melt inclusions capture a record of magmatic volatiles at the time of entrapment, and have been used widely to constrain volatile concentrations in different magmatic settings (e.g. Dunbar et al., 1989; Saal et al., 2002; Wallace, 2005; Stefano et al., 2011). However, recent studies have revealed the rapidity of H<sup>+</sup> diffusion in common igneous phenocrysts under magmatic conditions, relative to the timescales of pre-eruptive magma storage (e.g. Woods et al., 2000; Ingrin & Blanchard, 2006; Reubi et al., 2013). Melt inclusion H<sub>2</sub>O contents may diffusively re-equilibrate through their host crystals within hours to weeks, limiting this volatile record to the very final stages of magma storage and/or ascent (e.g. Danyushevsky et al., 2002; Portnyagin et al., 2008; Gaetani et al., 2012; Bucholz et al., 2013; Lloyd et al., 2013; Reubi et al., 2013; Preece et al., 2014). It has also been shown that significant amounts of CO2 may migrate from the inclusion melt into shrinkage bubbles, leading to underestimates of pressure (Moore et al., 2015). Understanding volatile systematics during latestage magma storage and the onset of magma ascent is essential for identifying eruption-triggering processes and understanding the 'warning' signs that would be observed at the Earth's surface in the build-up to an eruption.

Apatite [Ca<sub>4</sub>(PO<sub>4</sub>)<sub>3</sub>F, CI, OH] is a common accessory mineral in volcanic, plutonic and ore-forming environments (Piccoli & Candela, 2002) and has received increasing attention as a potential magmatic volatile 'probe', owing to its ability to incorporate all major magmatic volatiles into its crystal structure. Halogens and OH are essential structural constituents in apatite and are incorporated as part of a series of exchange equilibria (Candela, 1986; McCubbin *et al.*, 2015). Sulphate and CO<sub>3</sub><sup>2-</sup> may also substitute into apatite as

trace components (e.g. Pan & Fleet, 2002; Dietterich & de Silva, 2010; Riker et al., 2018). Recent work has focused on deciphering the relationship between apatite F-CI-OH compositions and their host melt volatile contents (e.g. Candela, 1986; Piccoli & Candela, 1994; Patiño Douce & Roden, 2006; Boyce & Hervig, 2009; McCubbin et al., 2011; Boyce et al., 2014; Stock et al., 2016). One advantage of apatite analysis is that phenocryst-hosted apatite inclusions can preserve a record of melt volatile compositions under conditions in which melt inclusions may have re-equilibrated (Stock et al., 2016). Because volatile re-equilibration in mineral-hosted apatite inclusions requires simultaneous diffusion of F, Cl and/or OH, this process will be rate-limited by halogen diffusivity in the host phenocrysts, which is significantly slower than that of H<sup>+</sup> (Bucholz et al., 2013; Lloyd et al., 2013). In contrast, F-CI-OH diffusion within apatite crystals is relatively rapid and microphenocrysts are therefore able to exchange volatiles with their host liquids on geologically short timescales [i.e. microphenocryst rims may equilibrate in weeks to years at magmatic temperatures (T); Brenan, 1993]. Because these timescales of apatite microphenocryst re-equilibration are longer than timescales of magma ascent (i.e. hours to days), apatite microphenocrysts may preserve a record of preeruptive conditions, even when matrix glasses degas at low pressure [(P); Stock et al., 2016].

In this study, we investigate apatite and glass compositions in juvenile samples from eight eruptions of the Campi Flegrei volcano (Italy), to determine magmatic volatile systematics and processes in the build-up to eruptions. Campi Flegrei was selected as the focus of this study because its melts are known to be apatitebearing and volatile-rich (D'Antonio et al., 1999; Cannatelli et al., 2007; Arienzo et al., 2010, 2016). It has also recently shown signs of unrest (Chiodini et al., 2012; Moretti et al., 2017). Building on the work of Candela (1986) and Piccoli & Candela (1994) we develop thermodynamic models that predict the theoretical compositional evolution of apatite as a function of changing magma compositions during fractional crystallization in the presence or absence of different fluid phases. Different populations of apatite inclusions (hosted in biotite and clinopyroxene) and microphenocrysts are identified based on their volatile compositions and, through comparison with our thermodynamic models, we use these to constrain the pattern of magmatic volatile behaviour in the sub-volcanic plumbing system at Campi Flegrei. Although melt inclusions have re-equilibrated during magma ascent, coupled interpretation of apatite and glass compositions provides additional constraints on the structure of the Campi Flegrei plumbing system and the composition of the magmatic fluid phase prior to eruption. Finally, we discuss apparent variations in apatite volatile contents that are linked to different periods of eruptive activity at Campi Flegrei, and suggest that the volatile contents of the incoming parental magmas vary with time.

#### **GEOLOGICAL SETTING**

Campi Flegrei comprises a nested caldera system, defined by collapse scarps that formed during the Campanian Ignimbrite (~40 ka; Giaccio et al., 2017) and Neapolitan Yellow Tuff (NYT, ~15 ka; Deino et al., 2004) eruptions (Fig. 1). It is one of the most active volcanoes in Europe, having produced >60 eruptions in the past 15 kyr, from vents located within the NYT caldera (Smith et al., 2011). These are divided into three 'epochs' that represent periods of eruptive activity, separated by prolonged quiescence (Di Vito et al., 1999). Vents for these eruptions are located within the NYT caldera (Fig. 1; Di Vito et al., 1999; Isaia et al., 2009). Epoch 1 occurred from ~15 to 10.6 ka and produced  $\sim$ 30 explosive eruptions, with a typical inter-eruptive interval of ~70 years (Di Vito et al., 1999; Smith et al., 2011). Six low-magnitude explosive eruptions occurred in Epoch 2, between  $\sim$ 9.6 and 9.1ka, at an average interval of ~65 years (Di Vito et al., 1999; Smith et al., 2011), followed by a long period ( $\sim$ 4 kyr) of guiescence. Twenty-seven eruptions occurred within the short Epoch 3 time-period, between ~5.5 and 3.5 ka (Smith et al., 2011), with an average eruptive interval of ~75 years (Di Vito *et al.,* 1999). Most Epoch 3 eruptions were small, explosive events [typically producing 0·02–0·10 km³ of material, dense rock equivalent (DRE); Smith *et al.,* 2011]. However, uniquely within the past 15 kyr, Epoch 3 also includes four effusive lava domes (Melluso *et al.,* 1995; Di Vito *et al.,* 1999). The most recent Campi Flegrei eruption was at Monte Nuovo in 1538 CE. This occurred after a >3 kyr dormant period (Piochi *et al.,* 2005), substantially greater than the typical inter-eruption time interval, and is therefore not considered part of Epoch 3 (Smith *et al.,* 2011).

# Chemical diversity of eruptive products in the past 15 kyr

The most mafic (shoshonitic) melt inclusions identified in Campi Flegrei typify the mantle melts feeding the system (Mangiacapra et al., 2008; Vetere et al., 2011). Major and trace element studies of Campi Flegrei whole-rocks and glasses show an evolutionary trend from these mafic melts to evolved trachytes or phonolites, with the entire suite formed by fractional crystallization of a single parental magma composition, punctuated by periodic recharge events (Villemant,

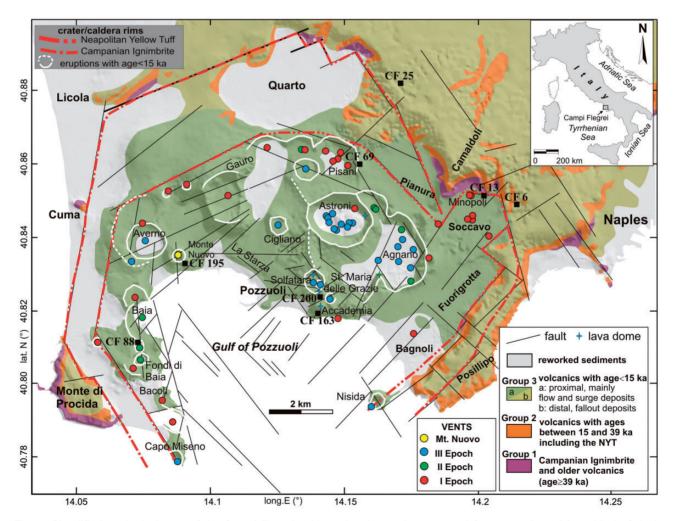


Fig. 1. Simplified geological map of the Campi Flegrei caldera showing major structural features, adapted from Vitale & Isaia (2014). The locations of sampling sites (black squares) and vents for each eruption analysed in this study (Table 1) are indicated.

1988; Civetta *et al.*, 1991; D'Antonio *et al.*, 1999; Pappalardo *et al.*, 1999; Di Renzo *et al.*, 2011; Di Vito *et al.*, 2011; Smith *et al.*, 2011; Fourmentraux *et al.*, 2012). Pre-NYT (older than 15 ka) melts encompass only the most evolved end of the compositional spectrum (Pappalardo *et al.*, 1999). Post-NYT (younger than 15 ka) magmas do not follow a continual compositional trend between eruptions (i.e. from primitive to evolved), but the most primitive matrix glasses and whole-rocks are derived from Epoch 1, with Epochs 2 and 3 largely indistinguishable based on their major element compositions (D'Antonio *et al.*, 1999; Smith *et al.*, 2011). Matrix glasses from Monte Nuovo are unlike previous eruptions, with notably higher Na<sub>2</sub>O concentrations (Smith *et al.*, 2011).

In contrast, isotopic heterogeneity in Campi Flegrei products suggests that the magmas cannot have formed through fractional crystallization alone (Pappalardo et al., 1999, 2002; D'Antonio et al., 2007; Di Renzo et al., 2011). The isotopic compositions of <15 ka Campi Flegrei melts reflect mixing between three distinct end-members, defined by Di Renzo et al. (2011) as the 'NYT component' (87Sr/86Sr 0.70750-0.70753, <sup>143</sup>Nd/<sup>144</sup>Nd  $\sim$ 0.51246,  $^{206}\text{Pb/}^{204}\text{Pb}\sim$  19·04,  $\delta^{11}B\sim$  –7·9%, the 'Minopoli 2 compo-<sup>143</sup>Nd/<sup>144</sup>Nd (<sup>87</sup>Sr/<sup>86</sup>Sr  $\sim$ 0.70860,  $\sim$ 0.51236,  $^{206}\text{Pb/}^{204}\text{Pb} \sim \! 18.90 , \, \delta^{11}\text{B} - \! 7.32\% )\! ,$  and the 'Astroni 6 com- $^{143}$ Nd/ $^{144}$ Nd  $\sim$ 0.51250, ponent' ( $^{87}$ Sr/ $^{86}$ Sr  $\sim 0.70726$ .  $^{.206}$ Pb/ $^{204}$ Pb  $\sim$ 19·08,  $\delta^{11}$ B –9·8 $^{\circ}_{\infty}$ ). The isotopic composition of Campi Flegrei magmas is defined in the deep crust, before significant crystal fractionation, but the major and trace element compositions of these primitive liquids must be closely similar to permit evolution along the same liquid line of descent (Pappalardo et al., 2002). The prevalence of these isotopic components correlates with the different epochs of activity: Epochs 1 and 2 show mixing between 'NYT' and 'Minopoli 2' components; Epoch 3 shows mixing between 'NYT', 'Minopoli 2' and 'Astroni 6' components; the Monte Nuovo eruption sampled near endmember 'Astroni 6 component' melts (Di Renzo et al., 2011). Enriched  $\delta^{11}B$  in Epochs 1 and 2 reflects a more metasomatized mantle source caused by a higher slabderived fluid input (D'Antonio et al., 2007) and/or a lower subducted sediment input into the mantle wedge (Tonarini et al., 2004; Di Renzo et al., 2011). Higher 87Sr/86Sr and lower 143Nd/144Nd and 206Pb/204Pb ratios demonstrate greater crustal assimilation in Epochs 1 and 2 than in Epoch 3 and Monte Nuovo (D'Antonio et al., 2007; Di Renzo et al., 2011). However, the Campi Flegrei liquid line of descent and extent of Sr and Pb isotopic heterogeneity is compatible only with very minor assimilation (D'Antonio et al., 2007; Fowler et al., 2007).

Water contents of Campi Flegrei melt inclusions typically vary from  $\sim$ 1 to  $\sim$ 4 wt % and show no systematic relationship with the degree of magma differentiation (Mangiacapra *et al.*, 2008; Arienzo *et al.*, 2016; Stock *et al.*, 2016). Campi Flegrei melt inclusions generally have very low CO<sub>2</sub> concentrations (<250 ppm), with a few analyses extending up to 400–500 ppm (Marianelli *et al.*, 2006; Arienzo *et al.*, 2010, 2016; Stock *et al.*, 2016).

#### SAMPLES AND METHODS

## **Eruptions studied**

The eruptions investigated cover the full range of melt compositions, eruption sizes and styles of activity from Campi Flegrei in the past 15 kyr (Table 1; for full stratigraphy see Smith et al., 2011). Vent locations, average matrix glass compositions and absolute eruption ages for the samples analysed in this study are given in Fig. 1 and Table 1. The Minopoli 1 tuff cone was sampled because it is a small, phono-tephritic to tephri-phonolitic explosive eruption, typical of the early Epoch 1 eruptions that followed the NYT event (Smith et al., 2011). Scoria was also sampled from the slightly larger, slightly more evolved Pisani 1 eruption, which occurred in mid-late Epoch 1 (Smith et al., 2011). Deposits from more recent eruptions are typically more evolved. We sampled Astroni 1 from Epoch 3, which represents the first of seven small explosive eruptions from the Astroni vent between 4 and 4.4 ka (Isaia et al., 2004; Smith et al., 2011). These deposits comprise phreatomagmatic surge beds interbedded with subordinate Strombolian pumice layers (Smith et al., 2011). Pomici Principali (PP) was the largest eruption in the last 15 kyr, generating a Plinian column and pyroclastic density currents (Di Vito et al., 1999; Smith et al., 2011; Bevilacqua et al., 2016). These were sampled to investigate relationships between eruption magnitude and apatite volatile compositions. The Baia-Fondi di Baia (B-FdB) and Monte Nuovo eruptions were sampled because they produced highly evolved melts from vents in the western NYT caldera. Baia-Fondi di Baia was the first eruption of Epoch 2 and was particularly explosive, owing to magma-H<sub>2</sub>O phreatomagmatic interaction, but it expelled only a relatively small volume of material (Pistolesi et al., 2017). The most evolved melts identified in Campi Flegrei were produced in the historical Monte Nuovo tuff cone eruption (Smith et al., 2011). The latitic Santa Maria delle Grazie (SMdG) lava, which is thought to be part of a shallow dyke that fed the SMdG scoria cone (Isaia et al., 2009), and the subsequent trachytic Accademia lava dome (Isaia et al., 2009; Melluso et al., 2012) were sampled to assess differences between apatite volatile compositions in these deposits and explosive units.

## **Samples**

All samples were collected from proximal deposits (Fig. 1), either by Smith *et al.* (2011) or during fieldwork for this study in March 2013 and September 2014. Minopoli 1, PP, Pisani 1, B–FdB and Astroni 1 samples are CF13, CF6, CF25, CF88 and CF69 of Smith *et al.* (2011). The Astroni 1 sample was investigated by Stock *et al.* (2016). CF88 is from the initial fallout (Baia) phase of the B–FdB eruption (Pistolesi *et al.*, 2017). The Monte Nuovo sample (CF195) is from the upper pyroclastic flow unit (Table 1; Unit II of Piochi *et al.*, 2005). The SMdG (CF200) and Accademia (CF163) samples were collected from the centre of the NYT caldera.

Table 1: Summary of eruption deposits analysed, including age estimate, deposit characteristics (i.e. eruption style), erupted volume (i.e. magnitude estimate), average matrix glass composition, vent location and sampling location

Epoch	Eruption	Sample number	Age (cal. years BP)	Deposit type	Erupted volume (km³)	Average matrix glass composition	Vent location*		Sampling location	
					(KIII )		UTM X	UTM Y	UTM X	UTM Y
	Monte Nuovo	CF195	1538 CE*	Pyroclastic density current deposit	~0.03*	Phonolite	423095	4520870	423277	4520614
Epoch 3	Astroni 1	CF69	4155–4357*	Ash fall with some lapilli layers	~0.06*	Phonolite	427999	4522122	428859*	4523584*
	Accademia	CF163	4205-4507*1	Lava dome	<0.01*	Trachyte <sup>§</sup>	427551	4519282	427597	4519562
	Santa Maria delle Grazie	CF200	4383–4507*	Shallow feeder dyke	<0.01*	Latite <sup>II</sup>	427625	4519806	427421	4518909
Epoch 2	Baia–Fondi di Baia	CF88	9525–9705*	Pumice fall	0·02–0·04 <sup>†</sup>	Trachyte	421855	4517930	421844*	4518253*
Epoch 1	Pisani 1	CF25	10516–12107* <sup>2</sup>	Scoria fall	0.1–0.3*	Tephri-phonolite- trachy-andesite	428260	4523970	430237*	4525994*
	Pomici Principali	CF6	11907–12107*	Plinian fall (B3 unit)	0·43–1·28 <sup>‡</sup>	Tephri-phonolite-	428698	4522284	433890*	4522346*
	Minopoli 1	CF13	11907–12725* <sup>3</sup>	Scoria fall	<0.1*	Phono-tephrite- tephri-phonolite	432400	4522670	432805*	4522632*

Erupted volumes are given as dense rock equivalent. Where exposure permitted, samples were collected vertically through volcaniclastic units and included both lapilli- and ash-sized material. The Pomici Principali sample represents only the B3 Plinian phase of the eruption and the Baia–Fondi di Baia sample comprises only the basal Baia fallout units. Samples from the deposits of these eruptions show only limited diversity in glass compositions (Smith *et al.*, 2011) and the analysed samples are therefore assumed representative of the bulk erupted material.

All samples have the major phase assemblage Kfeldspar + plagioclase + clinopyroxene + biotite + apatite + magnetite. Samples also contain fluorite ± sulphides ± sodalite ± leucite, with precipitation of these accessory phases constrained to late in magmatic evolution by their absence as phenocryst-hosted inclusions in natural samples, and by experimental studies and thermodynamic models (Fowler et al., 2007; Arzilli et al., 2016). Olivine is reported in mafic samples from Campi Flegrei (Cannatelli et al., 2007) but was not observed in this study. Crystal contents are typically <5–30% (from qualitative observations and Isaia et al., 2004; Piochi et al., 2005; Mastrolorenzo & Pappalardo, 2006) but are notably lower in B-FdB ( $\ll$ 1%; Mastrolorenzo & Pappalardo, 2006). In Monte Nuovo, two pyroxene populations can be identified in hand specimen: one black and one green, as in other eruptions at Campi Flegrei and Vesuvius (Cioni et al., 1998; D'Antonio et al., 1999). As apatite and melt inclusions show no systematic compositional difference between these pyroxene populations, they are not separated in the following discussion.

## **Analytical methods**

Clinopyroxene and biotite phenocrysts were hand-picked from the  $250-500\,\mu m$  size fraction in samples from explosive eruptions and the Accademia lava

dome. Heavy liquid and magnetic separation techniques were used to extract apatite microphenocrysts from the 44-250 µm size fraction. Crystals and matrix ash grains were mounted in epoxy, ground and polished for analysis. Lava samples were prepared as polished thin sections. Samples were examined using an FEI Quanta 650 FEG scanning electron microscope (SEM), operating with a 20 kV,  $\sim$ 6-7 nA beam, in the Department of Earth Sciences, University of Oxford. Only apatite inclusions away from cracks and fully enclosed within host phenocrysts were analysed, to ensure that they were trapped during phenocryst growth and were unable to subsequently re-equilibrate with melt or fluids. Melt inclusions were analysed only if they did not show visual evidence for post-entrapment crystallization and were located away from cracks in their host phenocryst. Ash, lapilli fragments and lavas were also assessed by SEM to identify microlite-free regions for analysis of the matrix glass compositions.

Mineral and glass major, trace and halogen element compositions were analysed using a JEOL 8600 electron microprobe at the Research Laboratory for Archaeology and the History of Art, University of Oxford. Samples were re-polished prior to electron probe microanalysis (EPMA) to remove any compositional modification induced by SEM electron-beam exposure (Stock et al., 2015), and subsequently carbon

<sup>\*</sup>Data from Smith et al. (2011). All ages of Smith et al. (2011) have been recalibrated using IntCal13 (Reimer et al., 2013).

<sup>&</sup>lt;sup>†</sup>Data from Pistolesi et al. (2017).

<sup>&</sup>lt;sup>‡</sup>Data from Bevilacqua et al. (2016).

<sup>&</sup>lt;sup>§</sup>Whole-rock composition from Melluso et al. (2012).

Whole-rock composition from Isaia et al. (2009).

<sup>&</sup>lt;sup>1</sup>Eruption not directly dated but stratigraphically older than the Solfatara eruption (4205–4427 years BP) and younger than Santa Maria delle Grazie.

<sup>&</sup>lt;sup>2</sup>Eruption not directly dated but stratigraphically older than Pisani 3 (10 517–10 760 years BP) and younger than Pomici Principali.

<sup>&</sup>lt;sup>3</sup>Eruption not directly dated but stratigraphically older than Pomici Principali and younger than Archiaverno (12579–12725 years BP).

coated along with secondary standards to avoid variable light element X-ray attenuation. Apatite was analysed using a defocused (5 µm), 15 kV, 10 nA beam, with halogens analysed first. Where possible, apatite crystals were analysed with the c-axis parallel to the plane of the mount. This routine limits the potential for timedependent variability in halogen X-ray counts during analysis (Stormer et al., 1993; Goldoff et al., 2012; Stock et al., 2015), while maintaining reasonable precision for low-concentration elements (i.e. CI). In glass, most elements were measured using a defocused (10 μm), 15 kV, 6 nA beam to minimize Na<sub>2</sub>O and SiO<sub>2</sub> migration (e.g. Humphreys et al., 2006); when in low abundance, SO<sub>2</sub>, P<sub>2</sub>O<sub>5</sub> and CI were measured in a second analysis using a 30 nA current. In both crystals and glass, count times were 20-30 s for major elements and 30-90 s for minor elements (120s for CI and SO<sub>2</sub> in apatite). Backgrounds were determined by counting for half of the on-peak count time on either side of the peak. Glass and apatite analytical totals were typically 95-99% and 96-102%, respectively. Data were filtered to remove analyses with totals <92%, and >101.5% in glass. Apatite totals >100% probably reflect minor electronbeam induced compositional modification (Stock et al., 2015) and totals significantly <100% probably result from the absence of trace elements [e.g. rare earth elements (REEs)] in the analytical routine. Apatite OH contents in EPMA data were calculated 'by difference', assuming stoichiometry. Typical analytical uncertainties are reported in Tables 2 and 3.

A subset of apatites and glasses was analysed by secondary ion mass spectrometry (SIMS) for H, F and CI using a Cameca ims-4f system at the Edinburgh Ion Micro-Probe Facility, University of Edinburgh, using the methods outlined by Stock et al. (2016). Samples were re-polished prior to SIMS analysis to remove any surficial halogen modification induced during EPMA (Stock et al., 2015). In apatite, F, CI and H<sub>2</sub>O concentrations were derived from working curves of <sup>19</sup>F/<sup>44</sup>Ca vs F,  $^{35}$ Cl/ $^{44}$ Ca vs Cl and  $^{1}$ H/ $^{44}$ Ca vs H<sub>2</sub>O, populated by independently characterized apatite (Supplementary Data Fig. S1; supplementary data are available for downloading at http://www.petrology. oxfordjournals.org). Glass H<sub>2</sub>O concentrations were similarly derived from working curves of <sup>1</sup>H/<sup>30</sup>Si vs H<sub>2</sub>O populated by well-characterized rhyolitic glasses (Supplementary Data Fig. S2). Working curves were created at the beginning of each day, analytical session or when beam conditions were changed. NIST SRM610 was used as a primary standard for glass F calibration, with 44Ca as the internal standard, based on EPMA of the same spot. SIMS backgrounds were monitored using anhydrous synthetic apatite or clinopyroxene crystals for apatite and glass, respectively. Working curves have a polynomial fit for H<sub>2</sub>O in apatite and linear fit for all other elements, with  $R^2$  typically >0.99. Typical analytical uncertainties are reported in Tables 2 and 3.

#### **APATITE AND GLASS COMPOSITIONS**

In total, >800 apatite and >250 glass analyses were acquired on samples from the eight Campi Flegrei eruptions (Table 1), using EPMA to identify broad compositional trends (full datasets are given in Supplementary Data Tables S1 and S2) and SIMS to measure a subset of crystals and glasses for H2O and halogens with lower analytical uncertainties (Tables 2 and 3). All glass data in the text and Figs 2, 3 and 4a-d are plotted normalized to 100% anhydrous. All apatite analyses were made close to the centre of the mineral grains. Glass compositions could not be obtained from lava deposits (Table 1), owing to melt inclusion devitrification and groundmass crystallization. Apatite inclusions from these eruptions were excluded as they were not isolated from melt or fluids after entrapment: biotites are partially (in the case of SMdG) or fully (for Accademia) broken down and pyroxenes are heavily fractured, such that all apatite inclusions are touching major cracks or are incompletely enclosed by their host crystals. Baia-Fondi di Baia melt inclusions were too small (typically <20 μm) to analyse by SIMS. Astroni 1 apatite data from Stock et al. (2016) were recalibrated using updated standard compositions, resulting in minor modification of absolute compositions but no change to the trends in this dataset. These were included with additional glass and apatite analyses from this eruption.

## Glass major element compositions

Melt inclusions and matrix glasses analysed in this study cover a wide compositional range from trachybasalts and basanites to phonolites and trachytes (Table 2; Supplementary Data Table S1), encompassing almost the entire compositional diversity reported in <15 ka Campi Flegrei deposits (Figs 2 and 3). Clinopyroxene-hosted melt inclusion compositions differ between eruptions: Minopoli 1 melt inclusions are the most mafic (50·46-56·84 wt % SiO<sub>2</sub>) and cover the largest compositional range; Pisani 1 and PP include the next most mafic melt inclusions, with SiO2 as low as 54-18 wt %; and melt inclusions from other eruptions are typically phonolitic to trachytic (Fig. 2). The same overall trend is seen for the matrix glasses (Fig. 3), although, in general, they are more evolved than melt inclusions from the same eruption, with higher SiO<sub>2</sub> (typically >53 wt %) and alkali element concentrations (Table 2; Supplementary Data Table S1).

For all elements measured, melt inclusion and matrix glass compositions plot on a single compositional trend, in agreement with literature data (Figs 2 and 3). Glass SiO<sub>2</sub>, Na<sub>2</sub>O and K<sub>2</sub>O concentrations typically increase with decreasing MgO, whereas CaO, FeOt (all Fe as FeO) and TiO<sub>2</sub> concentrations decrease (Figs 2 and 3; Table 2; Supplementary Data Table S1). However, glass K<sub>2</sub>O and SiO<sub>2</sub> concentrations decrease at low MgO contents ( $\lesssim$ 0.5 wt % MgO). Some low-MgO matrix glasses from Monte Nuovo and B–FdB are particularly K<sub>2</sub>O-

Table 2: Major and volatile element analyses of representative Campi Flegrei glasses

Analysis #:	Monte Nuovo CF195_cp x113_m6	Monte Nuovo CF195_cp x64_m1	Monte Nuovo CF195_cp x86_m2	Monte Nuovo CF195_cp x86_m3	Monte Nuovo CF195_cp x110_m1	Monte Nuovo CF195_cp x113_m7	Monte Nuovo CF195_cp x122_m1	Monte Nuovo CF195_cp x122_m2	Monte Nuovo CF195_cp x123_m3	Monte Nuovo CF195_cp x137_m5
Type:	mi	mi	mi	mi	mi	mi	mi	mi	mi	mi
SiO <sub>2</sub>	58.77	57.35	57.81	56-19	58.05	57.63	58-35	58-26	59.68	58.71
Na <sub>2</sub> O	4.45	4.04	4.11	3.94	5.04	4.65	4.20	4.08	5.03	4.42
MgO Al <sub>2</sub> O <sub>3</sub>	0⋅55 18⋅17	0.77 17.70	0⋅75 18⋅27	0·74 17·26	0⋅76 16⋅54	0·59 18·02	0.87 18.03	0⋅81 18⋅12	0⋅51 16⋅97	0·70 17·71
K <sub>2</sub> O	8.40	8.57	8.49	7.89	6.86	8.37	8.48	8.94	7.43	8.42
CaO	2.22	2.59	2.55	2.48	2.72	2.26	2.81	2.83	2.33	2.24
TiO <sub>2</sub>	0.43	0.48	0.48	0.42	0.55	0.46	0.50	0.44	0.53	0.49
MnO	0.09	0.09	0.11	0⋅15	0.19	0.08	0.10	0.12	0.12	0.11
FeOt	3.31	3.82	3.46	3.39	3.96	3.38	3.83	3.66	3.85	3.37
P <sub>2</sub> O <sub>5</sub> CI	0·08 0·95	0·15 0·66	0·13 0·73	0·14 0·64	0·08 0·85	0·12 0·84	0·13 0·67	0·13 0·66	0·10 0·94	0·11 0·77
F	0.25	0.22	0.23	0.22	0.22	0.28	0.25	0.22	0.23	0.28
H₂O	2.82	3.48	2.84	5.56	3.47	2.45	2.44	1.04	2.28	2.54
Total	100-47	99.93	99.96	99-01	99-29	99-12	100-67	99-31	100-01	99.87
Eruption:	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Astroni 1	Astroni 1	Astroni 1
A 1	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	0500	0500	0500
Analysis #:	CF195_cp	CF195_cp	VF55_cf19 5_1_1	VF55_cf19 5_2_1	VF55_cf19 5_3_1	VF55_cf19	VF55_cf19	CF69_cp	CF69_cp	CF69_cp x55_m1
Type:	x159_m4 mi	x188_m1 mi	b_I_I mg	5_∠_1 mg	5_3_1 mg	5_4_1 mg	5_4_2 mg	x22_m1 mi	x23_m2 mi	x55_m1 mi
SiO <sub>2</sub>	57·83	58.59	56.50	57·19	58.62	57·66	57·48	57·70	59·17	56.74
Na₂O MgO	4⋅85 0⋅66	4.77 0.68	6⋅88 0⋅21	7⋅88 0⋅25	7.56 0.28	7.69 0.17	7·29 0·18	4⋅39 0⋅65	4·25 0·42	4·26 0·54
Al <sub>2</sub> O <sub>3</sub>	17·10	17.52	18-28	18.79	19.04	18-69	18.87	17.90	17·84	18.07
Γ <sub>2</sub> Ο <sub>3</sub> Κ <sub>2</sub> Ο	7.18	8-25	6.88	7.23	7.25	7.20	7.30	8.08	8.36	8.89
CaO	2.59	2.52	1.81	1.51	1.66	1.69	1.75	2.47	2.19	2.17
TiO <sub>2</sub>	0.50	0.47	0.49	0.50	0.39	0.45	0.43	0.45	0.49	0.41
MnO	0.20	0.13	0.24	0.37	0.20	0.20	0.28	0.10	0.19	0.04
FeOt	3.91 0.15	3⋅38 0⋅11	3·08 0·01	3·30 0·00	2·74 0·00	2·89 0·05	3·19 0·07	3⋅36 0⋅12	3⋅65 0⋅10	3·15 0·08
P <sub>2</sub> O <sub>5</sub> Cl	0.15	0.11	0.90	0.00	0.00	1.05	1.17	0.12	1.00	0.08
F	0.24	0.32	1.21	0.65	0.73	0.82	0.68	0.28	0.28	0.26
H <sub>2</sub> O	3.71	2.38	0.88	0.76	0.97	1.06	1.08	2.37	2.22	2.12
Total	99.80	99.88	97.36	99-41	100-41	99-60	99.77	98-82	100-15	97.60
Eruption:	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1
Analysis #:	CF69_cpx	CF69_cpx	CF69_cpx	CF69_cpx	CF69_cpx	CF69_cpx	CF69_cpx	CF69_cpx	CF69_cpx	CF69_cpx
	55_m2	66_m1	66_m3 mi	82_m1 mi	91_m1 mi	92_m1 mi	97_m3 mi	113_m1 mi	116_m1 mi	127_m1 mi
Type	mi		1111				1111	1111		11111
Type:	mi	mi								
SiO <sub>2</sub>	56-99	54.97	56.56	56.78	58-21	57.38	57.73	56.74	56-44	56.55
SiO <sub>2</sub> Na <sub>2</sub> O	56·99 4·52	54·97 4·15	4.02	56·78 4·20	58·21 4·60	57·38 4·35	4.46	4.76	56·44 3·65	4.22
SiO <sub>2</sub> Na <sub>2</sub> O MgO	56·99 4·52 0·53	54·97 4·15 0·61	4·02 0·72	56·78 4·20 0·50	58·21 4·60 0·51	57·38 4·35 0·76	4·46 0·54	4·76 0·55	56·44 3·65 0·78	4·22 0·58
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub>	56·99 4·52	54·97 4·15	4.02	56·78 4·20	58·21 4·60	57·38 4·35	4.46	4.76	56·44 3·65	4.22
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO	56·99 4·52 0·53 17·82	54·97 4·15 0·61 17·60	4·02 0·72 18·24	56·78 4·20 0·50 17·90	58·21 4·60 0·51 17·67 8·37 2·03	57·38 4·35 0·76 18·60	4·46 0·54 17·91	4·76 0·55 17·51	56·44 3·65 0·78 18·40	4·22 0·58 17·65
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub>	56-99 4-52 0-53 17-82 8-50 2-15 0-44	54·97 4·15 0·61 17·60 8·55 2·43 0·47	4·02 0·72 18·24 8·83 2·62 0·47	56·78 4·20 0·50 17·90 8·35 2·07 0·45	58·21 4·60 0·51 17·67 8·37 2·03 0·48	57·38 4·35 0·76 18·60 8·44 2·47 0·41	4·46 0·54 17·91 8·27 2·34 0·45	4·76 0·55 17·51 8·57 2·17 0·45	56·44 3·65 0·78 18·40 8·84 2·37 0·51	4·22 0·58 17·65 8·28 2·34 0·45
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub>	56.99 4.52 0.53 17.82 8.50 2.15 0.44 0.13	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09	4·02 0·72 18·24 8·83 2·62 0·47 0·12	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20	57·38 4·35 0·76 18·60 8·44 2·47 0·41 0·11	4·46 0·54 17·91 8·27 2·34 0·45 0·12	4·76 0·55 17·51 8·57 2·17 0·45 0·15	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02	4·22 0·58 17·65 8·28 2·34 0·45 0·10
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt	56.99 4.52 0.53 17.82 8.50 2.15 0.44 0.13 3.37	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40	57·38 4·35 0·76 18·60 8·44 2·47 0·41 0·11 2·85	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub>	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40 0·13	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·08	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub>	56.99 4.52 0.53 17.82 8.50 2.15 0.44 0.13 3.37	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40	57·38 4·35 0·76 18·60 8·44 2·47 0·41 0·11 2·85	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13 0·69 0·23 2·58	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40 0·13 0·91 0·28 2·29	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·08 0·90 0·25 2·00	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51
$SiO_2$ $Na_2O$ $MgO$ $Al_2O_3$ $K_2O$ $CaO$ $TiO_2$ $MnO$ $FeOt$ $P_2O_5$ $CI$ $F$ $H_2O$	56.99 4.52 0.53 17.82 8.50 2.15 0.44 0.13 3.37 0.10 0.93 0.25	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13 0·69 0·23	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40 0·13 0·91 0·28	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31	57·38 4·35 0·76 18·60 8·44 2·47 0·41 0·11 2·85 0·17 0·84 0·30	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·08 0·90 0·25	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F H <sub>2</sub> O Total Eruption:	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28 Astroni 1	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13 0·69 0·23 2·58 99·41	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90 Astroni 1	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·08 0·90 0·25 2·00 98·43	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70
$SiO_{2}$ $Na_{2}O$ $MgO$ $Al_{2}O_{3}$ $K_{2}O$ $CaO$ $TiO_{2}$ $MnO$ $FeOt$ $P_{2}O_{5}$ $Cl$	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13 0·69 0·23 2·58 99·41 Astroni 1 CF69_cpx	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16 Astroni 1 CF69_cpx	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·08 0·90 0·25 2·00 98·43 Astroni 1 VF55_cf69	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71 Astroni 1 VF55_cf69	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf69
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> Cl F H <sub>2</sub> O Total  Eruption: Analysis #:	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28 Astroni 1	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13 0·69 0·23 2·58 99·41	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90 Astroni 1	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·98 0·90 0·25 2·00 98·43 Astroni 1 VF55_cf69 _2_1	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70
$SiO_2$ $Na_2O$ $MgO$ $Al_2O_3$ $K_2O$ $CaO$ $TiO_2$ $MnO$ $FeOt$ $P_2O_5$ $CI$ $F$ $H_2O$ $Total$ $Eruption:$ $Analysis \#:$	56.99 4.52 0.53 17.82 8.50 2.15 0.44 0.13 3.37 0.10 0.93 0.25 2.56 98.28 Astroni 1 CF69_cpx 127_m3 mi	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13 0·69 0·23 2·58 99·41 Astroni 1 CF69_cpx 138_m2 mi	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40 0·13 0·91 0·28 2·29 97·46 Astroni 1 CF69_cpx 156_m2 mi	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16 Astroni 1 CF69_cpx 182_m1 mi	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·08 0·90 0·25 2·00 98·43 Astroni 1 VF55_cf69 _2_1 mg	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71 Astroni 1 VF55_cf69 _5_1 mg	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf65 _6_1 mg
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F H <sub>2</sub> O Total Eruption: Analysis #: Type:	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1	4·02 0·72 18·24 8·83 2·62 0·47 0·12 4·20 0·13 0·69 0·23 2·58 99·41 Astroni 1 CF69_cpx 138_m2	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40 0·13 0·91 0·28 2·29 97·46  Astroni 1 CF69_cpx 156_m2	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16 Astroni 1 CF69_cpx 182_m1	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2	4·46 0·54 17·91 8·27 2·34 0·45 0·12 3·38 0·98 0·90 0·25 2·00 98·43 Astroni 1 VF55_cf69 _2_1	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf68 _6_1
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> Cl F H <sub>2</sub> O Total Eruption: Analysis #: Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO	56.99 4.52 0.53 17.82 8.50 2.15 0.44 0.13 3.37 0.10 0.93 0.25 2.56 98.28  Astroni 1 CF69_cpx 127_m3 mi  57.17 4.07 0.54	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55	4.02 0.72 18.24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99.41 Astroni 1 CF69_cpx 138_m2 mi 55.39 4.62 0.42	56·78 4·20 0·50 17·90 8·35 2·07 0·45 0·19 3·40 0·13 0·91 0·28 2·29 97·46  Astroni 1 CF69_cpx 156_m2 mi 58·52 4·56 0·45	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16 Astroni 1 CF69_cpx 182_m1 mi 55·35 3·76 0·71	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi 55-59 3-87 0-77	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.08 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf65 _6_1 mg 58·83 4·50 0·58
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> Cl F H <sub>2</sub> O Total Eruption: Analysis #: Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub>	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi  57-17 4-07 0-54 17-63	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi	4.02 0.72 18.24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99.41 Astroni 1 CF69_cpx 138_m2 mi 55.39 4.62 0.42 17.09	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi 58-52 4-56 0-45 17-12	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16  Astroni 1 CF69_cpx 182_m1 mi 55·35 3·76 0·71 17·52	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi  55-59 3-87 0-77 17-53	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.08 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf65 _6_1 mg 58·83 4·50 0·58 18·11
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F H <sub>2</sub> O Total Eruption: Analysis #:  Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi  57-17 4-07 0-54 17-63 8-36	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55 17.67 8.37	4.02 0.72 18.24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99.41 Astroni 1 CF69_cpx 138_m2 mi 55.39 4.62 0.42 17.09 8.48	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi 58-52 4-56 0-45 17-12 8-03	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16  Astroni 1 CF69_cpx 182_m1 mi 55·35 3·76 0·71 17·52 8·53	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi 55-59 3-87 0-77 17-53 8-48	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.08 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37 8.61	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85 8·47	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09 8·44	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf65 6_1 mg 58·83 4·50 0·58 18·11 8·56
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F H <sub>2</sub> O Total Eruption: Analysis #:  Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi  57-17 4-07 0-54 17-63 8-36 2-31	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55 17.67 8.37 2.30	4.02 0.72 18.24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99.41 Astroni 1 CF69_cpx 138_m2 mi 55.39 4.62 0.42 17.09 8.48 2.01	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi 58-52 4-56 0-45 17-12 8-03 2-23	58-21 4-60 0-51 17-67 8-37 2-03 0-48 0-20 3-40 0-08 1-03 0-31 2-27 99-16 Astroni 1 CF69_cpx 182_m1 mi 55-35 3-76 0-71 17-52 8-53 2-64	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi  55-59 3-87 0-77 17-53 8-48 2-70	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.08 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37 8.61 2.29	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85 8·47 2·28	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09 8·44 2·45	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf68 -6_1 mg 58·83 4·50 0·58 18·11 8·56 2·36
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> Cl F H <sub>2</sub> O Total Eruption: Analysis #: Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub>	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi 57-17 4-07 0-54 17-63 8-36 2-31 0-49	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55 17.67 8.37 2.30 0.51	4.02 0.72 18.24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99.41 Astroni 1 CF69_cpx 138_m2 mi 55.39 4.62 0.42 17.09 8.48 2.01 0.40	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi 58-52 4-56 0-45 17-12 8-03 2-23 0-48	58·21 4·60 0·51 17·67 8·37 2·03 0·48 0·20 3·40 0·08 1·03 0·31 2·27 99·16  Astroni 1 CF69_cpx 182_m1 mi 55·35 3·76 0·71 17·52 8·53 2·64 0·48	57.38 4.35 0.76 18.60 8.44 2.47 0.41 0.11 2.85 0.17 0.84 0.30 2.21 98.90  Astroni 1 CF69_cpx 182_m2 mi 55.59 3.87 0.77 17.53 8.48 2.70 0.50	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.08 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37 8.61 2.29 0.48	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85 8·47 2·28 0·42	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09 8·44 2·45 0·46	4-22 0-58 17-65 8-28 2-34 0-45 0-10 3-84 0-15 0-82 0-20 2-51 97-70 Astroni 1 VF55_cf68 _6_1 mg 58-83 4-50 0-58 18-11 8-56 2-36 0-37
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> Cl F H <sub>2</sub> O Total Eruption: Analysis #:  Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi  57-17 4-07 0-54 17-63 8-36 2-31	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55 17.67 8.37 2.30	4.02 0.72 18.24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99.41 Astroni 1 CF69_cpx 138_m2 mi 55.39 4.62 0.42 17.09 8.48 2.01	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi 58-52 4-56 0-45 17-12 8-03 2-23	58-21 4-60 0-51 17-67 8-37 2-03 0-48 0-20 3-40 0-08 1-03 0-31 2-27 99-16 Astroni 1 CF69_cpx 182_m1 mi 55-35 3-76 0-71 17-52 8-53 2-64	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi  55-59 3-87 0-77 17-53 8-48 2-70	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.08 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37 8.61 2.29	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85 8·47 2·28	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09 8·44 2·45	4·22 0·58 17·65 8·28 2·34 0·45 0·10 3·84 0·15 0·82 0·20 2·51 97·70 Astroni 1 VF55_cf68 -6_1 mg 58·83 4·50 0·58 18·11 8·56 2·36
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> Cl F H <sub>2</sub> O Total Eruption: Analysis #:  Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi  57-17 4-07 0-54 17-63 8-36 2-31 0-49 0-14 3-60 0-16	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55 17.67 8.37 2.30 0.51 0.11 3.40 0.10	4.02 0.72 18:24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99:41 Astroni 1 CF69_cpx 138_m2 mi 55:39 4.62 0.42 17:09 8.48 2.01 0.40 0.13 3.34 0.06	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi  58-52 4-56 0-45 17-12 8-03 2-23 0-48 0-16 3-53 0-09	58-21 4-60 0-51 17-67 8-37 2-03 0-48 0-20 3-40 0-08 1-03 0-31 2-27 99-16 Astroni 1 CF69_cpx 182_m1 mi 55-35 3-76 0-71 17-52 8-53 2-64 0-48 0-12 3-90 0-20	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi  55-59 3-87 0-77 17-53 8-48 2-70 0-50 0-17 3-39 0-12	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37 8.61 2.29 0.48 0.16 3.27 0.04	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85 8·47 2·28 0·42 0·17 3·31 0·09	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09 8·44 2·45 0·46 0·16 3·42 0·09	4-22 0-58 17-65 8-28 2-34 0-45 0-10 3-84 0-15 0-82 0-20 2-51 97-70 Astroni 1 VF55_cf68 6_1 mg 58-83 4-50 0-58 18-11 8-56 2-36 0-37 0-07 3-41 0-05
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F H <sub>2</sub> O Total Eruption: Analysis #:  Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F H <sub>2</sub> O Total Cruption: Analysis #:	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi  57-17 4-07 0-54 17-63 8-36 2-31 0-49 0-14 3-60 0-16 0-88	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55 17.67 8.37 2.30 0.51 0.11 3.40 0.93	4.02 0.72 18.24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99.41 Astroni 1 CF69_cpx 138_m2 mi 55.39 4.62 0.42 17.09 8.48 2.01 0.40 0.13	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi  58-52 4-56 0-45 17-12 8-03 2-23 0-48 0-16 3-53 0-09 0-91	58-21 4-60 0-51 17-67 8-37 2-03 0-48 0-20 3-40 0-08 1-03 0-31 2-27 99-16  Astroni 1 CF69_cpx 182_m1 mi  55-35 3-76 0-71 17-52 8-53 2-64 0-48 0-12 3-90 0-20 0-74	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi  55-59 3-87 0-77 17-53 8-48 2-70 0-50 0-17 3-39 0-12 0-73	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.08 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37 8.61 2.29 0.48 0.16 3.27 0.04 0.91	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85 8·47 2·28 0·42 0·17 3·31 0·09 0·91	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09 8·44 2·45 0·46 0·16 3·42 0·09 0·94	4-22 0-58 17-65 8-28 2-34 0-45 0-10 3-84 0-15 0-20 2-51 97-70 Astroni 1 VF55_cf65 6-1 mg 58-83 4-50 0-58 18-11 8-56 2-36 0-37 0-07 3-41 0-05 0-92
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub> CI F H <sub>2</sub> O Total Eruption: Analysis #:  Type: SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub>	56-99 4-52 0-53 17-82 8-50 2-15 0-44 0-13 3-37 0-10 0-93 0-25 2-56 98-28  Astroni 1 CF69_cpx 127_m3 mi  57-17 4-07 0-54 17-63 8-36 2-31 0-49 0-14 3-60 0-16	54.97 4.15 0.61 17.60 8.55 2.43 0.47 0.09 3.76 0.14 0.75 0.21 2.21 95.94 Astroni 1 CF69_cpx 130_m1 mi 56.80 4.51 0.55 17.67 8.37 2.30 0.51 0.11 3.40 0.10	4.02 0.72 18:24 8.83 2.62 0.47 0.12 4.20 0.13 0.69 0.23 2.58 99:41 Astroni 1 CF69_cpx 138_m2 mi 55:39 4.62 0.42 17:09 8.48 2.01 0.40 0.13 3.34 0.06	56-78 4-20 0-50 17-90 8-35 2-07 0-45 0-19 3-40 0-13 0-91 0-28 2-29 97-46  Astroni 1 CF69_cpx 156_m2 mi  58-52 4-56 0-45 17-12 8-03 2-23 0-48 0-16 3-53 0-09	58-21 4-60 0-51 17-67 8-37 2-03 0-48 0-20 3-40 0-08 1-03 0-31 2-27 99-16 Astroni 1 CF69_cpx 182_m1 mi 55-35 3-76 0-71 17-52 8-53 2-64 0-48 0-12 3-90 0-20	57-38 4-35 0-76 18-60 8-44 2-47 0-41 0-11 2-85 0-17 0-84 0-30 2-21 98-90  Astroni 1 CF69_cpx 182_m2 mi  55-59 3-87 0-77 17-53 8-48 2-70 0-50 0-17 3-39 0-12	4.46 0.54 17.91 8.27 2.34 0.45 0.12 3.38 0.90 0.25 2.00 98.43 Astroni 1 VF55_cf69 _2_1 mg 59.14 3.90 0.54 18.37 8.61 2.29 0.48 0.16 3.27 0.04	4·76 0·55 17·51 8·57 2·17 0·45 0·15 3·19 0·08 0·93 0·25 2·62 97·97 Astroni 1 VF55_cf69 _3_1 mg 57·69 4·38 0·50 17·85 8·47 2·28 0·42 0·17 3·31 0·09	56·44 3·65 0·78 18·40 8·84 2·37 0·51 0·02 3·27 0·15 0·73 0·22 2·33 97·71  Astroni 1 VF55_cf69 _5_1 mg  59·43 4·67 0·49 18·09 8·44 2·45 0·46 0·16 3·42 0·09	4-22 0-58 17-65 8-28 2-34 0-45 0-10 3-84 0-15 0-82 0-20 2-51 97-70 Astroni 1 VF55_cf68 6_1 mg 58-83 4-50 0-58 18-11 8-56 2-36 0-37 0-07 3-41 0-05

Eruption: Analysis #:	Astroni 1 VF55_cf69 _8_1	Astroni 1 VF55_cf69 _9_1	Astroni 1 VF55_cf69 _9_2	Baia–Fondi di Baia VF55 _cf88_1_2	Pisani 1 CF25_cpx 86_m3	Pisani 1 CF25_cpx 129_m3	Pisani 1 CF25_cpx 54_m3	Pisani 1 CF25_cpx 118_m3	Pisani 1 CF25_cpx 129_m1	Pisani 1 CF25_cpx 129_m2
Type:	mg	mg	mg	mg	mi	mi	mi	mi	mi	mi
SiO <sub>2</sub>	58.84	56.79	58.68	58.39	52.97	55.86	55.81	52.47	57.52	55.12
Na <sub>2</sub> O	4.26	2.87	4.52	4.95	3.53	3.36	3.46	3.55	3.57	3.55
MgO	0.48	0.52	0.59	0.34	1⋅34 17⋅97	1.00	1⋅17 17⋅89	1⋅05 17⋅60	0.97	1·07 18·01
Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O	18⋅19 8⋅65	18⋅09 8⋅09	18⋅01 8⋅67	16⋅03 6⋅79	8·20	17.53 8.14	7.88	7.70	18⋅25 7⋅98	7.99
CaO	2.10	2.41	2.49	2.14	3.43	3.57	3.58	3.65	3.57	3.58
TiO <sub>2</sub>	0.40	0.48	0.45	0.54	0.62	0.59	0.63	0.64	0.58	0.56
MnO	0.18	0.14	0.20	0.26	0.02	0.21	0.10	0.13	0.22	0.22
FeOt	3.06	3.37	3.33	4.29	4.70	4.76	4.91	5.03	5.23	5.07
P <sub>2</sub> O <sub>5</sub> CI	0.09 0.86	0·10 0·85	0·13 0·92	0.05 1.15	0·45 0·60	0·24 0·69	0·25 0·78	0·28 0·71	0·18 0·77	0·21 0·83
F	0.09	0.26	0.29	0.23	0.19	0.19	0.21	0.23	0.24	0.23
H₂O Total	0.63 97.83	0⋅91 94⋅89	0.95 99.23	0·77 95·92	2·17 96·17	2·23 98·39	2·37 99·04	2⋅30 95⋅35	2·12 101·22	2·32 98·75
Eruption:	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pomici	Pomici	Pomici	Pomici	Pomici	Pomici
·					Principali	Principali	Principali	Principali	Principali	Principali
Analysis #:	VF55_cf25	VF55_cf25	VF55_cf25	VF55_cf25	CF6_cpx	CF6_cpx 98_m1	CF6_cpx 110_m1	CF6_cpx 121_m2	CF6_cpx 192_m1	CF6_cpx
Type:	_2_1a mg	_2_1b mg	_4_1 mg	_4_2 mg	29_m1 mi	mi	mi	mi	mi	_137_2 mi
SiO <sub>2</sub>	56-27	56-27	55.42	54-66	57.09	56.58	56.59	57.57	57.05	56-34
Na <sub>2</sub> O	3.34	3.34	3.08	3.17	4.06	3.70	3.83	3.39	3.24	3.95
MgO	1.53	1.53	2.08	1.96	0.82	0.83	0·54	0.77	0.72	0·54
Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O	17⋅82 8⋅13	17⋅82 8⋅13	17.90 7.11	18⋅21 7⋅27	18⋅24 9⋅28	18⋅19 9⋅22	18⋅21 9⋅18	18⋅45 9⋅08	18⋅11 8⋅84	18⋅15 9⋅10
CaO	4.27	4.27	5·52	5·35	3.16	3.18	2.95	3.17	3.14	3.06
TiO <sub>2</sub>	0.65	0.65	0.75	0.81	0.41	0.45	0.47	0.48	0.46	0.55
MnO	0.09	0.09	0.13	0.06	0.09	0.16	0.17	0.04	0.14	0.14
FeOt P <sub>2</sub> O <sub>5</sub>	5⋅41 0⋅42	5⋅41 0⋅42	6·21 0·50	6·07 0·49	3⋅67 0⋅14	3⋅92 0⋅14	3.94 0.17	3⋅82 0⋅12	3⋅89 0⋅12	3⋅90 0⋅11
CI	0.76	0.76	0.56	0.61	0.74	0.73	0.67	0.73	0.72	0.86
F	0.22	0.20	0.31	0.27	0.23	0.22	0.25	0.22	0.28	0.24
H₂O Total	0·14 99·06	0·14 99·04	0·05 99·63	0.33 99.26	1⋅26 99⋅21	3·84 101·16	2·52 99·47	1.98 99.82	3⋅35 100⋅05	2·35 99·30
Eruption:	Pomici	Pomici	Pomici	Pomici	Pomici	Pomici	Minopoli 1	Minopoli 1	Minopoli 1	Minopoli 1
·	Principali	Principali	Principali	Principali	Principali	Principali	•	•	•	•
Analysis #:	CF6_cpx _167_1	VF55_cf6 _1_1	VF55_cf6 _1_2	VF55_cf6 _1_3	VF55_cf6 _1_4	VF55_cf6 _1_6	CF13_cpx 29_m1	CF13_cpx 142_m1	CF13_cpx _140_1	CF13_cpx _135_1
Type:	mi	mg	mg	mg	mg	mg	mi	mi	mi	mi
SiO <sub>2</sub>	56-56	56-31	55.90	56.52	55.79	56-16	49-19	49.92	53.36	54.30
Na <sub>2</sub> O	3.54	3.71	3.21	3.25	3.45	3.67	2.61	3.54	2.36	3.44
MgO	0⋅83 18⋅41	1⋅03 17⋅98	0.92 17.96	0.98 18.45	1⋅03 17⋅74	1.00 17.99	3⋅26 18⋅29	2·23 18·30	3⋅56 16⋅98	2·07 18·35
Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O	9.00	8.90	8.94	9.16	9.00	8.98	5.78	7·45	5.82	7.62
CaO	3.17	3.64	3.43	3.68	3.86	3.69	7.30	5.83	6.35	5.35
TiO <sub>2</sub>	0.44	0.45	0.47	0.47	0.50	0.42	0.98	0.86	0.66	0.83
MnO FeOt	0·17 3·56	0·13 4·24	0·19 4·13	0·12 4·18	0·04 4·18	0·17 4·24	0·12 7·42	0·17 7·02	0·11 4·85	0·14 4·82
P <sub>2</sub> O <sub>5</sub>	0·15	0.17	0.12	0.16	0.22	0.14	0.66	0.63	0.68	0.63
CĪ	0.84	0.69	0.75	0.71	0.69	0.73	0.48	0.51	0.59	0.77
F	0.27	0.24	0.27	0.26	0.24	0.14	0.29	0.30	0.28	0.22
H₂O Total	2·69 99·63	0⋅32 97⋅82	0·10 96·39	1·10 99·03	0·33 97·07	0·35 97·69	1.87 98.25	1⋅30 98⋅06	3·05 98·64	1·25 99·80
Eruption:	Minopoli 1	Minopoli 1	Minopoli 1	Minopoli 1	Minopoli 1	Minopoli 1	Minopoli 1	Minopoli 1		RP
Analysis #:	CF13_cpx _55_1	CF13_cpx 49 2	CF13_cpx 119_m1	VF55_cf13 _1_2	VF55_cf13 _3_1	VF55_cf13 _4_1	VF55_cf13 _5_1	VF55_cf13 _10_1		
-	mi	mi	mi	mg	mg	mg	mg	mg		
Type:		49.23	48.34	53.43	53.34	49.92	53.38	53.01		0.42
SiO <sub>2</sub>	52.65			3.15	2.91	3.01	2.96	2.87		0.26
SiO <sub>2</sub> Na <sub>2</sub> O	2.96	2.40	2.32		0.40		0 4 4	0.00		
SiO <sub>2</sub> Na <sub>2</sub> O MgO	2.96 2.29	3.62	3.57	2.35	2·40 17·06	2·19 15·73	2·14 18·49	2·28 17·27		0.06
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub>	2·96 2·29 17·30	3.62 17.53	3·57 17·55	2·35 16·79	2·40 17·06 7·47	15.73	18-49	17.27		0.06 0.23
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO	2·96 2·29 17·30 6·50 5·63	3.62 17.53 5.09 7.79	3.57 17.55 4.88 7.51	2·35 16·79 7·17 5·82	17·06 7·47 5·65	15⋅73 6⋅65 5⋅47	18·49 6·64 6·27	17·27 6·67 5·82		0.06 0.23 0.20 0.12
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub>	2·96 2·29 17·30 6·50 5·63 0·92	3.62 17.53 5.09 7.79 0.98	3.57 17.55 4.88 7.51 0.83	2·35 16·79 7·17 5·82 1·04	17·06 7·47 5·65 1·01	15·73 6·65 5·47 0·97	18·49 6·64 6·27 0·92	17·27 6·67 5·82 0·95		0.06 0.23 0.20 0.12 0.08
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO	2.96 2.29 17.30 6.50 5.63 0.92 0.18	3.62 17.53 5.09 7.79 0.98 0.20	3.57 17.55 4.88 7.51 0.83 0.18	2·35 16·79 7·17 5·82 1·04 0·19	17·06 7·47 5·65 1·01 0·18	15·73 6·65 5·47 0·97 0·13	18·49 6·64 6·27 0·92 0·17	17·27 6·67 5·82 0·95 0·25		0.06 0.23 0.20 0.12 0.08 0.09
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt	2·96 2·29 17·30 6·50 5·63 0·92 0·18 6·90	3.62 17.53 5.09 7.79 0.98 0.20 7.93	3·57 17·55 4·88 7·51 0·83 0·18 7·82	2·35 16·79 7·17 5·82 1·04 0·19 6·98	17·06 7·47 5·65 1·01 0·18 6·94	15·73 6·65 5·47 0·97 0·13 6·74	18·49 6·64 6·27 0·92 0·17 7·02	17·27 6·67 5·82 0·95 0·25 7·32		0.06 0.23 0.20 0.12 0.08 0.09 0.25
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub>	2·96 2·29 17·30 6·50 5·63 0·92 0·18 6·90 0·44 0·67	3·62 17·53 5·09 7·79 0·98 0·20 7·93 0·65 0·55	3·57 17·55 4·88 7·51 0·83 0·18 7·82 0·70 0·55	2·35 16·79 7·17 5·82 1·04 0·19 6·98 0·68	17·06 7·47 5·65 1·01 0·18 6·94 0·76 0·55	15·73 6·65 5·47 0·97 0·13 6·74 0·59 0·73	18·49 6·64 6·27 0·92 0·17 7·02 0·59 0·50	17·27 6·67 5·82 0·95 0·25 7·32 0·67 0·62		0.06 0.23 0.20 0.12 0.08 0.09 0.25 0.02
SiO <sub>2</sub> Na <sub>2</sub> O MgO Al <sub>2</sub> O <sub>3</sub> K <sub>2</sub> O CaO TiO <sub>2</sub> MnO FeOt P <sub>2</sub> O <sub>5</sub>	2·96 2·29 17·30 6·50 5·63 0·92 0·18 6·90 0·44	3·62 17·53 5·09 7·79 0·98 0·20 7·93 0·65	3.57 17.55 4.88 7.51 0.83 0.18 7.82 0.70	2·35 16·79 7·17 5·82 1·04 0·19 6·98 0·68	17·06 7·47 5·65 1·01 0·18 6·94 0·76	15·73 6·65 5·47 0·97 0·13 6·74 0·59	18·49 6·64 6·27 0·92 0·17 7·02 0·59	17·27 6·67 5·82 0·95 0·25 7·32 0·67		0.06 0.23 0.20 0.12 0.08 0.09 0.25 0.02

F and  $H_2O$  were measured by SIMS; all other elements were measured by EPMA. Analysis type refers to textural association: mi, melt inclusion; mg, matrix glass. Concentrations are in wt %. RP is representative absolute precision with 1SD based on counting statistics for F and elements measured by EPMA, and a 95% confidence interval based on regression of calibration standards for  $H_2O$ .

Table 3: Major and volatile element analyses of representative Campi Flegrei apatite inclusions and microphenocrysts

Table 5. Iviaj	or aria voiat	ile eleffiellt a	naryses or re	presentativ	c campine	grerapatite	inclusions a	na micropin	crioci yata	
Eruption:	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte
Crystal #:	Nuovo CFmap2_C	Nuovo CFmap2_C	Nuovo CFmap2_C	Nuovo CFmap2_C			Nuovo CFmap2_C	Nuovo CFmap2_C	Nuovo CFmap2_C	Nuovo CFmap2_
Type:	F195_19 mpc	F195_16 mpc	F195_17 mpc	F195_8 mpc	F195_49 mpc	F195_38 mpc	F195_35 mpc	F195_30 mpc	F195_24 mpc	CF195_22 mpc
P <sub>2</sub> O <sub>5</sub>	38-91	41-20	41.37	41.38	41.63	40.09	41.37	41.34	40.92	40.93
SiO <sub>2</sub>	1.60	0.58	0.40	0.42	0.66	0.85	0.44	0.58	0.55	0.69
CaO	53.97	54.09	54.28	54.00	53.42	52.59	53.49	54.32	53.79	53.28
SO <sub>2</sub>	1.34	0.70	0.42	0.41	0.14	0.63	0.42	0.19	0.56	0.44
F	2.30	2.03	2.03	2.13	2.83	2.08	2.09	2.58	2.29	2.56
CI	0.07	0.92	0.90	0.93	0.88	1.02	0.91	0.86	1.04	0.89
OH	1.57	1.09	1.08	1.20	0.57	1.19	1.13	0.57	0.75	0.66
O≡F, CI, OH Total	1·72 98·04	1·58 99·03	1·57 98·91	1·67 98·80	1⋅66 98⋅47	1.67 96.78	1⋅62 98⋅23	1·55 98·89	1.55 98.35	1.59 97.86
Eruption:	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte
	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo
Crystal #:	CFmap2_C	CFmap2_C	CFmap2_C	CFmap2_C	CFmap2_C		CFmap2_C		CF195_cpx	
-	F195_45	F195_31	F195_12	F195_97	F195_46	F195_64	F195_74	F195_98	63_ap1	80_ap4
Type:	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	срх	срх
$P_{2}O_{5}$	40.97	40.47	41.32	40.99	41.39	41.06	40.36	41.36	41.29	41.35
SiO <sub>2</sub>	0.68	1.32	0.61	0.54	0.41	0.43	0.90	0.63	0.61	0.50
CaO	53.20	52.93	53.68	54.32	53.37	54.19	53.78	54.67	53.85	54.13
SO <sub>2</sub>	0.25	0.17	0.56	0.23	0.27	0.42	0.13	0.16	0.39	0.36
F	2.28	2.02	2.47	2.36	2.50	2.08	2.94	3.30	2.12	2⋅35 0⋅92
CI OH	1·05 1·02	0·94 1·23	0.98 0.48	1⋅04 0⋅70	0⋅89 0⋅55	1·04 0·94	0·54 0·49	0·12 0·17	0.98 0.62	0.92 0.45
O≡F, CI, OH	1.68	1.64	1.49	1.56	1.51	1.55	1.59	1.50	1.41	1.41
Total	97.77	97.44	98-61	98-62	97.87	98-61	97.55	98-91	98-44	98-64
Eruption:	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte	Monte
Liuption.	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo	Nuovo
Crystal #:	CF195_cpx	CF195_cpx		CF195 cpx			CF195 cpx		CF195_cpx	
,	83_ap3	80_ap1	128_ap1	128_ap4	133_ap1	137_ap1	173_ap1	121_ap1	188_ap1	208_ap1
Type:	срх	срх	срх	срх	срх	срх	срх	срх	срх	срх
$P_{2}O_{5}$	41.93	41.30	40.73	41.02	40.57	40.96	39.65	40.21	40.89	40.80
SiO <sub>2</sub>	0.37	0.45	0.79	0.72	0.48	0.60	1.32	0.67	0.67	0.76
CaO	53.36	53.87	53.18	53.80	53.63	53.84	53.03	53.76	53.94	54.25
SO <sub>2</sub>	0.21	0.36	0.34	0.25	0.57	0.55	0.12	0.49	0.40	0.50
F	2.23	2.35	2.30	2.36	2.07	2.51	2.06	2.52	2.53	2.00
CI	0.94	0.91	1.00	0.79	0.97	0.92	0.91	0.92	0.79	0.91
OH O≡F, CI, OH	0·47 1·37	0·45 1·41	0·53 1·44	0·52 1·42	0·67 1·40	0·46 1·48	0⋅84 1⋅46	0·46 1·48	0·39 1·43	0.74 1.39
Total	98-13	98.29	97.42	98.05	97.56	98.35	96.47	97.54	98-19	98.57
Eruption:	Monte	Monte	Monte	Monte	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1
2.4500	Nuovo	Nuovo	Nuovo	Nuovo	7 101. 01	7 1011 0111 1	, 101.01.1	7.00.01	, 101. 01.1 1	7.00.0111
Crystal #:	CF195_cpx	CF195_	CF195_	CF195_	CFmap2_	CFmap2_	CFmap2_	CFmap2_	CFmap2_	CFmap2_
	250_ap4	bt40_ap1	bt46_ap1	bt64_ap3	CF69_47	CF69_32	CF69_29	CF69_10	CF69_26	CF69_6
Type:	срх	bt	bt	bt	mpc	mpc	mpc	mpc	mpc	mpc
$P_2O_5$	38.59	40.83	39.77	41.62	41.50	40.85	42.62	41.65	41.26	41.40
SiO <sub>2</sub>	1.46	0.80	0.52	0.54	0.64	0.80	0.06	0.61	0.76	0.56
CaO	52.53	53.12	53.90	53.72	54.55	52.57	53.77	53.38	53.59	53.84
SO <sub>2</sub>	1.16	0.32	0.34	0.36	0.39	0.10	0.15	0.45	0.22	0.63
F Cl	2·38 0·93	2·36 1·03	2·41 0·87	2.48	3⋅16 0⋅53	2.58	2·59 0·94	2.68	2·56 0·96	2·49 0·96
OH	0.93	0.74	0.63	0·91 0·56	0.53	1.06 0.70	0.94	0·92 0·71	0.96	0.74
O≡F, CI, OH	1.41	1.58	1.51	1.51	1.59	1.65	1.65	1.67	1.66	1.61
Total	96.07	97.63	96.93	98.68	99.47	97.01	99-21	98.73	98-46	99.01
Eruption:	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1	Astroni 1
Crystal #:	CFmap2_C	CFmap2_CF	CFmap2_C	CFmap2_C	CFmap2_C	CFmap2_C	CFmap2_C	CFmap2_C	CFmap2_C	CFmap2_C
	F69_19	69_22	F69_20	F69_21	F69_49	F69_46	F69_91	F69_85	F69_2	F69_82
Type:	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	Мрс
$P_2O_5$	41.16	42.17	41.16	41.78	40.05	41.46	40-65	40-64	41.64	41.08
SiO <sub>2</sub>	0.67	0.61	0.66	0.62	1.10	0.68	0.64	0.65	0.61	0.74
CaO	53 39	53.78	53.77	53.98	53.01	53.35	53.79	53.86	54.23	53.95
SO <sub>2</sub>	0.14	0.16	0.45	0.28	0.18	0.17	0.15	0.69	0.29	0.15
F	2.49	2.51	2.42	2.43	2.61	2.33	2.01	2.56	2.50	2.48
CI	0.97	0.93	0.95	0.94	1.01	0.99	1.07	0.99	0.98	1.01
OH CLOH	0·76	0.77 1.63	0.77 1.60	0.77 1.60	0.79 1.70	0.81 1.50	0.80 1.46	0·57	0·62	0.65 1.59
O≡F, CI, OH Total	1·62 97·96	1·63 99·30	1.60 98.58	1·60 99·20	1⋅70 97⋅05	1.59 98.20	1·46 97·65	1·57 98·39	1·57 99·30	1.58 98.48
ı Otai	37.30	55.50	30.30	33.20	37.00	30.20	37.00	30.33	55.50	30.40

Table 3: Continued

Eruption: Crystal #:	Astroni 1 CFmap2_C	Astroni 1 CFmap2_C	Astroni 1 CF69_cpx							
Type:	F69_18 mpc	F69_74 mpc	10_ap2 cpx	11_ap1 cpx	32_ap1 cpx	35_ap1	69_ap1	73_ap2 cpx	95_ap2 cpx	104_ap1
	-			<u> </u>	<u> </u>					
$P_2O_5$ SiO <sub>2</sub>	41·04 0·57	41⋅01 0⋅64	41⋅31 0⋅68	40.99 0.76	41·30 0·77	41.68 0.66	41·24 0·68	41⋅07 0⋅82	40⋅95 0⋅67	40·50 0·78
CaO	53.67	54.30	54.05	53.92	54.36	54.34	54.09	53.78	53.81	53.95
SO <sub>2</sub>	0.40	0.22	0.09	0.32	0.30	0.29	0.15	0.24	0.24	0.29
F	2.57	2.30	2.32	2.13	2.22	2.40	2.36	2.30	2.33	2.34
CI	1.01	0.57	0.90	0.99	0.92	0.91	0.91	0.88	0.89	0.89
ОН	0.60	1.24	0.42	0.58	0.60	0.50	0.53	0.52	0.49	0.47
O≡F, CI, OH	1.59	1.68	1.37	1.40	1.42	1.45	1.45	1.41	1.41	1.41
Total	98-27	98-60	98.38	98-29	99.05	99.33	98-51	98-20	97.96	97.81
Eruption: Crystal #:	Astroni 1 cpx110_	Astroni 1 CF69_cpx	Astroni 1 CF69_cpx	Astroni 1 CF69 cpx	Astroni 1 CF69_cpx					
•	ap1 _	132_ap1	137_ap1	139_ap2	153_ap2	151_ap1	112_ap1	65_ap2	21_ap2	44_ap6
Type:	срх									
P <sub>2</sub> O <sub>5</sub>	41.28	41.13	41.24	41.33	41.18	40.57	41.39	41.33	41.36	40.94
SiO <sub>2</sub>	0.64	0.61	0.61	0.57	0.59	0.94	0.50	0.72	0.55	0.56
CaO	53.31	53.37	53.72	54.14	53.71	53·52	54.33	54·62	54.03	53.97
SO <sub>2</sub> F	0·12 2·25	0⋅38 1⋅86	0⋅32 2⋅35	0·18 2·46	0⋅33 2⋅22	0·15 2·31	0·47 2·36	0·17 2·45	0·25 2·30	0·46 2·30
Cl	2·25 0·96	0.88	2·35 0·89	2·46 0·87	1.00	2·31 0·92	2·36 0·91	2·45 0·91	2·30 0·96	2·30 0·86
OH	0.51	0.46	0.58	0.57	0.68	0.56	0.56	0.58	0.43	0.53
O≡F, CI, OH	1.40	1.20	1.46	1.48	1.48	1.44	1.46	1.51	1.39	1.41
Total	97-67	97.48	98-24	98.59	98-22	97.53	99.06	99-27	98.50	98-21
Eruption:	Astroni 1									
Crystal #:	CF69_cpx	CF69_bt								
Type:	73_ap2 cpx	23_ap2 bt	23_ap3 bt	38_ap3 bt	48_ap1 bt	97_ap1 bt	117_ap1 bt	156_ap1 bt	167_ap1 bt	35_ap1 bt
P <sub>2</sub> O <sub>5</sub>	41.07	41.60	41.18	41.09	41.14	41.32	41.55	40.85	40.13	41.64
SiO <sub>2</sub>	0.82	0.56	0.59	0.69	0.59	0.60	0.53	0.51	0.59	0.47
CaO	53.78	53.97	54.18	53.86	54.14	54.30	54.51	54.49	53.81	54.06
SO <sub>2</sub>	0.24	0.34	0.38	0.58	0.31	0.56	0.38	0.40	0.35	0.42
F	2.34	2.40	2.49	2.52	2.42	2.40	2.54	2.37	2.40	2.43
CI	0.94	0.90	0.93	0.93	0.90	0.92	0.96	0.91	0.88	0.89
OH	0.56	0.59	0.65	0.55	0.66	0.58	0.65	0.57	0.57	0.59
O≡F, CI, OH Total	1·46 98·29	1·49 98·89	1·56 98·85	1·53 98·68	1·53 98·62	1·49 99·19	1·59 99·53	1·47 98·63	1·48 97·26	1·50 99·00
Eruption:	Astroni 1	Accademia	Accademia	Accademia						
Crystal #:	CF69_bt	CFmap2_C	CFmap2_C	CFmap2_C						
_	27_ap1	149_ap1	98_ap1	164_ap5	164_ap1	158_ap1	170_ap1	F163_14	F163_5	F163_42
Type:	bt	mpc	mpc	mpc						
P <sub>2</sub> O <sub>5</sub>	41.78	40.49	40.98	40.95	41.25	41.75	40.97	41.50	40.93	39.97
SiO <sub>2</sub>	0.50	0.74	0.58	0.60	0.65	0.45	0.40	0.68	0·51	1.19
CaO	53.99	54.37	53.87	54.35	54.54	54.59	54.01	54.24	53.55	54.42
SO <sub>2</sub> F	0·29 2·38	0·27 2·43	0·54 2·42	0·27 2·48	0·28 2·53	0·44 2·40	0·40 2·40	0·24 2·96	0·44 3·26	1⋅01 2⋅22
CI	0.89	0.90	0.91	0.91	0.93	0.92	0.91	0.66	0.48	0.06
OH	0.53	0.65	0.53	0.66	0.64	0.53	0.53	0.21	0.07	1.62
O≡F, CI, OH	1.45	1.53	1.48	1.56	1.58	1.47	1.46	1.49	1.51	1.71
Total	98-92	98-32	98-36	98-65	99.25	99-61	98-15	99.00	97.73	98.78
Eruption:	Accademia									
Crystal #:	CFmap2_C									
_	F163_25	F163_24	F163_28	F163_20	F163_46	F163_39	F163_32	F163_8	F163_6	F163_11
Type:	mpc									
P <sub>2</sub> O <sub>5</sub>	40.90	41.38	41.20	40.94	41.73	40.88	41.21	40.93	40.99	40.68
SiO <sub>2</sub>	0.57	0.66	0.77	0.63	0.35	0.77	0.62	0.57	0.58	0.67
CaO	54.08	54.08	53.47	53.47	53.90	53.81	53.64	53.72	53.68	53.35
SO <sub>2</sub>	0·54	0.27	0.22	0.17	0.35	0.14	0.16	0.65	0.33	0.25
F	3.62	2.54	2.49	2·61	2.53	2.62	2.52	2.32	2.64	2·63
CI OH	0·06 0·01	1⋅04 0⋅48	0.97 0.56	1⋅17 0⋅36	0⋅89 0⋅41	0·95 0·54	0.96 0.56	0·89 0·59	0⋅88 0⋅45	1⋅17 0⋅51
On O≡F, Cl, OH	1.54	1.53	1.53	1.53	1.46	1.57	1.54	1.46	1.52	1.61
Total	98.24	98.92	98.15	97.82	98.70	98.14	98.13	98.21	98.03	97.65
. 5.01	30-24	30 32	30 13	37.02	30.70	50 14	30 13	30-21	50.00	57.00

Table 3: Continued

Eruption:	Accademia	Accademia	Accademia	Accademia	Accademia	Accademia	Accademia	Santa Maria delle Grazie	Santa Maria delle Grazie	Santa Maria delle Grazie
Crystal #: Type:	CFmap2_C F163_10 mpc	mpc_C F16310 mpc	mpc-C F163_9 mpc	mpc-C F163_8 mpc	CFmap2_C F163_100 mpc	CFmap2_C F163_63 mpc	CFmap2_C F163_12 mpc	CF200_ map30 mpc	CF200_ map29 mpc	CF200_ map35 mpc
$P_2O_5$	40.51	41.44	40.36	41.10	41.06	41.15	41.37	40.65	40.71	39.27
SiO <sub>2</sub>	0.76	0.61	0.82	0.67	0.77	0.60	0.53	0.66	0.61	1.53
CaO SO <sub>2</sub>	53⋅20 0⋅23	54·41 0·17	53.97 0.10	54·22 0·23	54·29 0·13	54⋅37 0⋅31	53.97 0.16	54·06 0·46	53·94 0·37	52⋅93 0⋅24
F	2.83	3.29	3.69	3.78	3.79	3.77	3.60	2.85	3.54	2.67
CI	0.92	0.09	0.06	0.10	0.05	0.05	0.07	0.68	0.23	1.04
OH O≡F, CI, OH	0·46 1·62	0·00 1·41	0·00 1·57	0.00	0.00	0.00 1.60	0.00 1.53	0·18 1·44	0.00	0⋅33 1⋅51
Total	97.29	98-60	97.43	1.61 98.49	1.61 98.48	98-65	98.17	98.10	1⋅54 97⋅86	96.50
Eruption:	Santa Maria delle Grazie		Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia
Crystal #:	CF200_		. –	CFmap1_C	. –	. –	. –	. –	. –	. –
Type:	map36 mpc	F88_42 mpc	F88_49 mpc	F88_34 mpc	F88_16 mpc	F88_11 mpc	F88_10 mpc	F88_18 mpc	F88_33 mpc	F88_43 mpc
P <sub>2</sub> O <sub>5</sub>	41.14	39.99	40.97	41.23	40.52	40.42	39.43	41.07	39.95	40.29
SiO <sub>2</sub>	0.59	0.49	0.50 54.14	0.42	0.61 54.01	0.54	1.02	0.60	0.94	0.79 52.10
CaO SO₂	53.86 0.33	53·95 0·39	54·14 0·34	53.98 0.43	54·01 0·20	54·05 0·39	53·80 0·90	53·64 0·16	52⋅89 0⋅56	53⋅19 0⋅13
F F	3.57	2.06	2.07	2.29	2.10	2.17	2.18	1.99	1.98	2.32
CI	0.09	0.85	0.89	0.87	0.86	0.78	0.72	0.87	0.86	0.91
OH CLOU	0.00	0.90	1.01	0.95	0.97	0.82	0.88	0.82	0.83	0.91
O≡F, CI, OH Total	1·52 98·06	1·48 97·15	1·55 98·37	1.61 98.56	1·53 97·74	1·48 97·69	1·49 97·44	1·42 97·73	1·42 96·59	1.61 96.93
Eruption:	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia	Baia–Fondi di Baia
Crystal #:		CFmap1_C								
,	F88_30	F88_3	F88_8	F88_50	F88_44	F88_31	F88_45	F88_22	F88_17	F88_12
Type:	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc
P <sub>2</sub> O <sub>5</sub> SiO <sub>2</sub>	41·05 0·40	41·61 0·00	40·04 0·79	39·79 0·86	40·04 0·80	38·69 1·40	40·03 1·00	39·84 1·15	39·93 1·20	40·03 0·97
CaO	54.35	55.46	54.43	53.96	54.28	54.42	53.65	53.36	52.64	53.86
SO <sub>2</sub>	0.40	0.00	0.89	0.85	0.77	1.38	0.05	0.48	0.03	0.52
F	2.00	2.50	2.77	2.80	2.34	2.57	3.02	2.69	2.77	2.56
CI OH	0.89 0.88	0·03 0·92	0·09 0·93	0·09 0·79	0.06 1.16	0·04 0·97	0·42 0·47	0.62 0.63	0.55 0.68	0·61 0·59
O≡F, CI, OH	1.46	1.49	1.62	1.57	1.54	1.55	1.59	1.57	1.61	1.49
Total	98-51	99.03	98.32	97.57	97.91	97.92	97.05	97.20	96-19	97.65
Eruption:	di Baia	Baia-Fondi di Baia	di Baia	di Baia	di Baia	di Baia	di Baia	Pisani 1	Pisani 1	Pisani 1
Crystal #:	CFmap1_C F88_1	CFmap1_C F88_7	CF88_cpx 118_ap1	CF88_cpx 136_ap1	CF88_cpx 85_ap1	CF88_cpx 116_ap1	CF88_bt 36_ap1	F25_41	CFmap1_C F25_29	F25_19
Type:	mpc	mpc	срх	срх	cpx	срх	bt	mpc	mpc	mpc
P <sub>2</sub> O <sub>5</sub>	40.15	40.73	39.71	40.49	40.02	39.79	40.38	40.55	41-63	40.39
SiO <sub>2</sub>	0.72	0.45	1.50	1.33	1.25	1.28	1.08	0.77	0.42	0.66
CaO SO₂	54·23 0·72	54·03 0·41	52·48 0·00	52·90 0·00	53.55 0.56	52.78 0.00	53.70 0.53	53.02 0.36	54·27 0·37	53·43 0·68
F F	2.69	2.02	2.79	2.88	2.48	2.84	2.72	2.18	2.14	2.14
CI	0.61	0.89	0.42	0.43	0.63	0.41	0.54	1.15	0.94	1.10
OH	0.67	0.96	1.03	1.06	1.42	1.00	0.91	0.82	1.04	1.09
O≡F, CI, OH Total	1.59 98.20	1.50 97.99	1.76 96.18	1⋅81 97⋅28	1·86 98·05	1⋅76 96⋅34	1.69 98.16	1.56 97.29	1⋅60 99⋅21	1.66 97.83
Eruption:	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1
Crystal #:		CFmap1_C								
Type:	F25_7 mpc	F25_44 mpc	F25_5 mpc	F25_2 mpc	F25_22 mpc	F25_30 mpc	F25_18 mpc	F25_46 mpc	F25_37 mpc	F25_50 mpc
P <sub>2</sub> O <sub>5</sub>	41.14	42.19	39.38	41.82	38.27	38.17	40 64	41.45	42.14	41.11
SiO <sub>2</sub>	0.59	0.34	1.16	0.29	1.46	1.59	0.75	0.29	0.24	0.68
CaO	53.65	53.91	54.16	54.10	52.97	53.33	53.70	53.52	54.00	53.91
SO <sub>2</sub> F	0.69 1.96	0·34 2·12	1⋅21 2⋅07	0.06 2.83	1⋅18 3⋅61	1⋅24 3⋅34	0.63 1.98	0·09 2·07	0·03 2·86	0·13 1·95
Cl	1.96	2·12 0·92	0.06	2·83 0·55	ا 0۰12 0۰12	0.09	0.58	1.00	2·86 0·56	0.89
OH	0.95	0.99	1.39	0.35	0.33	0.40	1.21	0.76	0.32	0.96
$O \equiv F$ , CI, OH	1.51	1.57	1.54	1.48	1.70	1.61	1.53	1.45	1.48	1.47
Total	98.52	99-24	97.89	98.52	96-24	96.55	97.96	97.73	98-67	98-16

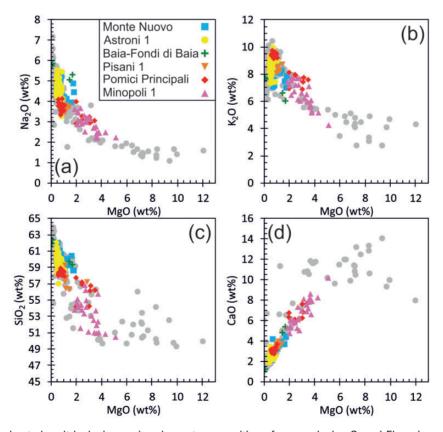
Table 3: Continued

Table 3: Con	itinuea									
Eruption: Crystal #:	Pisani 1 CFmap1_C F25_11	Pisani 1 CFmap1_C F25_9	Pisani 1 CFmap1_C F25_3	Pisani 1 CFmap1_C F25_1	Pisani 1 CFmap1_C F25_16	Pisani 1 CFmap1_C F25_27	Pisani 1 CFmap1_C F25_49	Pisani 1 CF25_cpx 98_ap2	Pisani 1 CF25_cpx 98_ap5	Pisani 1 CF25_cpx 121_ap1
Type:	mpc	mpc	mpc	mpc	mpc	mpc	mpc	срх	срх	срх
$P_2O_5$	39.29	39.66	40.88	42.41	42.02	42.09	41.94	40.79	40.22	40.04
SiO <sub>2</sub>	0.91	1.10	0.67	0.00	0.24	0.25	0.24	0.40	0.78	0.54
CaO	53.46	53.41	54.32	54.89	53.88	54-21	54.03	54.44	54.04	53.77
SO <sub>2</sub>	1.28	0.38	0.67	0.02	0.05	0.09	0.03	0.25	0.58	0.63
F Cl	2·02 1·10	2·58 0·74	1⋅96 1⋅00	2·59 0·02	3⋅03 0⋅45	2·93 0·06	2⋅80 0⋅51	1.92 0.82	1.92 0.91	2·03 1·02
OH	0.90	0.69	0.93	0.99	0.28	0.44	0.48	1.01	1.01	1.04
$O \equiv F$ , CI, OH	1.52	1.58	1.49	1.56	1.51	1.45	1.52	1.47	1.49	1.58
Total	97-44	96.98	98-94	99-36	98-44	98-62	98-51	98-16	97.97	97.50
Eruption:	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pisani 1	Pomici Principali	Pomici Principali	Pomici Principali
Crystal #:	CF25_cpx 122_ap1	CF25_cpx 142_ap2	CF25_bt 40_ap1	CF25_bt 67_ap2	CF25_bt 98_ap1	CF25_bt 112_ap1	CF25_bt 110_ap1	CFmap2_C F6_33	CFmap2_C F6_37	CFmap2_C F6_39
Type:	cpx	cpx	bt	bt	bt	bt	bt	mpc	mpc	mpc
P <sub>2</sub> O <sub>5</sub>	40.88	40.68	41.56	40.36	41.27	39-20	41.47	40.86	38.06	41.25
SiO <sub>2</sub>	0.40	0.42	0.37	0.39	0.49	0.50	0.37	0.77	1.86	0.66
CaO	54.52	54.45	54.81	54.55	54.81	54.46	54.13	53.59	52.96	54.30
SO <sub>2</sub> F	0·46 1·89	0⋅37 1⋅96	0⋅37 2⋅07	0·42 1·93	0·46 1·93	0⋅36 2⋅03	0⋅31 2⋅03	0.79 3.80	1·48 2·94	0.66 1.89
CI	0.89	1.90	1.01	0.84	0.84	0.99	1.01	0.03	0.07	0.58
OH	1.10	0.81	0.93	1.20	0.97	1.04	0.90	0.24	1.21	1.97
$O \equiv F$ , CI, OH	1.52	1.44	1.54	1.57	1.46	1.57	1.50	1.72	1.82	1.85
Total	98-63	98-25	99-57	98-12	99-31	97.02	98.72	98-36	96.76	99.46
Eruption:	Pomici	Pomici	Pomici	Pomici	Pomici	Pomici	Pomici	Pomici	Pomici	Pomici
Crystal #:	Principali CFmap2_C	Principali CFmap2_C	Principali CFmap2_C	Principali CFmap2_C	Principali CFmap2 C	Principali CFmap2 C	Principali CFmap2_C	Principali CFmap2_C	Principali CFmap2_C	Principali CFmap2 C
Crystal #.	F6_47	F6_25	F6_24	F6_30	F6_32	F6_16	F6_10	F6_13	F6_6	F6_5
Type:	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc
$P_2O_5$	41-61	41-66	41.79	42.03	40.00	40.46	41.71	41.03	42.00	42.06
SiO <sub>2</sub>	0.49	0.52	0.48	0.35	1.29	0.70	0.52	0.60	0.60	0.39
CaO	54.18	54.00	54.34	54·16 0·48	52.40	53.33	53.74	53.92	54.04	54·16 0·43
SO₂ F	0⋅38 1⋅90	0⋅39 2⋅66	0·44 2·43	2.14	0·18 2·40	0·23 2·10	0⋅34 2⋅06	0.75 2.07	0·22 2·07	2.08
CI	0.63	0.65	0.76	0.75	1.02	0.89	0.88	1.01	0.91	0.90
OH	1.50	0.91	1.01	1.31	1.16	1.09	1.12	1.12	1.15	1.14
O≡F, CI, OH	1.65	1.69	1.67	1.69	1.79	1.60	1.59	1.63	1.62	1.62
Total	99.04	99.10	99.58	99.53	96.66	97.20	98.78	98.87	99.37	99.54
Eruption:	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali
Crystal #:	CFmap2 C	CFmap2 C	CFmap2 C	CFmap2 C	CFmap2_C	CFmap2_C	CFmap2 C	CFmap2_C	CF6_cpx	CF6_cpx
,	F6_46	F6_21	F6_100	F6_15	F6_26	F6_98	F6_92	F6_62	19_ap1	64_ap1
Type:	mpc	mpc	mpc	mpc	mpc	mpc	mpc	mpc	срх	cpx
$P_2O_5$	41.36	42.04	39.99	41.04	41.45	40.93	40.23	40.05	41.35	40.50
SiO <sub>2</sub>	0.67	0.42	1.17	0.53	0.48	0.65	1.10	1.11	0.43	0.68
CaO SO₂	53·89 0·77	53·64 0·34	52⋅68 0⋅07	53·97 0·75	53⋅87 0⋅44	54·01 0·72	53·31 0·19	54·96 1·09	54·66 0·27	53⋅97 0⋅52
50 <sub>2</sub> F	0.77 2.21	2.06	2.39	0.75 2.12	2.84	2.06	1.97	2.20	2.08	2.00
CI	1.11	0.90	1.05	1.04	0.75	1.00	0.95	0.07	0.85	0.97
OH	0.95	1.02	0.85	0.88	0.19	0.89	0.99	1.30	0.93	1.35
O≡F, CI, OH	1.63	1.55	1.64	1.54	1.45	1.51	1.51	1.55	1.51	1.70
Total	99.33	98.87	96.56	98.79	98.57	98.75	97.23	99.23	99.06	98-28
Eruption:	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Pomici Principali	Minopoli 1				
Crystal #:	CF6_cpx 96_ap1	CF6_cpx 138_ap1	CF6_bt 94_ap1	CF6_bt 112_ap1	CF6_bt 26_ap1	CFmap1_C F13_39	CFmap1_C F13 50	CFmap1_C F13_15	CFmap1_C F13_16	CFmap1_C F13_3
Type:	срх	срх	bt	bt	bt	mpc	mpc	mpc	mpc	mpc
P <sub>2</sub> O <sub>5</sub>	40.30	40.41	40.31	40.72	40.75	40.93	40.66	40.58	39.82	40.85
SiO <sub>2</sub>	0.70	0.96	0.93	0.68	0.76	0.60	0.48	0.89	0.75	0.37
CaO	53.58	53.94	53.39	54.16	54.04	54.34	54.54	53.57	53.63	54.34
SO <sub>2</sub>	0.51	0·64	0.60	0·51	0·55	0.63	0·18	0.81	0.81	0.36
F CI	2·14 0·93	1.94 0.90	2·26 0·95	2·22 0·97	2·27 0·97	2·16 0·55	3·11 0·03	2·27 0·74	1.94 0.63	2·03 1·08
OH	0.33	1.08	1.14	1.16	1.18	1.08	0.39	1.51	1.38	0.78
$O \equiv F$ , $CI$ , $OH$	1.50	1.53	1.70	1.70	1.73	1.54	1.50	1.83	1.61	1.47
Total	97-47	98.34	97.87	98.72	98-80	98.75	97.89	98.54	97.35	98.34

Table 3: Continued

Eruption: Crystal #:	Minopoli 1 CFmap1_C									
	F13_4	F13_7	F13_22	F13_31	F13_27	F13_2	F13_10	F13_42	F13_19	F13_8
Type:	mpc									
P <sub>2</sub> O <sub>5</sub>	39.50	40.81	40.04	41.03	40.78	40.13	40.70	41.37	41.39	42.39
SiO <sub>2</sub>	0.84	0.51	1.20	0.42	0.63	1.00	0.45	0.41	0.50	0.06
CaO	54.03	53.92	53.26	54.06	53.77	52.90	53.86	53.81	54.92	54.64
$SO_2$	0.74	0.59	0.17	0.44	0.33	0.19	0.31	0.42	0.47	0.00
F	2.24	2.08	1.94	1.83	2.86	2.08	2.00	1.96	1.71	2.84
CI	0.05	0.85	0.90	0.81	0.93	0.94	0.88	0.86	0.41	0.12
OH	1.35	1.01	1.10	0.98	0.24	1.04	0.97	0.88	1.29	0.51
$O \equiv F$ , CI, OH	1.59	1.54	1.54	1.41	1.53	1.58	1.50	1.43	1.42	1.46
Total	97.16	98-23	97.07	98-16	98-01	96.70	97-67	98-28	99-27	99.10
Eruption:	Minopoli 1		RP							
Crystal #:	CFmap1_C	CFmap1_C	CFmap1_C	CFmap1_C	CFmap1_C	CFmap1_C	CF13_cpx	CF13_bt		
	F13_30	F13_43	F13_28	F13_41	F13_49	F13_32	169_ap1	114_ap1		
Type:	mpc	mpc	mpc	mpc	mpc	mpc	срх	bt		
P <sub>2</sub> O <sub>5</sub>	38-93	40.24	40.59	39.39	41.19	41.08	41.91	41.47		0.45
SiO <sub>2</sub>	1.34	0.60	0.58	1.20	0.39	0.53	0.42	0.55		0.05
CaO	54.05	54.82	53-61	54.10	53.81	54.51	54.40	54.15		0.50
SO <sub>2</sub>	1.36	0.44	0.74	0.97	0.11	0.44	0.41	0.55		0.03
F	2.45	3.24	2.32	1.81	2.02	2.12	1.94	1.94		0.18
CI	0.06	0.02	0.56	0.48	0.70	0.88	0.76	0.83		0.15
OH	1.30	0.45	0.97	1.37	0.81	1.01	0.89	1.06		0.08
$O \equiv F$ , CI, OH	1.66	1.58	1.56	1.51	1.39	1.57	1.41	1.50		
Total	97.83	98-23	97.81	97.81	97.64	99.00	99.32	99.04		

Halogens and OH were measured by SIMS; all other elements were measured by EPMA. Crystal type refers to textural association: cpx, clinopyroxene-hosted inclusion; bt, biotite-hosted inclusion; mpc, microphenocryst. Concentrations are in wt %. RP is representative absolute precision with 1SD based on counting statistics for elements measured by EPMA, and a 95% confidence interval based on regression of calibration standards for elements measured by SIMS.



**Fig. 2.** Clinopyroxene-hosted melt inclusion major element compositions from explosive Campi Flegrei eruptions. All data were measured by EPMA and are normalized to 100%. Data are distinguished by eruption (see legend). Grey points show a compilation of literature clinopyroxene- and olivine-hosted melt inclusion compositions in <15 ka Campi Flegrei deposits from Cannatelli *et al.* (2007), Mangiacapra *et al.* (2008), Arienzo *et al.* (2010, 2016) and Fourmentraux *et al.* (2012).

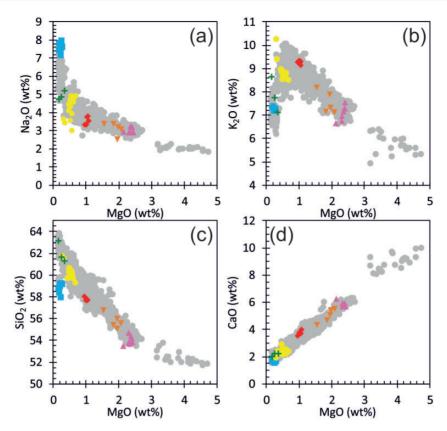


Fig. 3. Matrix glass major element compositions from explosive Campi Flegrei eruptions. All data were measured by EPMA and are normalized to 100%. Data are distinguished by eruption (see legend in Fig. 2). Grey points are literature matrix glass compositions in <15 ka Campi Flegrei deposits from Smith *et al.* (2011).

depleted and Na<sub>2</sub>O-rich (Fig. 3). This is consistent with fractional crystallization models, which indicate early fractionation of olivine, clinopyroxene and magnetite (Fowler *et al.*, 2007) but show a major change in compatibility for  $K_2O$  and  $Na_2O$  at  $\sim 0.5\,\mathrm{wt}$  % MgO, after plagioclase, biotite and K-feldspar come onto the liquidus (Fowler *et al.*, 2007; Cannatelli, 2012; Stock *et al.*, 2016).

## Glass volatile compositions

Campi Flegrei melt inclusions analysed in this study show that CI generally increases as MgO decreases, consistent with previous analyses (Fig. 4a). Minopoli 1 melt inclusions have the lowest CI concentrations (0.32-0.85 wt %), whereas melt inclusions from Pisani 1 and PP are intermediate (0.58-0.89 wt %). Melt inclusions from other eruptions have higher CI concentrations, extending up to 1.17 wt % (Fig. 4a; Supplementary Data Table S1). Matrix glass CI contents are typically similar to melt inclusions; the most CI-depleted matrix glasses (0.50-0.79 wt %) are from Minopoli 1 and Pisani 1, with the highest CI concentrations in matrix glasses from Monte Nuovo and B-FdB reaching 1.20 wt % (Fig. 4b; Supplementary Data Table S1). Some low-MgO (<0.5 wt %) matrix glasses from Astroni 1 and B-FdB have low CI contents (<0.6 wt %), in agreement with the greater variability in matrix glass CI concentrations at low MgO previously reported in the literature (Fig. 4b).

Mafic ( $\gtrsim$ 2 wt % MgO) melt inclusions from Campi Flegrei show scattered F contents, with concentrations up to 0.46 wt % F in Minopoli 1 (Fig. 4c). In more evolved melt inclusions ( $\lesssim$ 2 wt % MgO), F is typically higher in inclusions with lower MgO contents, from <0.1 wt % in Pisani 1 and PP to ~0.4 wt % in B–FdB (Fig. 4c; Supplementary Data Table S1). Fluorine concentrations up to 0.61 wt % have been reported in very low MgO melt inclusions (Fourmentraux *et al.*, 2012). In general, matrix glass F concentrations are approximately constant ( $\sim$ 0.1–0.3 wt %; Fig. 4d). However, F is significantly enriched in low-MgO (<0.5 wt %) matrix glasses from Monte Nuovo (typically 0.6–0.8 wt % F), with a single analysis reaching 1.26 wt %.

Melt inclusions analysed in this study typically contain  $\sim$ 2–4 wt % H<sub>2</sub>O and show no correlation between H<sub>2</sub>O and MgO (Fig. 4e), consistent with other Campi Flegrei eruptions (Arienzo *et al.*, 2016). Some inclusions have H<sub>2</sub>O concentrations as low as 1.04 wt %, with the largest variability at low MgO concentrations (<1.5 wt %), where measured H<sub>2</sub>O concentrations extend to >5 wt %, within the range previously reported for <15 kyr Campi Flegrei eruptions (up to 6.96 wt %; Cannatelli *et al.*, 2007). Matrix glass H<sub>2</sub>O contents are lower than in melt inclusions; mafic Minopoli 1 and Pisani 1 matrix glasses consistently have H<sub>2</sub>O concentrations <0.4 wt %, whereas more evolved (<1 wt % MgO) PP, Astroni 1 and Monte Nuovo matrix glasses

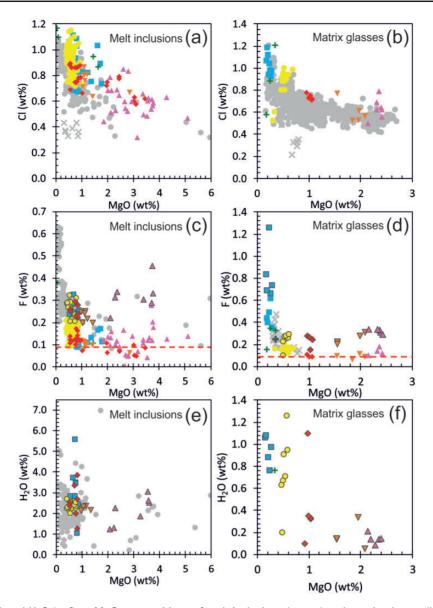


Fig. 4. Halogens (a–d) and  $H_2O$  (e, f) vs MgO compositions of melt inclusions (a, c, e) and matrix glasses (b, d, f) from explosive Campi Flegrei eruptions. MgO and CI were measured by EPMA,  $H_2O$  was measured by SIMS, and F was measured by EPMA (points with no outlines) and SIMS (points with grey outlines). The red dashed lines show EPMA F detection limits of 0·09 wt %. Data in (a)–(d) are normalized to 100%. Data are distinguished by eruption (see legend in Fig. 2). Grey circles show literature melt inclusion and matrix glass compositions in <15 ka Campi Flegrei deposits, from the same sources as Figs 2 and 3, respectively. The grey crosses show the compositions of Campanian Ignimbrite melt inclusions and matrix glasses from Signorelli *et al.* (1999).

have more variable  $H_2O$  concentrations, extending to >1 wt % (Fig. 4f).

## Apatite volatile compositions

## Apatite inclusions

Clinopyroxene-hosted apatite inclusions have measured F concentrations from 1.86 to 2.92 wt %, Cl concentrations from 0.41 to 1.27 wt % and OH concentrations from 0.39 to 1.42 wt % (Table 3; Supplementary Data Table S2). This translates to  $X_{\rm Cl}/X_{\rm OH}$  ratios from 0.19 to 1.22,  $X_{\rm F}/X_{\rm OH}$  from 1.33 to 5.79 and  $X_{\rm F}/X_{\rm Cl}$  from 2.96 to 12.8 (where  $X_{\rm F}$ ,  $X_{\rm Cl}$  and  $X_{\rm OH}$  are the mole fractions of F, Cl and OH, respectively). Variations in the apatite volatile site can typically be described by an F–OH

exchange, with an approximately constant CI component (Fig. 5a). Baia–Fondi di Baia is an exception, where CI in clinopyroxene-hosted apatite inclusions is notably depleted relative to inclusions in other eruptions. In general, clinopyroxene-hosted apatite inclusions show a positive correlation between  $X_{\rm CI}/X_{\rm OH}$  and  $X_{\rm F}/X_{\rm OH}$  and approximately constant  $X_{\rm F}/X_{\rm CI}$  (Fig. 5b and c).

Biotite-hosted apatite inclusions have a more restricted compositional range than clinopyroxene-hosted inclusions, with F, Cl and OH ranging from 1.81 to 2.76 wt %, 0.54 to 1.17 wt % and 0.53 to 1.20 wt %, respectively (Table 3; Supplementary Data Table S2). This translates to  $X_{\rm Cl}/X_{\rm OH}$  ratios from 0.28 to 0.95,  $X_{\rm F}/X_{\rm OH}$  from 1.20 to 4.11 and  $X_{\rm F}/X_{\rm Cl}$  from 3.31 to 9.47. Biotite-

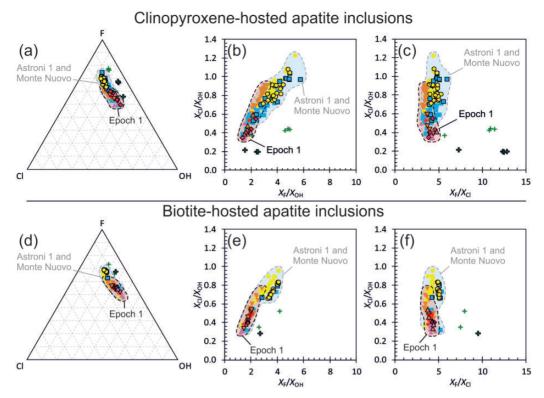


Fig. 5. Volatile compositions of clinopyroxene- (a–c) and biotite-hosted (d–f) apatite inclusions from explosive Campi Flegrei eruptions. Data are presented both in ternary space (a, d) and on binary plots of  $X_{\text{Cl}}/X_{\text{OH}}$  vs  $X_{\text{F}}/X_{\text{OH}}$  (b, e) and  $X_{\text{Cl}}/X_{\text{OH}}$  vs  $X_{\text{F}}/X_{\text{Cl}}$  (c, f). Points with black outlines were measured by SIMS and points without outlines were measured by EPMA. Data are distinguished by eruption (see legend in Fig. 2). The coloured fields highlight EPMA and SIMS data from Epoch 1 (red), and Astroni 1 (Epoch 3) and the historical Monte Nuovo eruptions (blue).

hosted inclusions plot on the same compositional trends as clinopyroxene-hosted inclusions (Fig. 5d–f). However, within an individual eruption, biotite-hosted inclusions are typically offset to more OH-rich compositions, with lower  $X_{\rm Cl}/X_{\rm OH}$  and  $X_{\rm F}/X_{\rm OH}$  ratios (Table 3; Supplementary Data Table S2).

## Apatite microphenocrysts

Apatite microphenocrysts show more compositional diversity than inclusions (Figs 5-7), with a small minority of F concentrations measured by EPMA reaching 4.54 wt %. These high F contents are almost exclusively in lava samples (Supplementary Data Table S2) but exceed the maximum stoichiometric limit of 3.76 wt % F (Pyle et al., 2002), probably reflecting electron beam-induced sample damage in near end-member fluorapatite (Stormer et al., 1993; Goldoff et al., 2012; Stock et al., 2015). The highest F concentrations measured by SIMS (3-80 wt %; Table 3) are stoichiometric within analytical uncertainty. In apatite microphenocrysts, the minimum F concentration is 1.67 wt %, CI ranges from 0.02 to 1.47 wt % and measured OH concentrations are <1.97 wt % (Table 3; Supplementary Data Table S2). This translates to overall  $X_{CI}/X_{OH}$ ,  $X_E/X_{OH}$  and  $X_E/X_{CI}$ ratios from  $\sim$ 0 to 16.5, 0.86 to 781 and 2.85 to 374, respectively, with  $X_{\rm Cl}/X_{\rm OH} \le 2.27$  and  $X_{\rm F}/X_{\rm OH} \le 25.8$  in explosive eruptions. Most apatite microphenocrysts plot on the same compositional trends as the inclusions (Figs 6 and 7) but typically extend to more OHrich compositions with lower  $X_{CI}/X_{OH}$  and  $X_F/X_{OH}$ ratios. Compositional differences between apatite inclusions and microphenocrysts along this main trend were used by Stock et al. (2016) to give temporal context to apatite data from Astroni 1. Two subsidiary sets of apatite microphenocrysts are observed in some explosive eruptions, both with lower Cl contents than the main trend (Figs 6 and 7). One has low  $X_{\rm CI}/X_{\rm OH}$  ratios with  $X_{\rm F}/X_{\rm OH}$  and  $X_{\rm F}/X_{\rm CI}$  ratios that extend to very high values (i.e. in Monte Nuovo, Astroni 1, B-FdB, Pisani 1, PP, Minopoli 1); the other has high  $X_{\rm CI}/X_{\rm OH}$  and  $X_{\rm F}/X_{\rm OH}$  ratios (i.e. in Astroni 1, B–FdB, PP, Minopoli 1). Additionally, a small number of inclusions from Pisani 1 define a trend towards high  $X_{\rm F}/$  $X_{\rm OH}$  at similar  $X_{\rm CI}/X_{\rm OH}$  ratios to the main population (Fig. 7d-f). Apatite microphenocrysts from lava deposits (Accademia and SMdG) are distinct; in ternary space they typically show an increasing F component at the expense of CI, trending first towards the F-CI join and then curving towards the F apex, but in binary space they are scattered with no discernible trend (Fig. 6g-I).

## Differences between eruptions

The SIMS analyses are more precise than the EPMA data, so subtle differences in apatite compositions

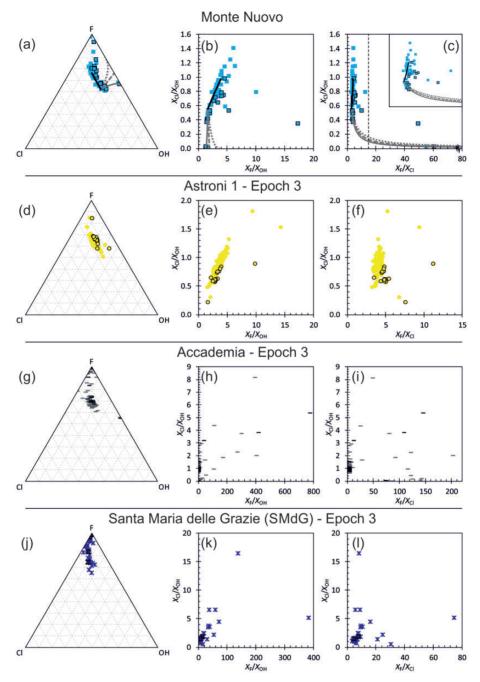


Fig. 6. Volatile compositions of apatite microphenocrysts from the historical Monte Nuovo (a–c) and Epoch 3 (d–l) eruptions of Campi Flegrei. Data are presented both in ternary space (a, d, g, j) and on binary plots of  $X_{\rm Cl}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm OH}$  (b, e, h, k) and  $X_{\rm Cl}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm Cl}$  (c, f, i, l). Inset in (c) shows an expansion of the area outlined by the black dashed box in the main panel. Points with black outlines were measured by SIMS and points without outlines were measured by EPMA. Non-stoichiometric analyses with F >3·76 wt % (i.e. in lava deposits) are plotted as end-member fluorapatite, with non-stoichiometry probably the result of F migration during EPMA (Stormer *et al.*, 1993; Goldoff *et al.*, 2012; Stock *et al.*, 2015). Data are distinguished by eruption (see headings). Lines in (a)–(c) show the theoretical trajectory of apatite compositional evolution predicted by our thermodynamic model that best fits natural apatite compositions from Campi Flegrei.  $C_{\rm Cl}^{\rm O}=0.4\,{\rm wt}$  %,  $C_{\rm F}^{\rm O}=0.2\,{\rm wt}$  %,  $C_{\rm H_{2O}}^{\rm O}=2.5\,{\rm wt}$  %,  $K_{\rm Cl-F}^{\rm ap-m}=75$ ,  $\bar{D}_{\rm F}^{\rm c/m}=0.99$ ,  $\bar{D}_{\rm Cl}^{\rm c/m}=0.8$ , and  $\bar{D}_{\rm H_{2O}}^{\rm c/m}=0.01$ .  $H_{\rm 2O}$  speciation is calculated after Zhang (1999) for a temperature of 1020° C. The black continuous line shows modelled apatite compositional evolution during subsequent  $H_{\rm 2O}$ -saturated crystallization, with  $D_{\rm Cl}^{\rm f/m}=20$  and  $D_{\rm Cl}^{\rm f/m}=0.7$ , after  $C_{\rm H_{2O}}^{\rm m}$  reaches  $C_{\rm H_{2O}}^{\rm sat}$  at 7.5 wt %. The grey line types illustrate the difference between isobaric (0%  $H_{\rm 2O}$  loss; continuous lines)  $H_{\rm 2O}$ -saturated crystallization and polybaric  $H_{\rm 2O}$ -saturated crystallization with 0.15 wt % (dashed lines) and 0.25 wt % (dotted lines)  $H_{\rm 2O}$  loss from the melt into the fluid phase per 1% crystallization.

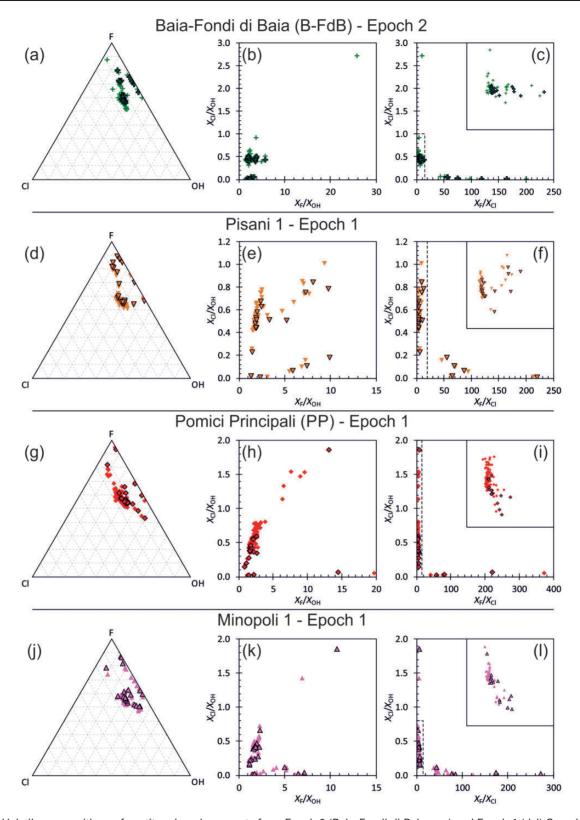


Fig. 7. Volatile compositions of apatite microphenocrysts from Epoch 2 (Baia–Fondi di Baia; a–c) and Epoch 1 (d–l) Campi Flegrei eruptions. Data are presented both in ternary space (a, d, g, j) and on binary plots of  $X_{\rm CI}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm OH}$  (b, e, h, k) and  $X_{\rm CI}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm CI}$  (c, f, i, l). Insets in (c), (f), (i) and (I) show expansions of the areas outlined by the black dashed boxes in the main panels. Points with black outlines were measured by SIMS and points without outlines were measured by EPMA. Data are distinguished by eruption (see headings).

between eruptions are more clearly Clinopyroxene- and biotite-hosted apatite inclusions from Epoch 1 eruptions (Minopoli 1, PP and Pisani 1) are typically more OH-rich than those from the more recent Astroni 1 (Epoch 3) and Monte Nuovo eruptions; this translates to lower  $X_{CI}/X_{OH}$  and  $X_F/X_{OH}$  and slightly lower  $X_{\rm F}/X_{\rm Cl}$  ratios in apatite inclusions from Epoch 1 than in those from younger eruptions (Fig. 5). The main population of apatite microphenocrysts from Astroni 1 is also typically more F-rich, with higher  $X_{CI}/X_{OH}$ ,  $X_{F}/X_{OH}$  $X_{OH}$  and  $X_F/X_{CI}$  ratios than those from Epoch 1 eruptions. Monte Nuovo microphenocrysts in this population extend to higher  $X_{CI}/X_{OH}$ ,  $X_F/X_{OH}$  and  $X_F/X_{CI}$  ratios than those from Epoch 1, but cover the entire compositional diversity of the main microphenocryst population from other eruptions (Figs 6 and 7). Baia-Fondi di Baia inclusions (Epoch 2) are relatively  $X_{CI}/X_{OH}$ depleted; B-FdB microphenocrysts are compositionally similar to those of Epoch 1 (i.e. comparable  $X_{CI}/X_{OH}$ ) but follow a different evolutionary trajectory. There is more overlap between eruptions in the lower-precision EPMA dataset but it is still clear that apatite inclusions and microphenocrysts from younger eruptions (Astroni 1 and Monte Nuovo) extend to more F-rich compositions, with higher  $X_{CI}/X_{OH}$ ,  $X_F/X_{OH}$  and  $X_F/X_{CI}$  ratios, than those from Epoch 1 (Figs 5-7).

# THERMODYNAMIC FRAMEWORK FOR APATITE VOLATILE INTERPRETATION

The composition of the apatite volatile site is described by a series of exchange equilibria between F, Cl and OH (e.g. Pan & Fleet, 2002); for example,

$$X_{\text{CI}}^{\text{Ap}} + X_{\text{OH}}^{\text{m}} \rightleftharpoons X_{\text{OH}}^{\text{Ap}} + X_{\text{CI}}^{\text{m}}$$
 (1)

where X denotes mole fraction in the apatite (Ap) and melt (m) phases. Exchange coefficients (K) for these equilibria are defined as

$$K_{\text{CI-OH}(P,T)}^{\text{Ap/m}} = \frac{a_{\text{CI}}^{\text{Ap}}}{a_{\text{CI}}^{\text{Ap}}} \cdot \frac{a_{\text{OH}}^{\text{m}}}{a_{\text{CI}}^{\text{m}}}$$
(2)

where a is activity. The evolution of F, Cl and  $\rm H_2O$  concentrations in the melt (i.e.  $a_{\rm OH}^{\rm m}$  and so on) during fractional crystallization depends on the nature of both crystal–melt partitioning and fluid–melt exchange. Therefore, we can predict how apatite compositions in equilibrium with the melt will evolve during different fractional crystallization scenarios using appropriate experimental  $K_{\rm D}$  values from the literature.

We have developed a thermodynamic model, relating apatite compositional evolution to changes in the volatile contents of silicate melts during crystallization in the presence or absence of different magmatic fluids, building on the work of Candela (1986), who treated halogens and H<sub>2</sub>O as perfectly incompatible during volatile-undersaturated crystallization. First, we

calculate the evolution of F, Cl and H<sub>2</sub>O in the melt owing to crystal ± fluid fractionation. We then compute the volatile composition of apatite in equilibrium with the melt at each stage of the crystallization model. We assume that the apatite solid solution is ideal and that the apatite-melt exchange coefficients are unaffected by the changing melt composition. The validity of these assumptions, definitions of terminology, and the values of key parameters, including mineral-melt and fluidmelt partition coefficients and the solubility of halogens and H<sub>2</sub>O in the silicate melt, are discussed in the Supplementary Data. Our approach can be applied generally to determine apatite compositions during fractional crystallization. The effects of pressure and temperature variations on apatite compositions are also outlined below.

## **Apatite fractional crystallization models**

The initial melt CI ( $C_{\text{CI}}^0$ ) and F ( $C_{\text{F}}^0$ ) concentrations are set at 0.4 wt % and 0.2 wt %, respectively, based on the mafic melt inclusion compositions given in Table 2 and Supplementary Data Table S1. The initial melt H<sub>2</sub>O concentration ( $C_{\text{H}_2\text{O}}^0$ ) is set at 2.5 wt %, similar to the predicted Campi Flegrei melt H<sub>2</sub>O concentration at apatite-in from Rhyolite-MELTS (Stock et al., 2016) and initial H<sub>2</sub>O concentrations used for modelling by Fowler et al. (2007) and Cannatelli (2012).

During volatile-undersaturated crystallization, the evolution of all volatile components in the melt is described by the Rayleigh equation:

$$C_{\rm i}^{\rm m} = C_{\rm i}^{\rm 0}.\phi^{\bar{D}_{\rm i}^{\rm c/m}-1}$$
 (3)

where  $C_i^{\rm m}$  is the concentration of i in the melt at a given melt fraction  $(\phi)$  and  $\bar{D}_i^{{\rm c/m}}$  is the bulk crystal–melt partition coefficient.

We simulate the onset of saturation with an H<sub>2</sub>Orich fluid by imposing a pre-defined H<sub>2</sub>O solubility limit  $(C_{H_2O}^{sat})$ , here taken as 7.5 wt %  $H_2O$ , which is appropriate for phonolite liquids at 200 MPa (Webster et al., 2014). Once this solubility limit is reached, the H<sub>2</sub>O concentration in the melt is held constant to simulate isobaric fractionation (or can be linearly decreased to illustrate the effect of polybaric decompression). Further H<sub>2</sub>O that exsolves owing to continuing fractionation is assigned to a separate fluid phase, following Candela & Holland (1986), and we calculate the F and CI concentration of that fluid phase using the fluid-melt partition coefficients (i.e.  $D_{\rm F}^{\rm f/m}$  and  $D_{\rm Cl}^{\rm f/m}$ , respectively). This fluid is iteratively extracted from the evolving melt by mass balance. We calculate mole fractions of CI, F and total H<sub>2</sub>O in the melt following Li & Hermann (2017) and use the H<sub>2</sub>O speciation model of Zhang (1999) for a temperature of 1020°C to calculate the mole fraction of OH in the melt. Finally, we calculate the mole fraction of F-apatite, Cl-apatite and OH-apatite in equilibrium with the calculated melt composition, following Candela (1986), and thus apatite  $X_{\rm CI}/X_{\rm OH}$ ,  $X_{\rm F}/X_{\rm OH}$  and

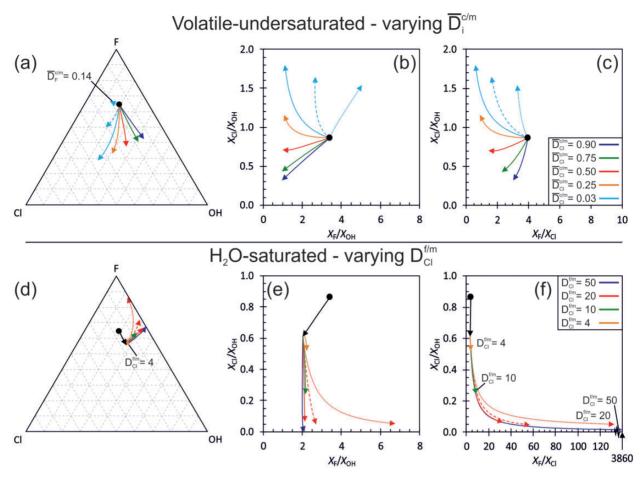


Fig. 8. Theoretical apatite compositional trajectories for different crystallization scenarios (not specific to Campi Flegrei). Data are presented both in ternary (a, d) space and on binary plots of  $X_{\rm Cl}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm OH}$  (b, e) and  $X_{\rm Cl}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm Cl}$  (c, f). In all models,  $C_{\rm Cl}^0=0.4$  wt %,  $C_{\rm H_2O}^0=2.9$  wt %,  $K_{\rm H_2O}^{\rm ep-m}=0.22$ ,  $K_{\rm Cl-Pl}^{\rm ep-m}=2.5$ ,  $K_{\rm F-OH}^{\rm ep-m}=7.5$ , and apatite crystallization begins at the black point (see discussion of input parameters in the text).  $H_2{\rm O}$  speciation is calculated after Zhang (1999) for a temperature of 1020°C. (a–c) Apatite compositional evolution during volatile-undersaturated fractional crystallization. Crystal compositions evolve in the arrow direction until 85% crystallization. Curves show the effect of varying  $\bar{D}_{\rm F}^{\rm i,m}$ , where  $\bar{D}_{\rm F}^{\rm c,m}>\bar{D}_{\rm Cl}^{\rm c,m}>\bar{D}_{\rm H_2O}^{\rm c,m}$  and  $1>\bar{D}_{\rm i}^{\rm c,m}>0$ . Line colours distinguish models with different  $\bar{D}_{\rm Cl}^{\rm c,m}$  values [see legend in (c)]. Line types distinguishes models with  $\bar{D}_{\rm F}^{\rm c,m}=0.99$  (continuous lines),  $\bar{D}_{\rm F}^{\rm c,m}=0.5$  (dashed lines) and  $\bar{D}_{\rm Cl}^{\rm c,m}=0.14$  (dotted lines).  $\bar{D}_{\rm H_2O}^{\rm c,m}$  is set at 0.003; varying  $\bar{D}_{\rm H_2O}^{\rm c,m}$  affects crystal  $X_{\rm F}/X_{\rm OH}$  and  $X_{\rm Cl}/X_{\rm OH}$  evolution in the opposing sense to  $\bar{D}_{\rm F}^{\rm c,m}$  and  $\bar{D}_{\rm Cl}^{\rm c,m}=0.14$  (dotted lines).  $\bar{D}_{\rm H_2O}^{\rm c,m}$  is set at 0.003; varying  $\bar{D}_{\rm H_2O}^{\rm c,m}$  affects crystal  $X_{\rm F}/X_{\rm OH}$  and  $X_{\rm Cl}/X_{\rm OH}$  and  $X_{\rm Cl}$ 

# Model outputs: apatite compositional evolution during fractional crystallization

During volatile-undersaturated crystallization for a range of parameters (see Supplementary Data), apatite compositions typically show a decrease in both  $X_F/X_{OH}$  and  $X_F/X_{CI}$  ratios, whereas  $X_{CI}/X_{OH}$  may increase or decrease depending on the exact values of the crystalmelt partition coefficients for F, CI and OH (Fig. 8b and c). In contrast,  $H_2O$ -saturated crystallization is

characterized by a strong decrease in apatite  $X_{\rm CI}/X_{\rm OH}$  ratios per unit crystallization, with increasing  $X_{\rm F}/X_{\rm CI}$  and variable  $X_{\rm F}/X_{\rm OH}$  (Fig. 8e and f). This is because fluid—melt partition coefficients for CI are typically much higher than for F (Borodulin *et al.*, 2009; Webster *et al.*, 2014). In ternary space, volatile-undersaturated trajectories are typified by apatite compositional evolution away from the F apex (Fig. 8a). Conversely, during H<sub>2</sub>O-saturated crystallization apatite compositions move

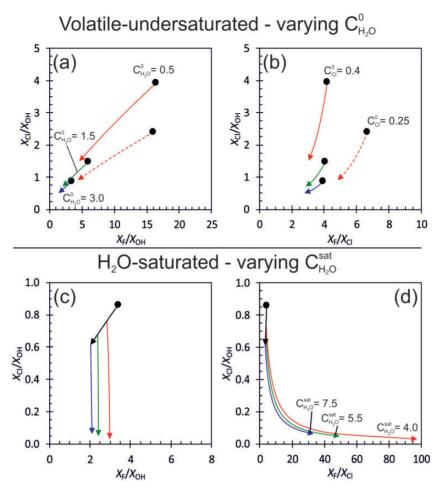


Fig. 9. The effect of varying input parameters in general (i.e. not specific to Campi Flegrei) thermodynamic models of apatite compositional evolution in  $X_{\rm Cl}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm OH}$  (a, c) and  $X_{\rm Cl}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm Cl}$  (b, d) binary space.  $C_{\rm F}^0$ ,  $\bar{D}_{\rm F}^{\rm c/m}$ ,  $\bar{D}_{\rm H_2O}^{\rm c/m}$ ,  $D_{\rm F}^{\rm f/m}$ ,  $K_{\rm Cl-P}^{\rm ap-m}$ ,  $K_{\rm Cl-PH}^{\rm ap-m}$ ,  $K_{\rm F-OH}^{\rm ap-m}$  and H<sub>2</sub>O speciation are as in Fig. 8d–f. Apatite crystallization begins at the black point and continues to 85% crystallization in all models. (a, b) Trajectories of volatile-undersaturated apatite compositional evolution with varying  $C_{\rm i}^0$ . Line colours distinguish models with  $C_{\rm H_2O}^0 = 3.0\,{\rm wt}$  % (purple),  $C_{\rm H_2O}^0 = 1.5\,{\rm wt}$  % (green) and  $C_{\rm H_2O}^0 = 0.5\,{\rm wt}$  % (red). Line types distinguish models with  $C_{\rm Cl}^0 = 0.4\,{\rm wt}$  % (continuous lines) and  $C_{\rm Cl}^0 = 0.25\,{\rm wt}$  % (dashed lines). (c, d) Trajectories of apatite compositional evolution with varying  $C_{\rm H_2O}^0$  are as in Fig. 8 and  $D_{\rm Cl}^{\rm f/m} = 20$ . The black arrows show an arbitrary trajectory of 61% volatile-undersaturated crystallization, as in Fig. 8e and f. Line colours show isobaric (0% H<sub>2</sub>O loss) H<sub>2</sub>O-saturated crystallization after 61%, 47% and 28% crystallization, where  $C_{\rm H_2O}^{\rm sat}$  is 7.5 wt % (purple), 5.5 wt % (green) and 4 wt % (red), respectively.

towards the F–OH binary (Fig. 8d). Under polybaric  $H_2O$ -saturated conditions, where  $H_2O$  is lost from the melt during depressurization, crystal compositions may then curve towards the F apex (Fig. 8d).

### The melt volatile composition

Varying the initial melt volatile concentrations ( $C_{\rm i}^0$ ) translates the starting point of the model in apatite compositional space, and changes the relative magnitudes of the calculated variations in apatite volatile components per unit crystallization. However, the slope or direction in binary or ternary space remains the same (Fig. 9a and b). In the general H<sub>2</sub>O-saturated case, decreasing  $C_{\rm H_2O}^{\rm sat}$  (e.g. at lower pressure or a different melt composition) reduces the extent of volatile-undersaturated crystallization before second boiling, but does not affect the trajectory of apatite compositional evolution after volatile-

saturation is achieved (Fig. 9c and d). Our model does not deal with saturation of either fluorite, which is present only as a late-stage accessory mineral, or brine. Fluorine and Cl act as incompatible elements during fractionation (i.e. higher halogen concentrations in low-MgO in glass analyses; Fig. 4a–d), which suggests that brine exsolution is also not a major factor at Campi Flegrei. We would anticipate that formation of fluorite and/or brine would change the mineral–melt and/or fluid–melt volatile partition coefficients at the late stages of fractional crystallization.

# Effect of pressure and temperature on apatite composition

As apatite—melt halogen—OH exchange coefficients depend on pressure and temperature (Piccoli & Candela, 1994; Riker *et al.*, 2018), mineral compositions may vary, even at a constant melt composition [e.g. equation (2)].

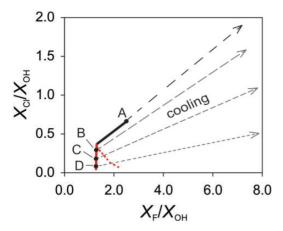


Fig. 10. Effect of re-equilibration on apatite compositions during cooling, shown in binary  $X_{\rm CI}/X_{\rm OH}$  vs  $X_{\rm F}/X_{\rm OH}$  space. The continuous lines represent an arbitrary trend of apatite compositional evolution during volatile-undersaturated (as in Fig. 8b; black line) followed by isobaric H<sub>2</sub>O-saturated (as in Fig. 8e; red line) fractional crystallization. The dotted red line illustrates polybaric H<sub>2</sub>O-saturated fractional crystallization (as in Fig. 8e). The dashed arrows show modelled trajectories of re-equilibration during cooling from 1000°C to 750°C, starting from different points on the volatile-undersaturated (point A) and isobaric H<sub>2</sub>O-saturated (points B-D) trend. In each cooling model, the melt composition is fixed and  $K_D$  values are calculated from the P-independent regressions of Riker et al. (2018). Cooling of a volatile-undersaturated system, where CI is retained within the melt (i.e. starting at point A), will cause apatite re-equilibration towards elevated  $X_{\rm C}/X_{\rm OH}$  and  $X_{\rm F}/X_{\rm OH}$  ratios (black dashed arrow). Extraction of CI from the melt into the fluid following H<sub>2</sub>O-saturation (i.e. starting sequentially from point B to point D) increasingly prevents elevation of  $X_{\rm CI}/X_{\rm OH}$ and drives apatite compositions to high  $X_F/X_{OH}$  ratios only (dashed grey lines). Re-equilibration of apatites that grew under volatile-undersaturated conditions to H<sub>2</sub>O-saturated conditions could potentially result in intermediate compositions in the space between A-B and B-D.

This is not treated in the fractional crystallization model detailed above but we illustrate the effects of cooling in Fig. 10, using the experimental regression of Riker et al. (2018); this is qualitatively consistent with the earlier work of Piccoli & Candela (1994). Cooling drives apatite compositions to higher  $X_{CI}/X_{OH}$ ,  $X_{F}/X_{OH}$  and  $X_{F}/X_{CI}$  ratios, reflecting an increasing preference for the smaller F anion at low temperatures. However, the slope (in binary space) of the cooling trend depends strongly on the CI concentration of the liquid; at low melt CI contents, apatite is driven towards high  $X_F/X_{OH}$  only (Fig. 10). Pressure has little effect on apatite compositions at constant melt composition (Piccoli & Candela, 1994; Riker et al., 2018), with depressurization primarily allowing a slight increase in the apatite CI component, which is reflected in a minor decrease in  $X_F/X_{CI}$  and increase in  $X_{CI}/X_{OH}$  ratios.

## **DISCUSSION**

# Application of apatite thermodynamic model to Campi Flegrei

The clear differences in model apatite trajectories between  $H_2O$ -saturated and volatile-undersaturated

conditions (Fig. 8) suggest that apatite may hold significant clues to understanding the late-stage variations in fluid systematics in magmatic—hydrothermal systems.

Texturally constrained analyses give a temporal dimension to the apatite dataset (Figs 5-7) and analysis of Astroni 1 has shown that the  $X_{\text{Halogen}}/X_{\text{OH}}$  ratios of crystals on the main compositional trend decrease during magmatic evolution (Stock et al., 2016). The same temporal trend is apparent in all eruptions analysed: apatite inclusions trapped in early formed clinopyroxene phenocrysts (clinopyroxene-in = 1080°C; Stock et al., 2016) extend to higher  $X_{\text{Halogen}}/X_{\text{OH}}$  ratios than in late-formed biotite phenocrysts (biotite-in = 910°C; Stock et al., 2016). Figure 6a-c shows that a good model fit can be achieved with our natural apatite analyses from Campi Flegrei using  $K_{\text{Cl-OH}}^{\text{Ap/m}} = 25$ ,  $K_{\text{F-OH}}^{\text{Ap/m}} = 75$  and  $K_{\text{Cl-F}}^{\text{Ap/m}} = 0.22$ , which are within the range identified experimentally (Riker et al., 2018) and the effects of varying input parameters on the model fit are shown in Figs 8 and 9, and Supplementary Data Figs S3-S6. The model fit (Fig. 6a-c) represents ~70% crystallization at volatileundersaturated conditions, with constant crystal-melt partition coefficients  $\bar{D}_{\rm F}^{{\rm c/m}} \approx 0.99$ ,  $\bar{D}_{\rm Cl}^{{\rm c/m}} \approx 0.8$  and  $ar{D}_{
m OH}^{
m c/m} pprox$  0.01 (Fig. 6a–c). These are realistic given the increase in glass halogen concentrations with decreasing MgO and the low modal abundance of apatite and biotite. The low  $X_{\rm CI}/X_{\rm OH}$  values identified in some apatite microphenocrysts can be partially reproduced by subsequent  $H_2O$ -saturated crystallization, with  $\bar{D}_{Cl}^{f/m}=20$  and  $\bar{D}_{F}^{f/m}=0.7$  (Fig. 6a–c), consistent with alkali melt– fluid halogen partitioning experiments under magmatic conditions (Signorelli & Carroll, 2000; Borodulin et al., 2009). However, our simple isobaric, isothermal H<sub>2</sub>Osaturated crystallization models cannot reproduce the range of  $X_{\rm F}/X_{\rm OH}$  ratios observed in the natural low  $X_{\rm CI}/$ X<sub>OH</sub> subsidiary population (Fig. 8d-f). Decompressioninduced degassing could account for some of this compositional variability but cooling, after the apatites have re-equilibrated under H2O-saturated conditions, appears to be the primary mechanism to drive compositions to the very high  $X_{\rm F}/X_{\rm OH}$  ratios observed in some eruptions (e.g. PP; Figs 7 and 10). The second, high  $X_{CI}$  $X_{OH}$  and  $X_F/X_{OH}$  subsidiary population (e.g. Minopoli 1) is consistent with re-equilibration of apatites during cooling of a less degassed melt (Fig. 10).

# Pre-eruptive volatile behaviour and the structure of the Campi Flegrei magmatic system

Our models reproduce the main compositional trend in apatite inclusions and microphenocrysts through protracted volatile-undersaturated crystallization (Figs 5–7). This suggests that the main zone of magma storage and crystallization remained volatile-undersaturated until late in magmatic evolution (i.e. after the entrapment of biotite-hosted inclusions; see Stock *et al.*, 2016) prior to past eruptions at Campi Flegrei. Thermodynamic (Rhyolite-MELTS) fractional crystallization models, combined with relevant phonolite

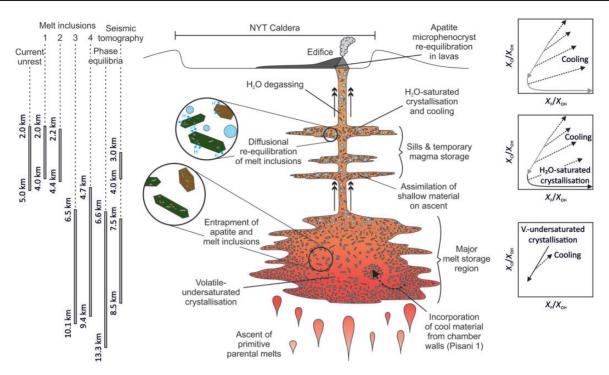


Fig. 11. Schematic diagram summarizing volatile systematics in the Campi Flegrei sub-volcanic plumbing system, based on apatite and glass compositions. The representative graphs (right) illustrate how apatite compositional trends relate to different parts of the magmatic system (black lines show apatite trends that form at a particular depth; grey lines show trends that form at a greater depth but are brought up to shallower levels by ascending magmas). The grey bars (left) display relative estimates of Campi Flegrei magma storage depths, derived from independent constraints: 'current unrest' estimates are from recent ground deformation (Woo & Kilburn, 2010; Amoruso et al., 2014b; D'Auria et al., 2015), fumarolic gas emissions (Carlino et al., 2015) and heat-flow measurements (Di Maio et al., 2015); melt inclusion estimates are for <15 ka eruptions and are from 1—Fourmentraux et al. (2012), 2—Vetere et al. (2011), 3—Arienzo et al. (2016) and 4—Arienzo et al. (2010); phase equilibria constraints are from Bohrson et al. (2006), Fowler et al. (2007) and Cannatelli (2012); seismic tomography data are from Zollo et al. (2008) and De Siena et al. (2010). Pressure estimates were converted to depths using a consistent crustal density of 2·3 kg cm<sup>-3</sup> after Rosi & Sbrana (1987). Not to scale.

solubility data, also show that H<sub>2</sub>O would remain undersaturated in Campi Flegrei melts until after biotite-in (i.e. ~910°C) at pressures ≥150 MPa (Carroll & Blank, 1997; Webster et al., 2014; Stock et al., 2016). Our conclusion that the system remained persistently volatileundersaturated until a late stage is therefore consistent with independent estimates of Campi Flegrei magma storage conditions, which suggest a high-pressure zone of major melt storage and crystallization (Fig. 11). Seismic tomography shows that the current region of major melt storage beneath Campi Flegrei is ~1 km thick and at a depth of 7.5 km (Zollo et al., 2008), which would equate to storage pressures of  $\sim$ 170–190 MPa. This broadly agrees with phase equilibria models, which indicate that mineral and glass compositions produced in past Campi Flegrei eruptions were stable at pressures between 150 and 300 MPa (i.e. 6-6-13-3 km; Bohrson et al., 2006; Fowler et al., 2007; Cannatelli, 2012), and the volatile contents of melt inclusions from <15 ka Campi Flegrei eruptions (Fig. 11), which give saturation pressures up to 229 MPa [~10 km; Arienzo et al., 2016; using a crustal density of 2.3 g cm<sup>-3</sup> after Rosi & Sbrana (1987)]. In particular, magma storage depths inferred from melt inclusion volatile contents for the Agnano-Monte Spina (Arienzo et al., 2010) and Nisida

(Arienzo *et al.*, 2016) eruptions overlap at 6·5–9·4 km, which agrees very well with the current depth of major magma storage (7·5–8·5 km; Zollo *et al.*, 2008). This appears to demonstrate that melt storage and crystallization before past eruptions occurred at very similar depths to today.

In addition to the main population of apatite inclusions and microphenocrysts, two subsidiary sets of lower-Cl microphenocrysts are present in some eruptions, as discussed above. Apatites with low  $X_{\rm Cl}/X_{\rm OH}$ ratios and high  $X_F/X_{OH}$  and  $X_F/X_{CI}$  ratios (i.e. in Monte Nuovo, Astroni 1, B-FdB, Pisani 1, PP, Minopoli 1; Figs 6 and 7) are consistent with H<sub>2</sub>O-saturated fractionation at high  $\mathcal{D}_{\mathrm{Cl}}^{\mathrm{f/m}},$  where CI is efficiently extracted from the melt into the fluid phase, but with cooling following extensive fluid exsolution. This requires equilibrating first with low CI concentrations in the melt owing to degassing, so that any later thermal re-equilibration drives apatites to high  $X_F/X_{OH}$  and  $X_F/X_{CI}$  only (Fig. 10). The second subsidiary set of apatites, with elevated  $X_{\rm CI}/X_{\rm OH}$ and/or  $X_{E}/X_{OH}$  ratios (i.e. in Astroni 1, B-FdB, PP, Minopoli 1; Figs 6 and 7), are consistent with equilibration during cooling of a melt that is not depleted in Cl (i.e. where CI has not been extracted into an exsolving fluid; Fig. 10).

These subsidiary compositions represent a small fraction of our dataset, but the analyses are from the cores of microphenocrysts, suggesting that they may have been fully equilibrated under different conditions. In contrast, microphenocrysts that plot on the main trend show no evidence of re-equilibration or overgrowth of new compositions, within the scale of our analyses [Stock et al. (2016) checked several grains for compositional zonation within >16 µm of the rim; although we anticipate that fine-scale zonation might be observable using higher-resolution methods]. The rapid experimental diffusivities of halogens in apatite (Brenan, 1993) suggest that these few anomalous grains must therefore have been spatially separated within the magmatic system. The simplest interpretation of our data is that the low-X<sub>CI</sub>/X<sub>OH</sub> apatites represent mushy material left over after incomplete evacuation of magma from separate, shallower storage regions (Fig. 11). Small, melt-rich magma bodies at ~2-4km beneath the current Campi Flegrei have been imaged by seismic tomography (De Siena et al., 2010) and inferred from modelling of long-term ground deformation (Amoruso et al., 2014a), fumarolic gas emissions (Carlino et al., 2015) and heat-flow measurements (Di Maio et al., 2015). These shallow magma bodies could derive from emplacement of sills between eruptions; recent seismic crises (1982-1984 and 2011-2013) have been attributed to sill intrusion at 2.75-5 km depth (Fig. 11; Woo & Kilburn, 2010; Amoruso et al., 2014b; D'Auria et al., 2015). Such small, shallow magma bodies would be saturated in volatiles and cooled by interaction with the country rock, allowing microphenocrysts to grow or re-equilibrate. Apatite inclusions, trapped before the magmas were emplaced in the shallow crust, would be 'armoured' by their host phenocrysts, preventing re-equilibration and allowing them to volatile-undersaturated signature. Inclusions that were incompletely enclosed by their host phenocrysts do show evidence of diffusional reequilibration, manifest as depleted CI contents relative to fully enclosed inclusions (Stock, 2016). We suggest that magmas ascending from the deep storage region may interact with these shallow melt bodies during eruptions, assimilating some of the residual material and acquiring a mixed cargo of microphenocrysts (Fig. 11). Magmas derived from the deep storage region degassed H<sub>2</sub>O during ascent, resulting in low matrix glass H<sub>2</sub>O contents (Fig. 4f). Similar F and Cl contents in matrix glasses and melt inclusions (Fig. 4a-d) demonstrate that halogens were not significantly degassed, despite  $D_{Cl}^{f/m} > 1$  (Signorelli & Carroll, 2000). Relatively undepleted matrix glass CI concentrations are consistent with kinetically induced vapour-melt disequilibrium during rapid ascent (Webster et al., 1993; Barclay et al., 1996; Shea et al., 2014).

The apatite microphenocrysts from Pisani 1 that branch away from the main population at moderate  $X_{\rm Cl}/X_{\rm OH}$  are consistent with cooling (Fig. 10), but before any H<sub>2</sub>O-saturated crystallization has taken place. We

suggest that these crystals formed in the main magma storage region at depth, possibly in cooler regions of the reservoir (e.g. close to the walls) and were incorporated before eruption (Fig. 11).

The majority of apatite inclusions and microphenocrysts from B-FdB are compositionally distinct, with significantly lower CI contents than other explosive eruptions (Figs 5 and 7a-c). This could be explained if part of the B-FdB magma crystallized from a melt with a different initial volatile content (i.e. lower Cl. higher H<sub>2</sub>O; Fig. 9a and b). The B-FdB eruption occurred at the start of Epoch 2 after a long hiatus in activity (Smith et al., 2011), and may have assimilated some of the country rock surrounding its magma reservoir during this time. Large volumes of residue from the Campanian Ignimbrite magma chamber underlie the Campi Flegrei caldera (D'Antonio, 2011) and represent a potential assimilant. Melt inclusions from Campanian Ignimbrite extend to high H<sub>2</sub>O contents and Campanian Ignimbrite glasses are consistently Cldeficient compared with eruptions in the past 15 kyr (see Fig. 4a and b; Signorelli et al., 1999; Marianelli et al., 2006). Furthermore, B-FdB glasses are depleted in Sr and Ba and enriched in Zr relative to other < 15 ka eruptions (Smith et al., 2011), again suggesting possible contamination from older (>15 ka) Campi Flegrei eruptions, including the Campanian Ignimbrite (Tomlinson et al., 2012).

Apatite microphenocrysts from Accademia and SMdG lava samples are distinct from those in explosive deposits. They do not show the same dominant compositional trend of F-OH exchange with an approximately constant CI, but rather define a trend with an increasing F component at the expense of CI, curving first towards the F-CI binary and then towards the F apex (Fig. 6g-I). This is analogous to the modelled trend of apatite compositional evolution during cooling. However, in binary space, apatite microphenocrysts from lava deposits are scattered with no discernible trend. We interpret this as variable extents of diffusive re-equilibration in the lavas (Fig. 10). Although microphenocrysts plot on the same cooling trend, they have different starting points and have variably rehomogenized. Such extensive reequilibration of microphenocryst compositions is unique to lava samples and we interpret that it occurred during slow cooling of the host magma at or near the surface (Fig. 11).

Apatite inclusions in clinopyroxene and biotite record volatile-undersaturated crystallization (see above) and we would therefore expect the  $H_2O$  contents of melt inclusions in the same phenocryst phases to correlate negatively with MgO, as  $H_2O$  behaves incompatibly during magmatic evolution. However, there is no correlation between clinopyroxene-hosted melt inclusion  $H_2O$  and MgO concentrations in any of the analysed eruptions, and most melt inclusions measured in this study contain  $\sim\!\!2\cdot0\!-\!3\cdot8$  wt %  $H_2O$  (Fig. 4e). This may indicate that melt inclusion  $H_2O$  contents have been reset after entrapment, reflecting the final ( $H_2O$ -saturated)

equilibration pressure at shallow crustal levels. Rapid H<sup>+</sup> diffusion through common phenocrysts causes reequilibration of melt inclusion H<sub>2</sub>O concentrations on short timescales (potentially hours to weeks under magmatic conditions) in response to any change in external conditions (e.g. Woods *et al.*, 2000; Ingrin & Blanchard, 2006; Reubi *et al.*, 2013). In contrast, reequilibration of apatite inclusions would be inhibited by the need for coupled halogen diffusion to retain stoichiometry (Bucholz *et al.*, 2013). Occasional volatile-rich melt inclusions in <15 ka Campi Flegrei deposits attest to evolved melts at depth (i.e. >5 wt % H<sub>2</sub>O in Fig. 4e).

Campi Flegrei melt inclusions consistently have very low CO2 contents (e.g. Marianelli et al., 2006; Arienzo et al., 2010; Stock et al., 2016), so we used a polynomial regression through published 850-950°C, CO2-free phonolite and trachyte solubility data (Carroll & Blank, 1997: Di Matteo et al., 2004: Larsen & Gardner, 2004: Schmidt & Behrens, 2008; Webster et al., 2014) to estimate a saturation pressure of  $\sim\!\!24\text{--}76\,\text{MPa}$  for the degassed melt inclusions (2.0-3.8 wt % H<sub>2</sub>O). This is equivalent to storage depths of ~1.1-3.4 km, which coincide with the depth of small, melt-rich magma bodies imaged beneath the current Campi Flegrei caldera (~2-4 km; De Siena et al., 2010). We infer that ascending magmas interacted with these shallow melt bodies during eruption, stalling or slowing such that melt inclusions from the deep reservoir re-equilibrated (Fig. 11), and incorporating phenocrysts with degassed melt inclusions (along with apatite microphenocrysts), which were stored in this shallow crystal residue.

# Inter-eruption variability in pre-eruptive magma volatile contents

Our data generally show the same qualitative trends in apatite inclusion and microphenocryst compositions in explosive eruptions, regardless of the erupted volume, melt composition or epoch of activity. The majority of apatite inclusions and microphenocrysts plot on a trend that is consistent with volatile-undersaturated fractionation. However, in different eruptions, this main apatite trend is offset depending on the epoch of activity, suggesting a temporal variation in melt volatile composition. This is particularly evident in SIMS data, which are more precise than the EPMA data. For example, apatite inclusions from Astroni 1 (Epoch 3) and the historical Monte Nuovo eruptions are typically F-rich and OH-depleted with higher  $X_{\text{Halogen}}/X_{\text{OH}}$  ratios than those in Epoch 1 (Minopoli 1, PP, Pisani 1; Fig. 5). The same is true for microphenocrysts, except that microphenocrysts from Monte Nuovo cover the entire compositional range measured in other eruptions (Figs 6 and 7). Epoch 2 (B-FdB) crystals have similar  $X_{CI}/X_{OH}$  ratios to those from Epoch 1 but cannot be compared directly given their distinct compositional trajectory (Fig. 7b and c).

Melt inclusion halogen contents are typically lower in Epoch 1 than in later explosive eruptions (Fig. 4a and

c), consistent with the more primitive Epoch 1 major element glass compositions (Figs 2 and 3). As  $\bar{D}_{\rm F}^{\rm c/m}>\bar{D}_{\rm Cl}^{\rm c/m}>\bar{D}_{\rm H_2O}^{\rm c/m}$  during apatite crystallization (see Supplementary Data), we would anticipate melt H2O contents to be proportionally lower for Epoch 1 magmas. Consequently, we would expect apatites forming from the more primitive Epoch 1 magmas to have higher  $X_{\text{Halogen}}/X_{\text{OH}}$  than those forming from Epoch 3 magmas, the opposite to what is observed (Fig. 5). This disparity can be explained by a systematic difference in the melt volatile composition at the point of initial apatite crystallization (i.e. apatite-in;  $C_i^0$ ), changing apatite compositions without significantly affecting mineral phase stabilities or the trajectory of apatite compositional evolution. Specifically, the change to higher X<sub>Halogen</sub>/X<sub>OH</sub> in Astroni 1 (Epoch 3) and Monte Nuovo suggests a decrease in melt H<sub>2</sub>O concentrations at apatite-in  $(C_{H_2O}^0)$  relative to Epoch 1 (Fig. 9a and b). At Campi Flegrei, Rhyolite-MELTS models demonstrate that  $C_{H_2O}^0$  depends on the  $H_2O$  content of the melt at the liquidus, and does not change significantly as a result of variations in pressure or oxygen fugacity (Stock et al., 2016). Variability in Campi Flegrei apatite compositions from different epochs is therefore likely to reflect temporal variations in the H<sub>2</sub>O concentration of parental melts feeding the system; Epoch 1 parental melts have higher apparent H<sub>2</sub>O contents than more recent eruptions. This agrees with isotopic evidence, which suggests that Epochs 1 and 2 were fed by the 'Minopoli 2' and 'NYT' end-members (i.e. a high slab-derived fluid and/or low subducted sediment mantle wedge input) primitive melts, whereas Epoch 3 and Monte Nuovo melts include the 'Astroni 6 end-member' component (i.e. a reduced slab-derived fluid and/or low subducted sediment mantle wedge input; D'Antonio et al., 2007; Di Renzo et al., 2011).

## **CONCLUSIONS**

Apatite analysis represents a robust approach for assessing magmatic volatile behaviour that has many major advantages over established techniques (e.g. melt inclusion analysis). Apatite microphenocrysts remain in equilibrium with their host melts on geologically short timescales but are unable to re-equilibrate on very short timescales during magma ascent. In contrast, apatite inclusions are isolated from the magma on entrapment and preserve a record of conditions earlier in crystallization. In explosive Campi Flegrei eruptions, inclusions record persistent undersaturated conditions until late in the crystallization history of the evolving magma, interpreted to occur in a deep magma reservoir at ~7.5-8.5 km, which corresponds with the current depth of major magma storage at Campi Flegrei. In addition to this main volatileundersaturated trend, apatite microphenocrysts also display subsidiary trends, related to cooling and H2Osaturated crystallization. The compositional diversity of apatite microphenocrysts is consistent with a mixed crystal cargo, attained as melts ascending from depth interact with crystal residue in discrete shallow-crustal magma bodies. Clinopyroxene-hosted melt inclusion H<sub>2</sub>O contents from Campi Flegrei commonly record saturation pressures equating to 1-3 km depth, owing to incorporation of phenocrysts that have been stored in the shallow crust and/or stalling of ascending magmas as they interact with shallow crustal magma bodies. This corresponds with small melt zones imaged seismically beneath the current Campi Flegrei caldera and the depth of sill emplacement during recent seismic crises. In lava deposits, apatite microphenocrysts diffusively re-equilibrate during cooling at the surface. A systematic offset in volatile-undersaturated apatite compositions between eruptions suggests that the H<sub>2</sub>O content of parental melts feeding the system was higher in Epoch 1 than in more recent times (i.e. Epoch 3 and Monte Nuovo). Although we have focused on Campi Flegrei, our approach of comparing thermodynamic models of apatite compositional evolution with empirical analyses can be applied widely to determine the temporal evolution of pre-eruptive volatile behaviour in other volcanic systems, providing new perspectives on sub-volcanic processes.

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## **SUPPLEMENTARY DATA**

Supplementary data for this paper are available at Journal of Petrology online.

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